

Conformal transformations
and
the beginning of the Universe.
Part III.

Pawel Nurowski

Centrum Fizyki Teoretycznej
Polska Akademia Nauk

GRIEG running seminar nr. 4, 12.01.2021

- A conformal transformation: $g \rightarrow \hat{g} = \frac{1}{\Omega^2} g$ results in the following transformation of the respective Ricci tensors:

$$\hat{R}_{\mu\nu} = \Omega^2 R_{\mu\nu} + 2\Omega \nabla_\nu \Omega_{,\mu} + g_{\mu\nu} (\Omega \square \Omega - 3g(\nabla \Omega, \nabla \Omega))$$

- Contracting we get the following transformation of the Ricci scalars:

$$\hat{R} = \Omega^2 R + 6\Omega \square \Omega - 12g(\nabla \Omega, \nabla \Omega).$$

- Interpreting \hat{g} as **the physical metric of spacetime**, whose conformal compactified metric g has \mathcal{I} where $\Omega \rightarrow 0$, we see that the **causal properties of \mathcal{I}** are governed by the formula:

$$\hat{R} = -12g(\nabla \Omega, \nabla \Omega).$$

- Recall that the signature is $(-, +, +, +)$, so \mathcal{I} **is spacelike** if $\nabla \Omega$ is timelike, i.e. **if the Ricci scalar \hat{R} of the physical metric is positive, $\hat{R} > 0$.**

- A conformal transformation: $g \rightarrow \hat{g} = \frac{1}{\Omega^2}g$ results in the following transformation of the respective Ricci tensors:

$$\hat{R}_{\mu\nu} = \Omega^2 R_{\mu\nu} + 2\Omega \nabla_\nu \Omega_{,\mu} + g_{\mu\nu} (\Omega \square \Omega - 3g(\nabla \Omega, \nabla \Omega))$$

- Contracting we get the following transformation of the Ricci scalars:

$$\hat{R} = \Omega^2 R + 6\Omega \square \Omega - 12g(\nabla \Omega, \nabla \Omega).$$

- Interpreting \hat{g} as **the physical metric of spacetime**, whose conformal compactified metric g has \mathcal{I} where $\Omega \rightarrow 0$, we see that the **causal properties of \mathcal{I}** are governed by the formula:

$$\hat{R} = -12g(\nabla \Omega, \nabla \Omega).$$

- Recall that the signature is $(-, +, +, +)$, so \mathcal{I} is **spacelike** if $\nabla \Omega$ is timelike, i.e. **if the Ricci scalar \hat{R} of the physical metric is positive, $\hat{R} > 0$.**

- A conformal transformation: $g \rightarrow \hat{g} = \frac{1}{\Omega^2}g$ results in the following transformation of the respective Ricci tensors:

$$\hat{R}_{\mu\nu} = \Omega^2 R_{\mu\nu} + 2\Omega \nabla_\nu \Omega_{,\mu} + g_{\mu\nu} (\Omega \square \Omega - 3g(\nabla \Omega, \nabla \Omega))$$

- Contracting we get the following transformation of the Ricci scalars:

$$\hat{R} = \Omega^2 R + 6\Omega \square \Omega - 12g(\nabla \Omega, \nabla \Omega).$$

- Interpreting \hat{g} as **the physical metric of spacetime**, whose conformal compactified metric g has \mathcal{I} where $\Omega \rightarrow 0$, we see that the **causal properties of \mathcal{I}** are governed by the formula:

$$\hat{R} = -12g(\nabla \Omega, \nabla \Omega).$$

- Recall that the signature is $(-, +, +, +)$, so \mathcal{I} **is spacelike** if $\nabla \Omega$ is timelike, i.e. **if the Ricci scalar \hat{R} of the physical metric is positive, $\hat{R} > 0$.**

- A conformal transformation: $g \rightarrow \hat{g} = \frac{1}{\Omega^2}g$ results in the following transformation of the respective Ricci tensors:

$$\hat{R}_{\mu\nu} = \Omega^2 R_{\mu\nu} + 2\Omega \nabla_\nu \Omega_{,\mu} + g_{\mu\nu} (\Omega \square \Omega - 3g(\nabla \Omega, \nabla \Omega))$$

- Contracting we get the following transformation of the Ricci scalars:

$$\hat{R} = \Omega^2 R + 6\Omega \square \Omega - 12g(\nabla \Omega, \nabla \Omega).$$

- Interpreting \hat{g} as **the physical metric of spacetime**, whose conformal compactified metric g has \mathcal{I} where $\Omega \rightarrow 0$, we see that the **causal properties of \mathcal{I}** are governed by the formula:

$$\hat{R} = -12g(\nabla \Omega, \nabla \Omega).$$

- Recall that the signature is $(-, +, +, +)$, so \mathcal{I} is **spacelike** if $\nabla \Omega$ is timelike, i.e. **if the Ricci scalar \hat{R} of the physical metric is positive, $\hat{R} > 0$.**

- A conformal transformation: $g \rightarrow \hat{g} = \frac{1}{\Omega^2}g$ results in the following transformation of the respective Ricci tensors:

$$\hat{R}_{\mu\nu} = \Omega^2 R_{\mu\nu} + 2\Omega \nabla_\nu \Omega_{,\mu} + g_{\mu\nu} (\Omega \square \Omega - 3g(\nabla \Omega, \nabla \Omega))$$

- Contracting we get the following transformation of the Ricci scalars:

$$\hat{R} = \Omega^2 R + 6\Omega \square \Omega - 12g(\nabla \Omega, \nabla \Omega).$$

- Interpreting \hat{g} as **the physical metric of spacetime**, whose conformal compactified metric g has \mathcal{I} where $\Omega \rightarrow 0$, we see that the **causal properties of \mathcal{I}** are governed by the formula:

$$\hat{R} = -12g(\nabla \Omega, \nabla \Omega).$$

- Recall that the signature is $(-, +, +, +)$, so \mathcal{I} is **spacelike** if $\nabla \Omega$ is timelike, i.e. **if the Ricci scalar \hat{R} of the physical metric is positive, $\hat{R} > 0$.**

- A conformal transformation: $g \rightarrow \hat{g} = \frac{1}{\Omega^2}g$ results in the following transformation of the respective Ricci tensors:

$$\hat{R}_{\mu\nu} = \Omega^2 R_{\mu\nu} + 2\Omega \nabla_\nu \Omega_{,\mu} + g_{\mu\nu} (\Omega \square \Omega - 3g(\nabla \Omega, \nabla \Omega))$$

- Contracting we get the following transformation of the Ricci scalars:

$$\hat{R} = \Omega^2 R + 6\Omega \square \Omega - 12g(\nabla \Omega, \nabla \Omega).$$

- Interpreting \hat{g} as **the physical metric of spacetime**, whose conformal compactified metric g has \mathcal{I} where $\Omega \rightarrow 0$, we see that the **causal properties of \mathcal{I}** are governed by the formula:

$$\hat{R} = -12g(\nabla \Omega, \nabla \Omega).$$

- Recall that the signature is $(-, +, +, +)$, so \mathcal{I} is **spacelike** if $\nabla \Omega$ is timelike, i.e. **if the Ricci scalar \hat{R} of the physical metric is positive, $\hat{R} > 0$.**

- A conformal transformation: $g \rightarrow \hat{g} = \frac{1}{\Omega^2}g$ results in the following transformation of the respective Ricci tensors:

$$\hat{R}_{\mu\nu} = \Omega^2 R_{\mu\nu} + 2\Omega \nabla_\nu \Omega_{,\mu} + g_{\mu\nu} (\Omega \square \Omega - 3g(\nabla\Omega, \nabla\Omega))$$

- Contracting we get the following transformation of the Ricci scalars:

$$\hat{R} = \Omega^2 R + 6\Omega \square \Omega - 12g(\nabla\Omega, \nabla\Omega).$$

- Interpreting \hat{g} as **the physical metric of spacetime**, whose conformal compactified metric g has \mathcal{I} where $\Omega \rightarrow 0$, we see that the **causal properties of \mathcal{I}** are governed by the formula:

$$\hat{R} = -12g(\nabla\Omega, \nabla\Omega).$$

- Recall that the signature is $(-, +, +, +)$, so \mathcal{I} is **spacelike** if $\nabla\Omega$ is timelike, i.e. **if the Ricci scalar \hat{R} of the physical metric is positive, $\hat{R} > 0$.**

- A conformal transformation: $g \rightarrow \hat{g} = \frac{1}{\Omega^2}g$ results in the following transformation of the respective Ricci tensors:

$$\hat{R}_{\mu\nu} = \Omega^2 R_{\mu\nu} + 2\Omega \nabla_\nu \Omega_{,\mu} + g_{\mu\nu} (\Omega \square \Omega - 3g(\nabla\Omega, \nabla\Omega))$$

- Contracting we get the following transformation of the Ricci scalars:

$$\hat{R} = \Omega^2 R + 6\Omega \square \Omega - 12g(\nabla\Omega, \nabla\Omega).$$

- Interpreting \hat{g} as **the physical metric of spacetime**, whose conformal compactified metric g has \mathcal{I} where $\Omega \rightarrow 0$, we see that the **causal properties of \mathcal{I}** are governed by the formula:

$$\hat{R} = -12g(\nabla\Omega, \nabla\Omega).$$

- Recall that the signature is $(-, +, +, +)$, so \mathcal{I} is **spacelike** if $\nabla\Omega$ is timelike, i.e. **if the Ricci scalar \hat{R} of the physical metric is positive, $\hat{R} > 0$.**

- A conformal transformation: $g \rightarrow \hat{g} = \frac{1}{\Omega^2}g$ results in the following transformation of the respective Ricci tensors:

$$\hat{R}_{\mu\nu} = \Omega^2 R_{\mu\nu} + 2\Omega \nabla_\nu \Omega_{,\mu} + g_{\mu\nu} (\Omega \square \Omega - 3g(\nabla \Omega, \nabla \Omega))$$

- Contracting we get the following transformation of the Ricci scalars:

$$\hat{R} = \Omega^2 R + 6\Omega \square \Omega - 12g(\nabla \Omega, \nabla \Omega).$$

- Interpreting \hat{g} as **the physical metric of spacetime**, whose conformal compactified metric g has \mathcal{I} where $\Omega \rightarrow 0$, we see that the **causal properties of \mathcal{I}** are governed by the formula:

$$\hat{R} = -12g(\nabla \Omega, \nabla \Omega).$$

- Recall that the signature is $(-, +, +, +)$, so \mathcal{I} is **spacelike** if $\nabla \Omega$ is timelike, i.e. **if the Ricci scalar \hat{R} of the physical metric is positive, $\hat{R} > 0$.**

- Using the Einstein equations satisfied by the metric \hat{g} ,

$$\hat{R}_{\mu\nu} - \frac{1}{2}\hat{R}\hat{g}_{\mu\nu} + \hat{\Lambda}\hat{g}_{\mu\nu} = \hat{T}_{\mu\nu}$$

or

$$\hat{R}_{\mu\nu} = \hat{T}_{\mu\nu} + (\hat{\Lambda} - \frac{1}{2}\hat{T})\hat{g}_{\mu\nu},$$

one can express this also in terms of the cosmological constant $\hat{\Lambda}$ of the physical spacetime and the trace of its energy momentum tensor \hat{T} :

$$4\hat{\Lambda} - \hat{T} > 0.$$

- This, in particular means that if close to \mathcal{I} the trace of the energy momentum \hat{T} vanishes, a **positive cosmological constant $\hat{\Lambda}$ makes \mathcal{I} spacelike.**

- Using the Einstein equations satisfied by the metric \hat{g} ,

$$\hat{R}_{\mu\nu} - \frac{1}{2}\hat{R}\hat{g}_{\mu\nu} + \hat{\Lambda}\hat{g}_{\mu\nu} = \hat{T}_{\mu\nu}$$

or

$$\hat{R}_{\mu\nu} = \hat{T}_{\mu\nu} + (\hat{\Lambda} - \frac{1}{2}\hat{T})\hat{g}_{\mu\nu},$$

one can express this also in terms of the cosmological constant $\hat{\Lambda}$ of the physical spacetime and the trace of its energy momentum tensor \hat{T} :

$$4\hat{\Lambda} - \hat{T} > 0.$$

- This, in particular means that if close to \mathcal{I} the trace of the energy momentum \hat{T} vanishes, a **positive cosmological constant $\hat{\Lambda}$ makes \mathcal{I} spacelike.**

- Using the Einstein equations satisfied by the metric \hat{g} ,

$$\hat{R}_{\mu\nu} - \frac{1}{2}\hat{R}\hat{g}_{\mu\nu} + \hat{\Lambda}\hat{g}_{\mu\nu} = \hat{T}_{\mu\nu}$$

or

$$\hat{R}_{\mu\nu} = \hat{T}_{\mu\nu} + (\hat{\Lambda} - \frac{1}{2}\hat{T})\hat{g}_{\mu\nu},$$

one can express this also in terms of the cosmological constant $\hat{\Lambda}$ of the physical spacetime and the trace of its energy momentum tensor \hat{T} :

$$4\hat{\Lambda} - \hat{T} > 0.$$

- This, in particular means that if close to \mathcal{I} the trace of the energy momentum \hat{T} vanishes, a **positive cosmological constant $\hat{\Lambda}$ makes \mathcal{I} spacelike.**

- Using the Einstein equations satisfied by the metric \hat{g} ,

$$\hat{R}_{\mu\nu} - \frac{1}{2}\hat{R}\hat{g}_{\mu\nu} + \hat{\Lambda}\hat{g}_{\mu\nu} = \hat{T}_{\mu\nu}$$

or

$$\hat{R}_{\mu\nu} = \hat{T}_{\mu\nu} + (\hat{\Lambda} - \frac{1}{2}\hat{T})\hat{g}_{\mu\nu},$$

one can express this also in terms of the cosmological constant $\hat{\Lambda}$ of the physical spacetime and the trace of its energy momentum tensor \hat{T} :

$$4\hat{\Lambda} - \hat{T} > 0.$$

- This, in particular means that if close to \mathcal{I} the trace of the energy momentum \hat{T} vanishes, a **positive cosmological constant $\hat{\Lambda}$ makes \mathcal{I} spacelike.**

- Using the Einstein equations satisfied by the metric \hat{g} ,

$$\hat{R}_{\mu\nu} - \frac{1}{2}\hat{R}\hat{g}_{\mu\nu} + \hat{\Lambda}\hat{g}_{\mu\nu} = \hat{T}_{\mu\nu}$$

or

$$\hat{R}_{\mu\nu} = \hat{T}_{\mu\nu} + (\hat{\Lambda} - \frac{1}{2}\hat{T})\hat{g}_{\mu\nu},$$

one can express this also in terms of the cosmological constant $\hat{\Lambda}$ of the physical spacetime and the trace of its energy momentum tensor \hat{T} :

$$4\hat{\Lambda} - \hat{T} > 0.$$

- This, in particular means that if close to \mathcal{S} the trace of the energy momentum \hat{T} vanishes, a **positive cosmological constant $\hat{\Lambda}$ makes \mathcal{S} spacelike.**

- Using the Einstein equations satisfied by the metric \hat{g} ,

$$\hat{R}_{\mu\nu} - \frac{1}{2}\hat{R}\hat{g}_{\mu\nu} + \hat{\Lambda}\hat{g}_{\mu\nu} = \hat{T}_{\mu\nu}$$

or

$$\hat{R}_{\mu\nu} = \hat{T}_{\mu\nu} + (\hat{\Lambda} - \frac{1}{2}\hat{T})\hat{g}_{\mu\nu},$$

one can express this also in terms of the cosmological constant $\hat{\Lambda}$ of the physical spacetime and the trace of its energy momentum tensor \hat{T} :

$$4\hat{\Lambda} - \hat{T} > 0.$$

- This, in particular means that if close to \mathcal{S} the trace of the energy momentum \hat{T} vanishes, a **positive cosmological constant $\hat{\Lambda}$ makes \mathcal{S} spacelike.**

- Using the Einstein equations satisfied by the metric \hat{g} ,

$$\hat{R}_{\mu\nu} - \frac{1}{2}\hat{R}\hat{g}_{\mu\nu} + \hat{\Lambda}\hat{g}_{\mu\nu} = \hat{T}_{\mu\nu}$$

or

$$\hat{R}_{\mu\nu} = \hat{T}_{\mu\nu} + (\hat{\Lambda} - \frac{1}{2}\hat{T})\hat{g}_{\mu\nu},$$

one can express this also in terms of the cosmological constant $\hat{\Lambda}$ of the physical spacetime and the trace of its energy momentum tensor \hat{T} :

$$4\hat{\Lambda} - \hat{T} > 0.$$

- This, in particular means that if close to \mathcal{I} the trace of the energy momentum \hat{T} vanishes, a **positive cosmological constant $\hat{\Lambda}$ makes \mathcal{I} spacelike.**

- **Perfect cosmological principle** (Copernicus ?): The Universe is the same **everywhere, in every direction, and at every moment of time.**
- A **General Relativity model** of such Universe is **due to Einstein** (1917):

$$M = \mathbb{R} \times \mathbb{S}^3, \text{ and } g_{Einst} = -dt^2 + \Omega^2 g_{\mathbb{S}^3},$$

with $g_{\mathbb{S}^3}$ is the standard metric on \mathbb{S}^3 of radius 1, and $\Omega = \text{const}$. This is the **Einstein's static Universe model.**

- The Ricci tensor for this spacetime is $Ricci = \frac{2}{\Omega^2} g_{\mathbb{S}^3}$. This satisfies Einstein's equations

$$Ricci - \frac{1}{2} Rg + \Lambda g = \mu u \otimes u,$$

with $u = -dt$, and $\Lambda = \frac{1}{\Omega^2}$, $\mu = \frac{2}{\Omega^2}$.

- Thus it is a solution to Einstein's field equations for homogeneous **dust** with **cosmological constant**.
- In particular the Ricci scalar $R = \frac{6}{\Omega^2}$ **is positive**. This means that Einstein's model has **spacelike** \mathcal{I} .

- **Perfect cosmological principle** (Copernicus ?): The Universe is the same **everywhere, in every direction, and at every moment of time.**
- A **General Relativity model** of such Universe is **due to Einstein** (1917):

$$M = \mathbb{R} \times \mathbb{S}^3, \text{ and } g_{Einst} = -dt^2 + \Omega^2 g_{\mathbb{S}^3},$$

with $g_{\mathbb{S}^3}$ is the standard metric on \mathbb{S}^3 of radius 1, and $\Omega = \text{const}$. This is the **Einstein's static Universe model**.

- The Ricci tensor for this spacetime is $Ricci = \frac{2}{\Omega^2} g_{\mathbb{S}^3}$. This satisfies Einstein's equations

$$Ricci - \frac{1}{2} Rg + \Lambda g = \mu u \otimes u,$$

with $u = -dt$, and $\Lambda = \frac{1}{\Omega^2}$, $\mu = \frac{2}{\Omega^2}$.

- Thus it is a solution to Einstein's field equations for homogeneous **dust** with **cosmological constant**.
- In particular the Ricci scalar $R = \frac{6}{\Omega^2}$ **is positive**. This means that Einstein's model has **spacelike** \mathcal{I} .

- **Perfect cosmological principle** (Copernicus ?): The Universe is the same **everywhere, in every direction, and at every moment of time.**
- **A General Relativity model** of such Universe is **due to Einstein** (1917):

$$M = \mathbb{R} \times \mathbb{S}^3, \text{ and } g_{Einst} = -dt^2 + \Omega^2 g_{\mathbb{S}^3},$$

with $g_{\mathbb{S}^3}$ is the standard metric on \mathbb{S}^3 of radius 1, and $\Omega = \text{const}$. This is the **Einstein's static Universe model.**

- The Ricci tensor for this spacetime is $Ricci = \frac{2}{\Omega^2} g_{\mathbb{S}^3}$. This satisfies Einstein's equations

$$Ricci - \frac{1}{2} Rg + \Lambda g = \mu u \otimes u,$$

with $u = -dt$, and $\Lambda = \frac{1}{\Omega^2}$, $\mu = \frac{2}{\Omega^2}$.

- Thus it is a solution to Einstein's field equations for homogeneous **dust** with **cosmological constant.**
- In particular the Ricci scalar $R = \frac{6}{\Omega^2}$ **is positive.** This means that Einstein's model has **spacelike** \mathcal{I} .

- **Perfect cosmological principle** (Copernicus?): The Universe is the same **everywhere, in every direction, and at every moment of time.**
- **A General Relativity model** of such Universe is **due to Einstein** (1917):

$$M = \mathbb{R} \times \mathbb{S}^3, \text{ and } g_{Einst} = -dt^2 + \Omega^2 g_{\mathbb{S}^3},$$

with $g_{\mathbb{S}^3}$ is the standard metric on \mathbb{S}^3 of radius 1, and $\Omega = \text{const}$. This is the **Einstein's static Universe model.**

- The Ricci tensor for this spacetime is $Ricci = \frac{2}{\Omega^2} g_{\mathbb{S}^3}$. This satisfies Einstein's equations

$$Ricci - \frac{1}{2} Rg + \Lambda g = \mu u \otimes u,$$

with $u = -dt$, and $\Lambda = \frac{1}{\Omega^2}$, $\mu = \frac{2}{\Omega^2}$.

- Thus it is a solution to Einstein's field equations for homogeneous **dust** with **cosmological constant.**
- In particular the Ricci scalar $R = \frac{6}{\Omega^2}$ **is positive.** This means that Einstein's model has **spacelike \mathcal{I} .**

- **Perfect cosmological principle** (Copernicus?): The Universe is the same **everywhere, in every direction, and at every moment of time.**
- A **General Relativity model** of such Universe is **due to Einstein** (1917):

$$M = \mathbb{R} \times \mathbb{S}^3, \text{ and } g_{Einst} = -dt^2 + \Omega^2 g_{\mathbb{S}^3},$$

with $g_{\mathbb{S}^3}$ is the standard metric on \mathbb{S}^3 of radius 1, and $\Omega = \text{const}$. This is the **Einstein's static Universe model**.

- The Ricci tensor for this spacetime is $Ricci = \frac{2}{\Omega^2} g_{\mathbb{S}^3}$. This satisfies Einstein's equations

$$Ricci - \frac{1}{2} Rg + \Lambda g = \mu u \otimes u,$$

with $u = -dt$, and $\Lambda = \frac{1}{\Omega^2}$, $\mu = \frac{2}{\Omega^2}$.

- Thus it is a solution to Einstein's field equations for homogeneous **dust** with **cosmological constant**.
- In particular the Ricci scalar $R = \frac{6}{\Omega^2}$ **is positive**. This means that Einstein's model has **spacelike** \mathcal{I} .

- **Perfect cosmological principle** (Copernicus?): The Universe is the same **everywhere, in every direction, and at every moment of time.**
- A **General Relativity model** of such Universe is **due to Einstein** (1917):

$$M = \mathbb{R} \times \mathbb{S}^3, \text{ and } g_{Einst} = -dt^2 + \Omega^2 g_{\mathbb{S}^3},$$

with $g_{\mathbb{S}^3}$ is the standard metric on \mathbb{S}^3 of radius 1, and $\Omega = \text{const.}$ This is the **Einstein's static Universe model.**

- The Ricci tensor for this spacetime is $Ricci = \frac{2}{\Omega^2} g_{\mathbb{S}^3}$. This satisfies Einstein's equations

$$Ricci - \frac{1}{2} Rg + \Lambda g = \mu u \otimes u,$$

with $u = -dt$, and $\Lambda = \frac{1}{\Omega^2}$, $\mu = \frac{2}{\Omega^2}$.

- Thus it is a solution to Einstein's field equations for homogeneous **dust** with **cosmological constant.**
- In particular the Ricci scalar $R = \frac{6}{\Omega^2}$ **is positive.** This means that Einstein's model has **spacelike \mathcal{I} .**

- **Perfect cosmological principle** (Copernicus ?): The Universe is the same **everywhere, in every direction, and at every moment of time.**
- A **General Relativity model** of such Universe is **due to Einstein** (1917):

$$M = \mathbb{R} \times \mathbb{S}^3, \text{ and } g_{Einst} = -dt^2 + \Omega^2 g_{\mathbb{S}^3},$$

with $g_{\mathbb{S}^3}$ is the standard metric on \mathbb{S}^3 of radius 1, and $\Omega = \text{const}$. This is the **Einstein's static Universe model.**

- The Ricci tensor for this spacetime is $Ricci = \frac{2}{\Omega^2} g_{\mathbb{S}^3}$. This satisfies Einstein's equations

$$Ricci - \frac{1}{2} Rg + \Lambda g = \mu u \otimes u,$$

with $u = -dt$, and $\Lambda = \frac{1}{\Omega^2}$, $\mu = \frac{2}{\Omega^2}$.

- Thus it is a solution to Einstein's field equations for homogeneous **dust** with **cosmological constant**.
- In particular the Ricci scalar $R = \frac{6}{\Omega^2}$ **is positive**. This means that Einstein's model has **spacelike** \mathcal{I} .

- **Perfect cosmological principle** (Copernicus ?): The Universe is the same **everywhere, in every direction, and at every moment of time.**
- A **General Relativity model** of such Universe is **due to Einstein** (1917):

$$M = \mathbb{R} \times \mathbb{S}^3, \text{ and } g_{Einst} = -dt^2 + \Omega^2 g_{\mathbb{S}^3},$$

with $g_{\mathbb{S}^3}$ is the standard metric on \mathbb{S}^3 of radius 1, and $\Omega = \text{const}$. This is the **Einstein's static Universe model.**

- The Ricci tensor for this spacetime is $Ricci = \frac{2}{\Omega^2} g_{\mathbb{S}^3}$. This satisfies Einstein's equations

$$Ricci - \frac{1}{2} Rg + \Lambda g = \mu u \otimes u,$$

with $u = -dt$, and $\Lambda = \frac{1}{\Omega^2}$, $\mu = \frac{2}{\Omega^2}$.

- Thus it is a solution to Einstein's field equations for homogeneous **dust** with **cosmological constant**.
- In particular the Ricci scalar $R = \frac{6}{\Omega^2}$ **is positive**. This means that Einstein's model has **spacelike** \mathcal{I} .

- **Perfect cosmological principle** (Copernicus ?): The Universe is the same **everywhere, in every direction, and at every moment of time.**
- A **General Relativity model** of such Universe is **due to Einstein** (1917):

$$M = \mathbb{R} \times \mathbb{S}^3, \text{ and } g_{Einst} = -dt^2 + \Omega^2 g_{\mathbb{S}^3},$$

with $g_{\mathbb{S}^3}$ is the standard metric on \mathbb{S}^3 of radius 1, and $\Omega = \text{const}$. This is the **Einstein's static Universe model.**

- The Ricci tensor for this spacetime is $Ricci = \frac{2}{\Omega^2} g_{\mathbb{S}^3}$. This satisfies Einstein's equations

$$Ricci - \frac{1}{2} Rg + \Lambda g = \mu u \otimes u,$$

with $u = -dt$, and $\Lambda = \frac{1}{\Omega^2}$, $\mu = \frac{2}{\Omega^2}$.

- Thus it is a solution to Einstein's field equations for homogeneous **dust** with **cosmological constant**.
- In particular the Ricci scalar $R = \frac{6}{\Omega^2}$ **is positive**. This means that Einstein's model has **spacelike \mathcal{I}** .

- **Perfect cosmological principle** (Copernicus?): The Universe is the same **everywhere, in every direction, and at every moment of time.**
- A **General Relativity model** of such Universe is **due to Einstein** (1917):

$$M = \mathbb{R} \times \mathbb{S}^3, \text{ and } g_{Einst} = -dt^2 + \Omega^2 g_{\mathbb{S}^3},$$

with $g_{\mathbb{S}^3}$ is the standard metric on \mathbb{S}^3 of radius 1, and $\Omega = \text{const}$. This is the **Einstein's static Universe model.**

- The Ricci tensor for this spacetime is $Ricci = \frac{2}{\Omega^2} g_{\mathbb{S}^3}$. This satisfies Einstein's equations

$$Ricci - \frac{1}{2} Rg + \Lambda g = \mu u \otimes u,$$

with $u = -dt$, and $\Lambda = \frac{1}{\Omega^2}$, $\mu = \frac{2}{\Omega^2}$.

- Thus it is a solution to Einstein's field equations for homogeneous **dust** with **cosmological constant**.
- In particular the Ricci scalar $R = \frac{6}{\Omega^2}$ **is positive**. This means that Einstein's model has **spacelike \mathcal{I}** .

- **Perfect cosmological principle** (Copernicus ?): The Universe is the same **everywhere, in every direction, and at every moment of time.**
- A **General Relativity model** of such Universe is **due to Einstein** (1917):

$$M = \mathbb{R} \times \mathbb{S}^3, \text{ and } g_{Einst} = -dt^2 + \Omega^2 g_{\mathbb{S}^3},$$

with $g_{\mathbb{S}^3}$ is the standard metric on \mathbb{S}^3 of radius 1, and $\Omega = \text{const}$. This is the **Einstein's static Universe model.**

- The Ricci tensor for this spacetime is $Ricci = \frac{2}{\Omega^2} g_{\mathbb{S}^3}$. This satisfies Einstein's equations

$$Ricci - \frac{1}{2} Rg + \Lambda g = \mu u \otimes u,$$

with $u = -dt$, and $\Lambda = \frac{1}{\Omega^2}$, $\mu = \frac{2}{\Omega^2}$.

- Thus it is a solution to Einstein's field equations for homogeneous **dust** with **cosmological constant.**
- In particular the Ricci scalar $R = \frac{6}{\Omega^2}$ **is positive.** This means that Einstein's model has **spacelike** \mathcal{I} .

- **Perfect cosmological principle** (Copernicus ?): The Universe is the same **everywhere, in every direction, and at every moment of time.**
- A **General Relativity model** of such Universe is **due to Einstein** (1917):

$$M = \mathbb{R} \times \mathbb{S}^3, \text{ and } g_{Einst} = -dt^2 + \Omega^2 g_{\mathbb{S}^3},$$

with $g_{\mathbb{S}^3}$ is the standard metric on \mathbb{S}^3 of radius 1, and $\Omega = \text{const}$. This is the **Einstein's static Universe model.**

- The Ricci tensor for this spacetime is $Ricci = \frac{2}{\Omega^2} g_{\mathbb{S}^3}$. This satisfies Einstein's equations

$$Ricci - \frac{1}{2} Rg + \Lambda g = \mu u \otimes u,$$

with $u = -dt$, and $\Lambda = \frac{1}{\Omega^2}$, $\mu = \frac{2}{\Omega^2}$.

- Thus it is a solution to Einstein's field equations for homogeneous **dust** with **cosmological constant.**
- In particular the Ricci scalar $R = \frac{6}{\Omega^2}$ **is positive.** This means that Einstein's model has **spacelike \mathcal{I} .**

- **Perfect cosmological principle** (Copernicus ?): The Universe is the same **everywhere**, **in every direction**, and **at every moment of time**.
- A **General Relativity model** of such Universe is **due to Einstein** (1917):

$$M = \mathbb{R} \times \mathbb{S}^3, \text{ and } g_{Einst} = -dt^2 + \Omega^2 g_{\mathbb{S}^3},$$

with $g_{\mathbb{S}^3}$ is the standard metric on \mathbb{S}^3 of radius 1, and $\Omega = \text{const}$. This is the **Einstein's static Universe model**.

- The Ricci tensor for this spacetime is $Ricci = \frac{2}{\Omega^2} g_{\mathbb{S}^3}$. This satisfies Einstein's equations

$$Ricci - \frac{1}{2} Rg + \Lambda g = \mu u \otimes u,$$

with $u = -dt$, and $\Lambda = \frac{1}{\Omega^2}$, $\mu = \frac{2}{\Omega^2}$.

- Thus it is a solution to Einstein's field equations for homogeneous **dust** with **cosmological constant**.
- In particular the Ricci scalar $R = \frac{6}{\Omega^2}$ **is positive**. This means that Einstein's model has **spacelike** \mathcal{I} .

- After **Lemaître-Hubble's** discovery of **expansion of Galaxies** in 1927-29, it became clear that the **perfect cosmological principle** is **not correct**, and it got changed into: The **Universe is the same everywhere and in every direction**.
- The models of the Universe compatible with this principle, due to **Friedman, Lemaître, Robertson and Walker** (FLRW) are such that the **spacetime manifold M** is a **bundle $S \rightarrow M \rightarrow \mathbb{R}$** , with $S = \mathbb{H}^3$, or \mathbb{R}^3 , or \mathbb{S}^3 - the **spaces of constant curvature**, with the **spacetime metric $g = -dt^2 + \Omega^2(t)g_S$** , in which g_S is either **the standard metric on the hyperbolic space \mathbb{H}^3** , or **the flat metric on \mathbb{R}^3** , or **the standard metric on the unit sphere \mathbb{S}^3** .

- After **Lemaître-Hubble's** discovery of **expansion of Galaxies** in 1927-29, it became clear that the **perfect cosmological principle** is **not correct**, and it got changed into: The **Universe is the same everywhere and in every direction**.
- The models of the Universe compatible with this principle, due to **Friedman, Lemaître, Robertson and Walker** (FLRW) are such that the **spacetime manifold M** is a **bundle $S \rightarrow M \rightarrow \mathbb{R}$** , with $S = \mathbb{H}^3$, or \mathbb{R}^3 , or \mathbb{S}^3 - the **spaces of constant curvature**, with the **spacetime metric $g = -dt^2 + \Omega^2(t)g_S$** , in which g_S is either the **standard metric on the hyperbolic space \mathbb{H}^3** , or the **flat metric on \mathbb{R}^3** , or the **standard metric on the unit sphere \mathbb{S}^3** .

- After **Lemaître-Hubble's** discovery of **expansion of Galaxies** in 1927-29, it became clear that the **perfect cosmological principle** is **not correct**, and it got changed into: The **Universe is the same everywhere** and **in every direction**.
- The models of the Universe compatible with this principle, due to **Friedman, Lemaître, Robertson and Walker** (FLRW) are such that the **spacetime manifold M** is a **bundle $S \rightarrow M \rightarrow \mathbb{R}$** , with $S = \mathbb{H}^3$, or \mathbb{R}^3 , or \mathbb{S}^3 - the **spaces of constant curvature**, with the **spacetime metric $g = -dt^2 + \Omega^2(t)g_S$** , in which g_S is either the **standard metric on the hyperbolic space \mathbb{H}^3** , or the **flat metric on \mathbb{R}^3** , or the **standard metric on the unit sphere \mathbb{S}^3** .

- After **Lemaître-Hubble's** discovery of **expansion of Galaxies** in 1927-29, it became clear that the **perfect cosmological principle** is **not correct**, and it got changed into: The **Universe is the same everywhere** and **in every direction**.
- The models of the Universe compatible with this principle, due to **Friedman, Lemaître, Robertson and Walker** (FLRW) are such that the **spacetime manifold M** is a **bundle $S \rightarrow M \rightarrow \mathbb{R}$** , with $S = \mathbb{H}^3$, or \mathbb{R}^3 , or \mathbb{S}^3 - the **spaces of constant curvature**, with the **spacetime metric $g = -dt^2 + \Omega^2(t)g_S$** , in which g_S is either the **standard metric on the hyperbolic space \mathbb{H}^3** , or the **flat metric on \mathbb{R}^3** , or the **standard metric on the unit sphere \mathbb{S}^3** .

- After **Lemaître-Hubble's** discovery of **expansion of Galaxies** in 1927-29, it became clear that the **perfect cosmological principle** is **not correct**, and it got changed into: The **Universe is the same everywhere** and **in every direction**.
- The models of the Universe compatible with this principle, due to **Friedman, Lemaître, Robertson and Walker** (FLRW) are such that the **spacetime manifold M** is a **bundle $S \rightarrow M \rightarrow \mathbb{R}$** , with $S = \mathbb{H}^3$, or \mathbb{R}^3 , or \mathbb{S}^3 - the **spaces of constant curvature**, with the **spacetime metric $g = -dt^2 + \Omega^2(t)g_S$** , in which g_S is either the **standard metric on the hyperbolic space \mathbb{H}^3** , or the **flat metric on \mathbb{R}^3** , or the **standard metric on the unit sphere \mathbb{S}^3** .

- After **Lemaître-Hubble's** discovery of **expansion of Galaxies** in 1927-29, it became clear that the **perfect cosmological principle** is **not correct**, and it got changed into: The **Universe is the same everywhere** and **in every direction**.
- The models of the Universe compatible with this principle, due to **Friedman, Lemaître, Robertson** and **Walker** (FLRW) are such that the **spacetime manifold M** is a **bundle $S \rightarrow M \rightarrow \mathbb{R}$** , with $S = \mathbb{H}^3$, or \mathbb{R}^3 , or \mathbb{S}^3 - the **spaces of constant curvature**, with the **spacetime metric $g = -dt^2 + \Omega^2(t)g_S$** , in which g_S is either the **standard metric on the hyperbolic space \mathbb{H}^3** , or the **flat metric on \mathbb{R}^3** , or the **standard metric on the unit sphere \mathbb{S}^3** .

- After **Lemaître-Hubble's** discovery of **expansion of Galaxies** in 1927-29, it became clear that the **perfect cosmological principle** is **not correct**, and it got changed into: The **Universe is the same everywhere** and **in every direction**.
- The models of the Universe compatible with this principle, due to **Friedman, Lemaître, Robertson and Walker** (FLRW) are such that the **spacetime manifold M** is a **bundle $S \rightarrow M \rightarrow \mathbb{R}$** , with $S = \mathbb{H}^3$, or \mathbb{R}^3 , or \mathbb{S}^3 - the **spaces of constant curvature**, with the **spacetime metric $g = -dt^2 + \Omega^2(t)g_S$** , in which g_S is either the **standard metric on the hyperbolic space \mathbb{H}^3** , or the **flat metric on \mathbb{R}^3** , or the **standard metric on the unit sphere \mathbb{S}^3** .

- After **Lemaître-Hubble's** discovery of **expansion of Galaxies** in 1927-29, it became clear that the **perfect cosmological principle** is **not correct**, and it got changed into: The **Universe is the same everywhere** and **in every direction**.
- The models of the Universe compatible with this principle, due to **Friedman, Lemaître, Robertson and Walker** (FLRW) are such that the **spacetime manifold M** is a **bundle $S \rightarrow M \rightarrow \mathbb{R}$** , with $S = \mathbb{H}^3$, or \mathbb{R}^3 , or \mathbb{S}^3 - the **spaces of constant curvature**, with the **spacetime metric $g = -dt^2 + \Omega^2(t)g_S$** , in which g_S is either the **standard metric on the hyperbolic space \mathbb{H}^3** , or the **flat metric on \mathbb{R}^3** , or the **standard metric on the unit sphere \mathbb{S}^3** .

- After **Lemaître-Hubble's** discovery of **expansion of Galaxies** in 1927-29, it became clear that the **perfect cosmological principle** is **not correct**, and it got changed into: The **Universe is the same everywhere** and **in every direction**.
- The models of the Universe compatible with this principle, due to **Friedman, Lemaître, Robertson and Walker** (FLRW) are such that the **spacetime manifold M** is a **bundle $S \rightarrow M \rightarrow \mathbb{R}$** , with $S = \mathbb{H}^3$, or \mathbb{R}^3 , or \mathbb{S}^3 - the **spaces of constant curvature**, with the **spacetime metric $g = -dt^2 + \Omega^2(t)g_S$** , in which g_S is either the **standard metric on the hyperbolic space \mathbb{H}^3** , or the **flat metric on \mathbb{R}^3** , or the **standard metric on the unit sphere \mathbb{S}^3** .

- After **Lemaître-Hubble's** discovery of **expansion of Galaxies** in 1927-29, it became clear that the **perfect cosmological principle** is **not correct**, and it got changed into: The **Universe is the same everywhere** and **in every direction**.
- The models of the Universe compatible with this principle, due to **Friedman, Lemaître, Robertson and Walker** (FLRW) are such that the **spacetime manifold M** is a **bundle $S \rightarrow M \rightarrow \mathbb{R}$** , with $S = \mathbb{H}^3$, or \mathbb{R}^3 , or \mathbb{S}^3 - the **spaces of constant curvature**, with the **spacetime metric $g = -dt^2 + \Omega^2(t)g_S$** , in which g_S is either the **standard metric on the hyperbolic space \mathbb{H}^3** , or the **flat metric on \mathbb{R}^3** , or the **standard metric on the unit sphere \mathbb{S}^3** .

- After **Lemaître-Hubble's** discovery of **expansion of Galaxies** in 1927-29, it became clear that the **perfect cosmological principle** is **not correct**, and it got changed into: The **Universe is the same everywhere** and **in every direction**.
- The models of the Universe compatible with this principle, due to **Friedman, Lemaître, Robertson and Walker** (FLRW) are such that the **spacetime manifold M** is a **bundle $S \rightarrow M \rightarrow \mathbb{R}$** , with $S = \mathbb{H}^3$, or \mathbb{R}^3 , or \mathbb{S}^3 - the **spaces of constant curvature**, with the **spacetime metric $g = -dt^2 + \Omega^2(t)g_S$** , in which g_S is either the **standard metric on the hyperbolic space \mathbb{H}^3** , or the **flat metric on \mathbb{R}^3** , or the **standard metric on the unit sphere \mathbb{S}^3** .

- After **Lemaître-Hubble's** discovery of **expansion of Galaxies** in 1927-29, it became clear that the **perfect cosmological principle** is **not correct**, and it got changed into: The **Universe is the same everywhere** and **in every direction**.
- The models of the Universe compatible with this principle, due to **Friedman, Lemaître, Robertson and Walker** (FLRW) are such that the **spacetime manifold M** is a **bundle $S \rightarrow M \rightarrow \mathbb{R}$** , with $S = \mathbb{H}^3$, or \mathbb{R}^3 , or \mathbb{S}^3 - the **spaces of constant curvature**, with the **spacetime metric $g = -dt^2 + \Omega^2(t)g_S$** , in which g_S is either the **standard metric on the hyperbolic space \mathbb{H}^3** , or the **flat metric on \mathbb{R}^3** , or the **standard metric on the unit sphere \mathbb{S}^3** .

- After **Lemaître-Hubble's** discovery of **expansion of Galaxies** in 1927-29, it became clear that the **perfect cosmological principle** is **not correct**, and it got changed into: The **Universe is the same everywhere** and **in every direction**.
- The models of the Universe compatible with this principle, due to **Friedman, Lemaître, Robertson and Walker** (FLRW) are such that the **spacetime manifold M** is a **bundle $S \rightarrow M \rightarrow \mathbb{R}$** , with $S = \mathbb{H}^3$, or \mathbb{R}^3 , or S^3 - the **spaces of constant curvature**, with the **spacetime metric $g = -dt^2 + \Omega^2(t)g_S$** , in which g_S is either the **standard metric on the hyperbolic space \mathbb{H}^3** , or the **flat metric on \mathbb{R}^3** , or the **standard metric on the unit sphere S^3** .

- After **Lemaître-Hubble's** discovery of **expansion of Galaxies** in 1927-29, it became clear that the **perfect cosmological principle** is **not correct**, and it got changed into: The **Universe is the same everywhere** and **in every direction**.
- The models of the Universe compatible with this principle, due to **Friedman, Lemaître, Robertson and Walker** (FLRW) are such that the **spacetime manifold M** is a **bundle $S \rightarrow M \rightarrow \mathbb{R}$** , with $S = \mathbb{H}^3$, or \mathbb{R}^3 , or S^3 - the **spaces of constant curvature**, with the **spacetime metric $g = -dt^2 + \Omega^2(t)g_S$** , in which g_S is either the **standard metric on the hyperbolic space \mathbb{H}^3** , or the **flat metric on \mathbb{R}^3** , or the **standard metric on the unit sphere S^3** .

- After **Lemaître-Hubble's** discovery of **expansion of Galaxies** in 1927-29, it became clear that the **perfect cosmological principle** is **not correct**, and it got changed into: The **Universe is the same everywhere** and **in every direction**.
- The models of the Universe compatible with this principle, due to **Friedman, Lemaître, Robertson and Walker** (FLRW) are such that the **spacetime manifold M** is a **bundle $S \rightarrow M \rightarrow \mathbb{R}$** , with $S = \mathbb{H}^3$, or \mathbb{R}^3 , or S^3 - the **spaces of constant curvature**, with the **spacetime metric $g = -dt^2 + \Omega^2(t)g_S$** , in which g_S is either the **standard metric on the hyperbolic space \mathbb{H}^3** , or the **flat metric on \mathbb{R}^3** , or the **standard metric on the unit sphere S^3** .

- After **Lemaître-Hubble's** discovery of **expansion of Galaxies** in 1927-29, it became clear that the **perfect cosmological principle** is **not correct**, and it got changed into: The **Universe is the same everywhere** and **in every direction**.
- The models of the Universe compatible with this principle, due to **Friedman, Lemaître, Robertson and Walker** (FLRW) are such that the **spacetime manifold M** is a **bundle $S \rightarrow M \rightarrow \mathbb{R}$** , with $S = \mathbb{H}^3$, or \mathbb{R}^3 , or S^3 - the **spaces of constant curvature**, with the **spacetime metric $g = -dt^2 + \Omega^2(t)g_S$** , in which g_S is either **the standard metric on the hyperbolic space \mathbb{H}^3** , or **the flat metric on \mathbb{R}^3** , or **the standard metric on the unit sphere S^3** .

- After **Lemaître-Hubble's** discovery of **expansion of Galaxies** in 1927-29, it became clear that the **perfect cosmological principle** is **not correct**, and it got changed into: The **Universe is the same everywhere and in every direction**.
- The models of the Universe compatible with this principle, due to **Friedman, Lemaître, Robertson and Walker** (FLRW) are such that the **spacetime manifold M** is a **bundle $S \rightarrow M \rightarrow \mathbb{R}$** , with $S = \mathbb{H}^3$, or \mathbb{R}^3 , or S^3 - the **spaces of constant curvature**, with the **spacetime metric $g = -dt^2 + \Omega^2(t)g_S$** , in which g_S is either **the standard metric on the hyperbolic space \mathbb{H}^3** , or **the flat metric on \mathbb{R}^3** , or **the standard metric on the unit sphere S^3** .

- After **Lemaître-Hubble's** discovery of **expansion of Galaxies** in 1927-29, it became clear that the **perfect cosmological principle** is **not correct**, and it got changed into: The **Universe is the same everywhere and in every direction**.
- The models of the Universe compatible with this principle, due to **Friedman, Lemaître, Robertson and Walker** (FLRW) are such that the **spacetime manifold M** is a **bundle $S \rightarrow M \rightarrow \mathbb{R}$** , with $S = \mathbb{H}^3$, or \mathbb{R}^3 , or \mathbb{S}^3 - the **spaces of constant curvature**, with the **spacetime metric $g = -dt^2 + \Omega^2(t)g_S$** , in which g_S is either **the standard metric on the hyperbolic space \mathbb{H}^3** , or **the flat metric on \mathbb{R}^3** , or **the standard metric on the unit sphere \mathbb{S}^3** .

- **Universe spacetime:** $S \rightarrow M \rightarrow \mathbb{R}$, with $g = -dt^2 + \Omega^2(t)g_S$.
- For example, there exists local coordinates (x, y, z) on S such that $g_S = \frac{dx^2 + dy^2 + dz^2}{(1 + \frac{\kappa}{4}(x^2 + y^2 + z^2))^2}$, and $\kappa = -1, 0, 1$ corresponds to \mathbb{H}^3 , \mathbb{R}^3 and \mathbb{S}^3 , respectively.
- Special solutions:
 - $\Omega(t) = \text{const}$ and $\kappa = 1$: **Einstein's static Universe**; homogeneous dust with positive Λ .
 - $\Omega(t) = \alpha \cosh(\frac{t}{\alpha})$ and $\kappa = 1$: **deSitter Universe**; vacuum solution with $\Lambda = \frac{3}{\alpha^2}$; **spacetime** of constant curvature.
 - $\Omega(t)$ satisfies $\Omega(\Omega'^2 + \kappa) = \frac{M}{4\pi} + \frac{1}{3}\Lambda\Omega^3$, with $M = \text{const} > 0$, $\Lambda = \text{const}$; **Friedman-Lemaître Universe** with Λ of arbitrary sign, filled with **dust** of energy density $\mu = \frac{M}{\frac{4}{3}\pi\Omega^3}$, comoving with 4-velocity $u = -dt$. For the Ricci scalar being positive one needs $\Lambda > -\mu/4$.
- **Note that all FLRW metrics are conformally flat!**

- **Universe spacetime:** $S \rightarrow M \rightarrow \mathbb{R}$, with $g = -dt^2 + \Omega^2(t)g_S$.
- For example, there exists local coordinates (x, y, z) on S such that $g_S = \frac{dx^2 + dy^2 + dz^2}{(1 + \frac{\varkappa}{4}(x^2 + y^2 + z^2))^2}$, and $\varkappa = -1, 0, 1$ corresponds to \mathbb{H}^3 , \mathbb{R}^3 and \mathbb{S}^3 , respectively.
- Special solutions:
 - $\Omega(t) = \text{const}$ and $\varkappa = 1$: **Einstein's static Universe**; homogeneous dust with positive Λ .
 - $\Omega(t) = \alpha \cosh(\frac{t}{\alpha})$ and $\varkappa = 1$: **deSitter Universe**; vacuum solution with $\Lambda = \frac{3}{\alpha^2}$; **spacetime** of constant curvature.
 - $\Omega(t)$ satisfies $\Omega(\Omega'^2 + \varkappa) = \frac{M}{4\pi} + \frac{1}{3}\Lambda\Omega^3$, with $M = \text{const} > 0$, $\Lambda = \text{const}$; **Friedman-Lemaître Universe** with Λ of arbitrary sign, filled with **dust** of energy density $\mu = \frac{M}{\frac{4}{3}\pi\Omega^3}$, comoving with 4-velocity $u = -dt$. For the Ricci scalar being positive one needs $\Lambda > -\mu/4$.
- **Note that all FLRW metrics are conformally flat!**

- **Universe spacetime:** $S \rightarrow M \rightarrow \mathbb{R}$, with $g = -dt^2 + \Omega^2(t)g_S$.
- For example, there exists local coordinates (x, y, z) on S such that $g_S = \frac{dx^2 + dy^2 + dz^2}{(1 + \frac{\varkappa}{4}(x^2 + y^2 + z^2))^2}$, and $\varkappa = -1, 0, 1$ corresponds to \mathbb{H}^3 , \mathbb{R}^3 and \mathbb{S}^3 , respectively.
- Special solutions:
 - $\Omega(t) = \text{const}$ and $\varkappa = 1$: **Einstein's static Universe**; homogeneous dust with positive Λ .
 - $\Omega(t) = \alpha \cosh(\frac{t}{\alpha})$ and $\varkappa = 1$: **deSitter Universe**; vacuum solution with $\Lambda = \frac{3}{\alpha^2}$; **spacetime** of constant curvature.
 - $\Omega(t)$ satisfies $\Omega(\Omega'^2 + \varkappa) = \frac{M}{4\pi} + \frac{1}{3}\Lambda\Omega^3$, with $M = \text{const} > 0$, $\Lambda = \text{const}$; **Friedman-Lemaître Universe** with Λ of arbitrary sign, filled with **dust** of energy density $\mu = \frac{M}{\frac{4}{3}\pi\Omega^3}$, comoving with 4-velocity $u = -dt$. For the Ricci scalar being positive one needs $\Lambda > -\mu/4$.
- **Note that all FLRW metrics are conformally flat!**

- **Universe spacetime:** $S \rightarrow M \rightarrow \mathbb{R}$, with $g = -dt^2 + \Omega^2(t)g_S$.
- For example, there exists local coordinates (x, y, z) on S such that $g_S = \frac{dx^2 + dy^2 + dz^2}{(1 + \frac{\varkappa}{4}(x^2 + y^2 + z^2))^2}$, and $\varkappa = -1, 0, 1$ corresponds to \mathbb{H}^3 , \mathbb{R}^3 and \mathbb{S}^3 , respectively.
- Special solutions:
 - $\Omega(t) = \text{const}$ and $\varkappa = 1$: **Einstein's static Universe**; homogeneous dust with positive Λ .
 - $\Omega(t) = \alpha \cosh(\frac{t}{\alpha})$ and $\varkappa = 1$: **deSitter Universe**; vacuum solution with $\Lambda = \frac{3}{\alpha^2}$; **spacetime** of constant curvature.
 - $\Omega(t)$ satisfies $\Omega(\Omega'^2 + \varkappa) = \frac{M}{4\pi} + \frac{1}{3}\Lambda\Omega^3$, with $M = \text{const} > 0$, $\Lambda = \text{const}$; **Friedman-Lemaître Universe** with Λ of arbitrary sign, filled with **dust** of energy density $\mu = \frac{M}{\frac{4}{3}\pi\Omega^3}$, comoving with 4-velocity $u = -dt$. For the Ricci scalar being positive one needs $\Lambda > -\mu/4$.
- **Note that all FLRW metrics are conformally flat!**

- **Universe spacetime:** $S \rightarrow M \rightarrow \mathbb{R}$, with $g = -dt^2 + \Omega^2(t)g_S$.
- For example, there exists local coordinates (x, y, z) on S such that $g_S = \frac{dx^2 + dy^2 + dz^2}{(1 + \frac{\varkappa}{4}(x^2 + y^2 + z^2))^2}$, and $\varkappa = -1, 0, 1$ corresponds to \mathbb{H}^3 , \mathbb{R}^3 and \mathbb{S}^3 , respectively.
- Special solutions:
 - $\Omega(t) = \text{const}$ and $\varkappa = 1$: **Einstein's static Universe**; homogeneous dust with positive Λ .
 - $\Omega(t) = \alpha \cosh(\frac{t}{\alpha})$ and $\varkappa = 1$: **deSitter Universe**; vacuum solution with $\Lambda = \frac{3}{\alpha^2}$; **spacetime** of constant curvature.
 - $\Omega(t)$ satisfies $\Omega(\Omega'^2 + \varkappa) = \frac{M}{4\pi} + \frac{1}{3}\Lambda\Omega^3$, with $M = \text{const} > 0$, $\Lambda = \text{const}$; **Friedman-Lemaître Universe** with Λ of arbitrary sign, filled with **dust** of energy density $\mu = \frac{M}{\frac{4}{3}\pi\Omega^3}$, comoving with 4-velocity $u = -dt$. For the Ricci scalar being positive one needs $\Lambda > -\mu/4$.
- **Note that all FLRW metrics are conformally flat!**

- **Universe spacetime:** $S \rightarrow M \rightarrow \mathbb{R}$, with $g = -dt^2 + \Omega^2(t)g_S$.
- For example, there exists local coordinates (x, y, z) on S such that $g_S = \frac{dx^2 + dy^2 + dz^2}{(1 + \frac{\varkappa}{4}(x^2 + y^2 + z^2))^2}$, and $\varkappa = -1, 0, 1$ corresponds to \mathbb{H}^3 , \mathbb{R}^3 and \mathbb{S}^3 , respectively.
- Special solutions:
 - $\Omega(t) = \text{const}$ and $\varkappa = 1$: **Einstein's static Universe**; homogeneous dust with positive Λ .
 - $\Omega(t) = \alpha \cosh(\frac{t}{\alpha})$ and $\varkappa = 1$: **deSitter Universe**; vacuum solution with $\Lambda = \frac{3}{\alpha^2}$; **spacetime** of constant curvature.
 - $\Omega(t)$ satisfies $\Omega(\Omega'^2 + \varkappa) = \frac{M}{4\pi} + \frac{1}{3}\Lambda\Omega^3$, with $M = \text{const} > 0$, $\Lambda = \text{const}$; **Friedman-Lemaître Universe** with Λ of arbitrary sign, filled with **dust** of energy density $\mu = \frac{M}{\frac{4}{3}\pi\Omega^3}$, comoving with 4-velocity $u = -dt$. For the Ricci scalar being positive one needs $\Lambda > -\mu/4$.
- **Note that all FLRW metrics are conformally flat!**

- **Universe spacetime:** $S \rightarrow M \rightarrow \mathbb{R}$, with $g = -dt^2 + \Omega^2(t)g_S$.
- For example, there exists local coordinates (x, y, z) on S such that $g_S = \frac{dx^2 + dy^2 + dz^2}{(1 + \frac{\varkappa}{4}(x^2 + y^2 + z^2))^2}$, and $\varkappa = -1, 0, 1$ corresponds to \mathbb{H}^3 , \mathbb{R}^3 and \mathbb{S}^3 , respectively.
- Special solutions:
 - $\Omega(t) = \text{const}$ and $\varkappa = 1$: **Einstein's static Universe**; homogeneous dust with positive Λ .
 - $\Omega(t) = \alpha \cosh(\frac{t}{\alpha})$ and $\varkappa = 1$: **deSitter Universe**; vacuum solution with $\Lambda = \frac{3}{\alpha^2}$; **spacetime** of constant curvature.
 - $\Omega(t)$ satisfies $\Omega(\Omega'^2 + \varkappa) = \frac{M}{4\pi} + \frac{1}{3}\Lambda\Omega^3$, with $M = \text{const} > 0$, $\Lambda = \text{const}$; **Friedman-Lemaître Universe** with Λ of arbitrary sign, filled with **dust** of energy density $\mu = \frac{M}{\frac{4}{3}\pi\Omega^3}$, comoving with 4-velocity $u = -dt$. For the Ricci scalar being positive one needs $\Lambda > -\mu/4$.
- **Note that all FLRW metrics are conformally flat!**

- **Universe spacetime:** $S \rightarrow M \rightarrow \mathbb{R}$, with $g = -dt^2 + \Omega^2(t)g_S$.
- For example, there exists local coordinates (x, y, z) on S such that $g_S = \frac{dx^2 + dy^2 + dz^2}{(1 + \frac{\kappa}{4}(x^2 + y^2 + z^2))^2}$, and $\kappa = -1, 0, 1$ corresponds to \mathbb{H}^3 , \mathbb{R}^3 and \mathbb{S}^3 , respectively.
- Special solutions:
 - $\Omega(t) = \text{const}$ and $\kappa = 1$: **Einstein's static Universe**; homogeneous dust with positive Λ .
 - $\Omega(t) = \alpha \cosh(\frac{t}{\alpha})$ and $\kappa = 1$: **deSitter Universe**; vacuum solution with $\Lambda = \frac{3}{\alpha^2}$; **spacetime** of constant curvature.
 - $\Omega(t)$ satisfies $\Omega(\Omega'^2 + \kappa) = \frac{M}{4\pi} + \frac{1}{3}\Lambda\Omega^3$, with $M = \text{const} > 0$, $\Lambda = \text{const}$; **Friedman-Lemaître Universe** with Λ of arbitrary sign, filled with **dust** of energy density $\mu = \frac{M}{\frac{4}{3}\pi\Omega^3}$, comoving with 4-velocity $u = -dt$. For the Ricci scalar being positive one needs $\Lambda > -\mu/4$.
- **Note that all FLRW metrics are conformally flat!**

- **Universe spacetime:** $S \rightarrow M \rightarrow \mathbb{R}$, with $g = -dt^2 + \Omega^2(t)g_S$.
- For example, there exists local coordinates (x, y, z) on S such that $g_S = \frac{dx^2 + dy^2 + dz^2}{(1 + \frac{\kappa}{4}(x^2 + y^2 + z^2))^2}$, and $\kappa = -1, 0, 1$ corresponds to \mathbb{H}^3 , \mathbb{R}^3 and \mathbb{S}^3 , respectively.
- Special solutions:
 - $\Omega(t) = \text{const}$ and $\kappa = 1$: **Einstein's static Universe**; homogeneous dust with positive Λ .
 - $\Omega(t) = \alpha \cosh(\frac{t}{\alpha})$ and $\kappa = 1$: **deSitter Universe**; vacuum solution with $\Lambda = \frac{3}{\alpha^2}$; **spacetime** of constant curvature.
 - $\Omega(t)$ satisfies $\Omega(\Omega'^2 + \kappa) = \frac{M}{4\pi} + \frac{1}{3}\Lambda\Omega^3$, with $M = \text{const} > 0$, $\Lambda = \text{const}$; **Friedman-Lemaître Universe** with Λ of arbitrary sign, filled with **dust** of energy density $\mu = \frac{M}{\frac{4}{3}\pi\Omega^3}$, comoving with 4-velocity $u = -dt$. For the Ricci scalar being positive one needs $\Lambda > -\mu/4$.
- **Note that all FLRW metrics are conformally flat!**

- **Universe spacetime:** $S \rightarrow M \rightarrow \mathbb{R}$, with $g = -dt^2 + \Omega^2(t)g_S$.
- For example, there exists local coordinates (x, y, z) on S such that $g_S = \frac{dx^2 + dy^2 + dz^2}{(1 + \frac{\kappa}{4}(x^2 + y^2 + z^2))^2}$, and $\kappa = -1, 0, 1$ corresponds to \mathbb{H}^3 , \mathbb{R}^3 and \mathbb{S}^3 , respectively.
- Special solutions:
 - $\Omega(t) = \text{const}$ and $\kappa = 1$: **Einstein's static Universe**; homogeneous dust with positive Λ .
 - $\Omega(t) = \alpha \cosh(\frac{t}{\alpha})$ and $\kappa = 1$: **deSitter Universe**; vacuum solution with $\Lambda = \frac{3}{\alpha^2}$; **spacetime** of constant curvature.
 - $\Omega(t)$ satisfies $\Omega(\Omega'^2 + \kappa) = \frac{M}{4\pi} + \frac{1}{3}\Lambda\Omega^3$, with $M = \text{const} > 0$, $\Lambda = \text{const}$; **Friedman-Lemaître Universe** with Λ of arbitrary sign, filled with **dust** of energy density $\mu = \frac{M}{\frac{4}{3}\pi\Omega^3}$, comoving with 4-velocity $u = -dt$. For the Ricci scalar being positive one needs $\Lambda > -\mu/4$.
- **Note that all FLRW metrics are conformally flat!**

- **Universe spacetime:** $S \rightarrow M \rightarrow \mathbb{R}$, with $g = -dt^2 + \Omega^2(t)g_S$.
- For example, there exists local coordinates (x, y, z) on S such that $g_S = \frac{dx^2 + dy^2 + dz^2}{(1 + \frac{\kappa}{4}(x^2 + y^2 + z^2))^2}$, and $\kappa = -1, 0, 1$ corresponds to \mathbb{H}^3 , \mathbb{R}^3 and \mathbb{S}^3 , respectively.
- Special solutions:
 - $\Omega(t) = \text{const}$ and $\kappa = 1$: **Einstein's static Universe**; homogeneous dust with positive Λ .
 - $\Omega(t) = \alpha \cosh(\frac{t}{\alpha})$ and $\kappa = 1$: **deSitter Universe**; vacuum solution with $\Lambda = \frac{3}{\alpha^2}$; **spacetime** of constant curvature.
 - $\Omega(t)$ satisfies $\Omega(\Omega'^2 + \kappa) = \frac{M}{4\pi} + \frac{1}{3}\Lambda\Omega^3$, with $M = \text{const} > 0$, $\Lambda = \text{const}$; **Friedman-Lemaître Universe** with Λ of arbitrary sign, filled with **dust** of energy density $\mu = \frac{M}{\frac{4}{3}\pi\Omega^3}$, comoving with 4-velocity $u = -dt$. For the Ricci scalar being positive one needs $\Lambda > -\mu/4$.
- **Note that all FLRW metrics are conformally flat!**

- **Universe spacetime:** $S \rightarrow M \rightarrow \mathbb{R}$, with $g = -dt^2 + \Omega^2(t)g_S$.
- For example, there exists local coordinates (x, y, z) on S such that $g_S = \frac{dx^2 + dy^2 + dz^2}{(1 + \frac{\kappa}{4}(x^2 + y^2 + z^2))^2}$, and $\kappa = -1, 0, 1$ corresponds to \mathbb{H}^3 , \mathbb{R}^3 and \mathbb{S}^3 , respectively.
- Special solutions:
 - $\Omega(t) = \text{const}$ and $\kappa = 1$: **Einstein's static Universe**; homogeneous dust with positive Λ .
 - $\Omega(t) = \alpha \cosh(\frac{t}{\alpha})$ and $\kappa = 1$: **deSitter Universe**; vacuum solution with $\Lambda = \frac{3}{\alpha^2}$; **spacetime** of constant curvature.
 - $\Omega(t)$ satisfies $\Omega(\Omega'^2 + \kappa) = \frac{M}{4\pi} + \frac{1}{3}\Lambda\Omega^3$, with $M = \text{const} > 0$, $\Lambda = \text{const}$; **Friedman-Lemaître Universe** with Λ of arbitrary sign, filled with **dust** of energy density $\mu = \frac{M}{\frac{4}{3}\pi\Omega^3}$, comoving with 4-velocity $u = -dt$. For the Ricci scalar being positive one needs $\Lambda > -\mu/4$.
- **Note** that **all FLRW metrics are conformally flat!**

The scheme of Penrose's CCC is as follows:¹

¹See: **P. Tod** (2015), 'The equations of Conformal Cyclic Cosmology', *Gen. Rel. Grav.* **47**, <https://doi.org/10.1007/s10714-015-1859-7>, for details.

The scheme of **Penrose's CCC** is as follows:¹

¹See: **P. Tod** (2015), 'The equations of Conformal Cyclic Cosmology', *Gen. Rel. Grav.* **47**, <https://doi.org/10.1007/s10714-015-1859-7>, for details.

The scheme of **Penrose's CCC** is as follows:¹

¹See: **P. Tod** (2015), 'The equations of Conformal Cyclic Cosmology', *Gen. Rel. Grav.* **47**, <https://doi.org/10.1007/s10714-015-1859-7>, for details.

- The Universe consists of **eons**, each being a **time oriented** spacetime, whose **conformal compactifications** have **spacelike** \mathcal{I} . The **Weyl tensor** of the metric on each \mathcal{I} is **zero**.
- Eons are ordered, and the **conformal compactifications** of consecutive eons, say $\#i$ and $\#(i+1)$, are **glued together** along \mathcal{I}^+ of the eon $\#i$, and \mathcal{I}^- of the eon $\#(i+1)$.
- The **vicinity of the matching surface** (the **wound**) of eons $\#i$ and $\#(i+1)$, this region Penrose calls **bandaged region** for the two eons, is equipped with the following **three metrics**, which are **conformally flat** at the wound:
 - a Lorentzian metric g which is regular everywhere,
 - a Lorentzian metric \check{g} , which represents the physical metric of the eon $\#(i+1)$, and which is singular at the wound,
 - a Lorentzian metric \hat{g} , which represents the physical metric of the eon $\#i$, and which expands at the wound.

- The Universe consists of **eons**, each being a **time oriented** spacetime, whose **conformal compactifications** have **spacelike** \mathcal{I} . The **Weyl tensor** of the metric on each \mathcal{I} is **zero**.
- Eons are ordered, and the **conformal compactifications** of consecutive eons, say $\#i$ and $\#(i+1)$, are **glued together** along \mathcal{I}^+ of the eon $\#i$, and \mathcal{I}^- of the eon $\#(i+1)$.
- The **vicinity of the matching surface** (the **wound**) of eons $\#i$ and $\#(i+1)$, this region Penrose calls **bandaged region** for the two eons, is equipped with the following **three metrics**, which are **conformally flat** at the wound:
 - a Lorentzian metric g which is regular everywhere,
 - a Lorentzian metric \check{g} , which represents the physical metric of the eon $\#(i+1)$, and which is singular at the wound,
 - a Lorentzian metric \hat{g} , which represents the physical metric of the eon $\#i$, and which expands at the wound.

- The Universe consists of **eons**, each being a **time oriented** spacetime, whose **conformal compactifications** have **spacelike** \mathcal{I} . The **Weyl tensor** of the metric on each \mathcal{I} is **zero**.
- Eons are ordered, and the **conformal compactifications** of consecutive eons, say $\#i$ and $\#(i+1)$, are **glued together** along \mathcal{I}^+ of the eon $\#i$, and \mathcal{I}^- of the eon $\#(i+1)$.
- The **vicinity of the matching surface** (the **wound**) of eons $\#i$ and $\#(i+1)$, this region Penrose calls **bandaged region** for the two eons, is equipped with the following **three metrics**, which are **conformally flat** at the wound:
 - a Lorentzian metric g which is regular everywhere,
 - a Lorentzian metric \check{g} , which represents the physical metric of the eon $\#(i+1)$, and which is singular at the wound,
 - a Lorentzian metric \hat{g} , which represents the physical metric of the eon $\#i$, and which expands at the wound.

- The Universe consists of **eons**, each being a **time oriented** spacetime, whose **conformal compactifications** have **spacelike** \mathcal{I} . The **Weyl tensor** of the metric on each \mathcal{I} is **zero**.
- Eons are ordered, and the **conformal compactifications** of consecutive eons, say $\#i$ and $\#(i+1)$, are **glued together** along \mathcal{I}^+ of the eon $\#i$, and \mathcal{I}^- of the eon $\#(i+1)$.
- The **vicinity of the matching surface** (the **wound**) of eons $\#i$ and $\#(i+1)$, this region Penrose calls **bandaged region** for the two eons, is equipped with the following **three metrics**, which are **conformally flat** at the wound:
 - a Lorentzian metric g which is regular everywhere,
 - a Lorentzian metric \check{g} , which represents the physical metric of the eon $\#(i+1)$, and which is singular at the wound,
 - a Lorentzian metric \hat{g} , which represents the physical metric of the eon $\#i$, and which expands at the wound.

- The Universe consists of **eons**, each being a **time oriented** spacetime, whose **conformal compactifications** have **spacelike** \mathcal{I} . The **Weyl tensor** of the metric on each \mathcal{I} is **zero**.
- Eons are ordered, and the **conformal compactifications** of consecutive eons, say $\#i$ and $\#(i+1)$, are **glued together** along \mathcal{I}^+ of the eon $\#i$, and \mathcal{I}^- of the eon $\#(i+1)$.
- The **vicinity of the matching surface** (the **wound**) of eons $\#i$ and $\#(i+1)$, this region Penrose calls **bandaged region** for the two eons, is equipped with the following **three metrics**, which are **conformally flat** at the wound:
 - a Lorentzian metric g which is regular everywhere,
 - a Lorentzian metric \check{g} , which represents the physical metric of the eon $\#(i+1)$, and which is singular at the wound,
 - a Lorentzian metric \hat{g} , which represents the physical metric of the eon $\#i$, and which expands at the wound.

- The Universe consists of **eons**, each being a **time oriented** spacetime, whose **conformal compactifications** have **spacelike** \mathcal{I} . The **Weyl tensor** of the metric on each \mathcal{I} is **zero**.
- Eons are ordered, and the **conformal compactifications** of consecutive eons, say $\#i$ and $\#(i+1)$, are **glued together** along \mathcal{I}^+ of the eon $\#i$, and \mathcal{I}^- of the eon $\#(i+1)$.
- The **vicinity of the matching surface** (the **wound**) of eons $\#i$ and $\#(i+1)$, this region Penrose calls **bandaged region** for the two eons, is equipped with the following **three metrics**, which are **conformally flat** at the wound:
 - a Lorentzian metric g which is regular everywhere,
 - a Lorentzian metric \check{g} , which represents the physical metric of the eon $\#(i+1)$, and which is singular at the wound,
 - a Lorentzian metric \hat{g} , which represents the physical metric of the eon $\#i$, and which expands at the wound.

- The Universe consists of **eons**, each being a **time oriented** spacetime, whose **conformal compactifications** have **spacelike** \mathcal{I} . The **Weyl tensor** of the metric on each \mathcal{I} is **zero**.
- Eons are ordered, and the **conformal compactifications** of consecutive eons, say $\#i$ and $\#(i+1)$, are **glued together** along \mathcal{I}^+ of the eon $\#i$, and \mathcal{I}^- of the eon $\#(i+1)$.
- The **vicinity of the matching surface** (the **wound**) of eons $\#i$ and $\#(i+1)$, this region Penrose calls **bandaged region** for the two eons, is equipped with the following **three metrics**, which are **conformally flat** at the wound:
 - a Lorentzian metric g which is regular everywhere,
 - a Lorentzian metric \check{g} , which represents the physical metric of the eon $\#(i+1)$, and which is singular at the wound,
 - a Lorentzian metric \hat{g} , which represents the physical metric of the eon $\#i$, and which expands at the wound.

- The Universe consists of **eons**, each being a **time oriented** spacetime, whose **conformal compactifications** have **spacelike** \mathcal{I} . The **Weyl tensor** of the metric on each \mathcal{I} is **zero**.
- Eons are ordered, and the **conformal compactifications** of consecutive eons, say $\#i$ and $\#(i+1)$, are **glued together** along \mathcal{I}^+ of the eon $\#i$, and \mathcal{I}^- of the eon $\#(i+1)$.
- The **vicinity of the matching surface** (the **wound**) of eons $\#i$ and $\#(i+1)$, this region Penrose calls **bandaged region** for the two eons, is equipped with the following **three metrics**, which are **conformally flat** at the wound:
 - a Lorentzian metric g which is regular everywhere,
 - a Lorentzian metric \check{g} , which represents the physical metric of the eon $\#(i+1)$, and which is singular at the wound,
 - a Lorentzian metric \hat{g} , which represents the physical metric of the eon $\#i$, and which expands at the wound.

- The Universe consists of **eons**, each being a **time oriented** spacetime, whose **conformal compactifications** have **spacelike** \mathcal{I} . The **Weyl tensor** of the metric on each \mathcal{I} is **zero**.
- Eons are ordered, and the **conformal compactifications** of consecutive eons, say $\#i$ and $\#(i+1)$, are **glued together** along \mathcal{I}^+ of the eon $\#i$, and \mathcal{I}^- of the eon $\#(i+1)$.
- The **vicinity of the matching surface** (the **wound**) of eons $\#i$ and $\#(i+1)$, this region Penrose calls **bandaged region** for the two eons, is equipped with the following **three metrics**, which are **conformally flat** at the wound:
 - a Lorentzian metric g which is regular everywhere,
 - a Lorentzian metric \check{g} , which represents the physical metric of the eon $\#(i+1)$, and which is singular at the wound,
 - a Lorentzian metric \hat{g} , which represents the physical metric of the eon $\#i$, and which expands at the wound.

- The Universe consists of **eons**, each being a **time oriented** spacetime, whose **conformal compactifications** have **spacelike** \mathcal{I} . The **Weyl tensor** of the metric on each \mathcal{I} is **zero**.
- Eons are ordered, and the **conformal compactifications** of consecutive eons, say $\#i$ and $\#(i+1)$, are **glued together** along \mathcal{I}^+ of the eon $\#i$, and \mathcal{I}^- of the eon $\#(i+1)$.
- The **vicinity of the matching surface** (the **wound**) of eons $\#i$ and $\#(i+1)$, this region Penrose calls **bandaged region** for the two eons, is equipped with the following **three metrics**, which are **conformally flat** at the wound:
 - a Lorentzian metric g which is regular everywhere,
 - a Lorentzian metric \check{g} , which represents the physical metric of the eon $\#(i+1)$, and which is singular at the wound,
 - a Lorentzian metric \hat{g} , which represents the physical metric of the eon $\#i$, and which expands at the wound.

- The Universe consists of **eons**, each being a **time oriented** spacetime, whose **conformal compactifications** have **spacelike** \mathcal{I} . The **Weyl tensor** of the metric on each \mathcal{I} is **zero**.
- Eons are ordered, and the **conformal compactifications** of consecutive eons, say $\#i$ and $\#(i + 1)$, are **glued together** along \mathcal{I}^+ of the eon $\#i$, and \mathcal{I}^- of the eon $\#(i + 1)$.
- The **vicinity of the matching surface** (the **wound**) of eons $\#i$ and $\#(i + 1)$, this region Penrose calls **bandaged region** for the two eons, is equipped with the following **three metrics**, which are **conformally flat** at the wound:
 - a Lorentzian metric g which is regular everywhere,
 - a Lorentzian metric \check{g} , which represents the physical metric of the eon $\#(i + 1)$, and which is singular at the wound,
 - a Lorentzian metric \hat{g} , which represents the physical metric of the eon $\#i$, and which expands at the wound.

- The Universe consists of **eons**, each being a **time oriented** spacetime, whose **conformal compactifications** have **spacelike** \mathcal{I} . The **Weyl tensor** of the metric on each \mathcal{I} is **zero**.
- Eons are ordered, and the **conformal compactifications** of consecutive eons, say $\#i$ and $\#(i+1)$, are **glued together** along \mathcal{I}^+ of the eon $\#i$, and \mathcal{I}^- of the eon $\#(i+1)$.
- The **vicinity of the matching surface** (the **wound**) of eons $\#i$ and $\#(i+1)$, this region Penrose calls **bandaged region** for the two eons, is equipped with the following **three metrics**, which are **conformally flat** at the wound:
 - a Lorentzian metric g which is regular everywhere,
 - a Lorentzian metric \check{g} , which represents the physical metric of the eon $\#(i+1)$, and which is singular at the wound,
 - a Lorentzian metric \hat{g} , which represents the physical metric of the eon $\#i$, and which expands at the wound.

- In any bandaged region, the **three metrics** g , \check{g} and \hat{g} , are **conformally related**.
- How to make this relation specific is debatable, but Penrose proposes that
$$\check{g} = \Omega^2 g, \text{ and } \hat{g} = \frac{1}{\Omega^2} g, \text{ with } \Omega \rightarrow 0 \text{ on the wound.}$$
- The metric \check{g} in eon $\#(i+1)$ is a **physical metric there**. Likewise, the metric \hat{g} in eon $\#i$ is a **physical metric there**.
- Of course, the metric \check{g} in eon $\#(i+1)$, and the metric \hat{g} in eon $\#i$, as **physical spacetime metrics, should satisfy Einstein's equations** in $\#(i+1)$ and $\#i$, respectively.

- In any bandaged region, the **three metrics** g , \check{g} and \hat{g} , are **conformally related**.
- How to make this relation specific is debatable, but Penrose proposes that
$$\check{g} = \Omega^2 g, \text{ and } \hat{g} = \frac{1}{\Omega^2} g, \text{ with } \Omega \rightarrow 0 \text{ on the wound.}$$
- The metric \check{g} in eon $\#(i+1)$ is a **physical metric there**. Likewise, the metric \hat{g} in eon $\#i$ is a **physical metric there**.
- Of course, the metric \check{g} in eon $\#(i+1)$, and the metric \hat{g} in eon $\#i$, as **physical spacetime metrics**, should satisfy **Einstein's equations** in $\#(i+1)$ and $\#i$, respectively.

- In any bandaged region, the **three metrics** g , \check{g} and \hat{g} , are **conformally related**.
- How to make this relation specific is debatable, but Penrose proposes that
$$\check{g} = \Omega^2 g, \text{ and } \hat{g} = \frac{1}{\Omega^2} g, \text{ with } \Omega \rightarrow 0 \text{ on the wound.}$$
- The metric \check{g} in eon $\#(i+1)$ is a **physical metric there**. Likewise, the metric \hat{g} in eon $\#i$ is a **physical metric there**.
- Of course, the metric \check{g} in eon $\#(i+1)$, and the metric \hat{g} in eon $\#i$, as **physical spacetime metrics, should satisfy Einstein's equations** in $\#(i+1)$ and $\#i$, respectively.

- In any bandaged region, the **three metrics** g , \check{g} and \hat{g} , are **conformally related**.
- How to make this relation specific is debatable, but Penrose proposes that

$$\check{g} = \Omega^2 g, \text{ and } \hat{g} = \frac{1}{\Omega^2} g, \text{ with } \Omega \rightarrow 0 \text{ on the wound.}$$

- The metric \check{g} in eon $\#(i+1)$ is a **physical metric there**. Likewise, the metric \hat{g} in eon $\#i$ is a **physical metric there**.
- Of course, the metric \check{g} in eon $\#(i+1)$, and the metric \hat{g} in eon $\#i$, as **physical spacetime metrics**, should satisfy **Einstein's equations** in $\#(i+1)$ and $\#i$, respectively.

- In any bandaged region, the **three metrics** g , \check{g} and \hat{g} , are **conformally related**.
- How to make this relation specific is debatable, but Penrose proposes that
$$\check{g} = \Omega^2 g, \text{ and } \hat{g} = \frac{1}{\Omega^2} g, \text{ with } \Omega \rightarrow 0 \text{ on the wound.}$$
- The metric \check{g} in eon $\#(i+1)$ is a **physical metric there**. Likewise, the metric \hat{g} in eon $\#i$ is a **physical metric there**.
- Of course, the metric \check{g} in eon $\#(i+1)$, and the metric \hat{g} in eon $\#i$, as **physical spacetime metrics**, should satisfy **Einstein's equations** in $\#(i+1)$ and $\#i$, respectively.

- In any bandaged region, the **three metrics** g , \check{g} and \hat{g} , are **conformally related**.
- How to make this relation specific is debatable, but Penrose proposes that
$$\check{g} = \Omega^2 g, \text{ and } \hat{g} = \frac{1}{\Omega^2} g, \text{ with } \Omega \rightarrow 0 \text{ on the wound.}$$
- The metric \check{g} in eon $\#(i+1)$ is a **physical metric there**. Likewise, the metric \hat{g} in eon $\#i$ is a **physical metric there**.
- Of course, the metric \check{g} in eon $\#(i+1)$, and the metric \hat{g} in eon $\#i$, as **physical spacetime metrics**, should satisfy **Einstein's equations** in $\#(i+1)$ and $\#i$, respectively.

- In any bandaged region, the **three metrics** g , \check{g} and \hat{g} , are **conformally related**.
- How to make this relation specific is debatable, but Penrose proposes that
$$\check{g} = \Omega^2 g, \text{ and } \hat{g} = \frac{1}{\Omega^2} g, \text{ with } \Omega \rightarrow 0 \text{ on the wound.}$$
- The metric \check{g} in eon $\#(i+1)$ is a **physical metric there**. Likewise, the metric \hat{g} in eon $\#i$ is a **physical metric there**.
- Of course, the metric \check{g} in eon $\#(i+1)$, and the metric \hat{g} in eon $\#i$, as **physical spacetime metrics**, should satisfy **Einstein's equations** in $\#(i+1)$ and $\#i$, respectively.

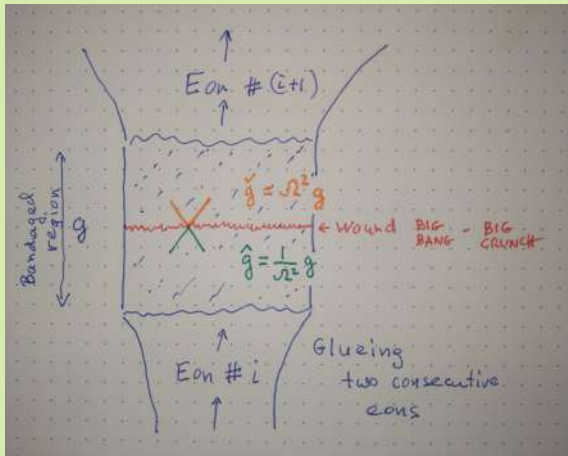
- In any bandaged region, the **three metrics** g , \check{g} and \hat{g} , are **conformally related**.
- How to make this relation specific is debatable, but Penrose proposes that
$$\check{g} = \Omega^2 g, \text{ and } \hat{g} = \frac{1}{\Omega^2} g, \text{ with } \Omega \rightarrow 0 \text{ on the wound.}$$
- The metric \check{g} in eon $\#(i+1)$ is a **physical metric there**. Likewise, the metric \hat{g} in eon $\#i$ is a **physical metric there**.
- Of course, the metric \check{g} in eon $\#(i+1)$, and the metric \hat{g} in eon $\#i$, as **physical spacetime metrics, should satisfy Einstein's equations** in $\#(i+1)$ and $\#i$, respectively.

- In any bandaged region, the **three metrics** g , \check{g} and \hat{g} , are **conformally related**.
- How to make this relation specific is debatable, but Penrose proposes that
$$\check{g} = \Omega^2 g, \text{ and } \hat{g} = \frac{1}{\Omega^2} g, \text{ with } \Omega \rightarrow 0 \text{ on the wound.}$$
- The metric \check{g} in eon $\#(i+1)$ is a **physical metric there**. Likewise, the metric \hat{g} in eon $\#i$ is a **physical metric there**.
- Of course, the metric \check{g} in eon $\#(i+1)$, and the metric \hat{g} in eon $\#i$, as **physical spacetime metrics, should satisfy Einstein's equations** in $\#(i+1)$ and $\#i$, respectively.

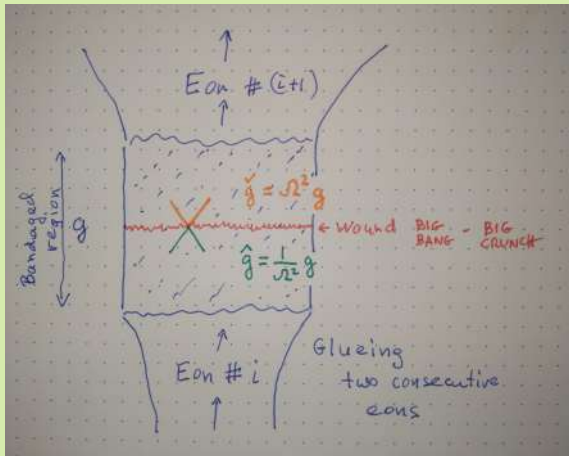
- In any bandaged region, the **three metrics** g , \check{g} and \hat{g} , are **conformally related**.
- How to make this relation specific is debatable, but Penrose proposes that
$$\check{g} = \Omega^2 g, \text{ and } \hat{g} = \frac{1}{\Omega^2} g, \text{ with } \Omega \rightarrow 0 \text{ on the wound.}$$
- The metric \check{g} in eon $\#(i+1)$ is a **physical metric there**. Likewise, the metric \hat{g} in eon $\#i$ is a **physical metric there**.
- Of course, the metric \check{g} in eon $\#(i+1)$, and the metric \hat{g} in eon $\#i$, as **physical spacetime metrics**, should satisfy **Einstein's equations** in $\#(i+1)$ and $\#i$, respectively.

- In any bandaged region, the **three metrics** g , \check{g} and \hat{g} , are **conformally related**.
- How to make this relation specific is debatable, but Penrose proposes that
$$\check{g} = \Omega^2 g, \text{ and } \hat{g} = \frac{1}{\Omega^2} g, \text{ with } \Omega \rightarrow 0 \text{ on the wound.}$$
- The metric \check{g} in eon $\#(i+1)$ is a **physical metric there**. Likewise, the metric \hat{g} in eon $\#i$ is a **physical metric there**.
- Of course, the metric \check{g} in eon $\#(i+1)$, and the metric \hat{g} in eon $\#i$, as **physical spacetime metrics, should satisfy Einstein's equations** in $\#(i+1)$ and $\#i$, respectively.

Penrose's Conformal Cyclic Cosmology



Penrose's Conformal Cyclic Cosmology



- **Question:** How to make a model of Penrose's bandaged region?
- One needs a function Ω , **vanishing on some spacelike hypersurface**, such that if $\check{g} = \Omega^2 g$ satisfies **Einstein equations** with some physically reasonable energy momentum tensor, then $\hat{g} = \frac{1}{\Omega^2} g$ **also satisfies Einstein equations** with possibly different, but still physically reasonable energy momentum tensor.
- Similar, but seems to me simpler, than a problem of **Brinkman**, who in 1925 asked a question '**when in a conformal class of metrics can be two different Einstein metrics?**'. Brinkman found all such metrics.
- Why not to start with **conformally flat situation** (reasonable, because compatible with the **cosmological principle/FLRW paradigm**), and (various) **perfect fluids?**

- **Question:** How to make a model of Penrose's bandaged region?
- One needs a function Ω , **vanishing on some spacelike hypersurface**, such that if $\check{g} = \Omega^2 g$ satisfies **Einstein equations** with some physically reasonable energy momentum tensor, then $\hat{g} = \frac{1}{\Omega^2} g$ **also satisfies Einstein equations** with possibly different, but still physically reasonable energy momentum tensor.
- Similar, but seems to me simpler, than a problem of **Brinkman**, who in 1925 asked a question '**when in a conformal class of metrics can be two different Einstein metrics?**'. Brinkman found all such metrics.
- Why not to start with **conformally flat situation** (reasonable, because compatible with the **cosmological principle/FLRW paradigm**), and (various) **perfect fluids?**

- **Question:** How to make a model of Penrose's bandaged region?
- One needs a function Ω , **vanishing on some spacelike hypersurface**, such that if $\check{g} = \Omega^2 g$ satisfies **Einstein equations** with some physically reasonable energy momentum tensor, then $\hat{g} = \frac{1}{\Omega^2} g$ **also satisfies Einstein equations** with possibly different, but still physically reasonable energy momentum tensor.
- Similar, but seems to me simpler, than a problem of **Brinkman**, who in 1925 asked a question '**when in a conformal class of metrics can be two different Einstein metrics?**'. Brinkman found all such metrics.
- Why not to start with **conformally flat situation** (reasonable, because compatible with the **cosmological principle/FLRW paradigm**), and (various) **perfect fluids?**

- **Question:** How to make a model of Penrose's bandaged region?
- One needs a function Ω , **vanishing on some spacelike hypersurface**, such that if $\check{g} = \Omega^2 g$ satisfies **Einstein equations** with some physically reasonable energy momentum tensor, then $\hat{g} = \frac{1}{\Omega^2} g$ **also satisfies Einstein equations** with possibly different, but still physically reasonable energy momentum tensor.
- Similar, but seems to me simpler, than a problem of **Brinkman**, who in 1925 asked a question '**when in a conformal class of metrics can be two different Einstein metrics?**'. Brinkman found all such metrics.
- Why not to start with **conformally flat situation** (reasonable, because compatible with the **cosmological principle/FLRW paradigm**), and (various) **perfect fluids?**

- **Question:** How to make a model of Penrose's bandaged region?
- One needs a function Ω , **vanishing on some spacelike hypersurface**, such that **if $\check{g} = \Omega^2 g$ satisfies Einstein equations** with some physically reasonable energy momentum tensor, **then $\hat{g} = \frac{1}{\Omega^2} g$ also satisfies Einstein equations** with possibly different, but still physically reasonable energy momentum tensor.
- Similar, but seems to me simpler, than a problem of **Brinkman**, who in 1925 asked a question '**when in a conformal class of metrics can be two different Einstein metrics?**'. Brinkman found all such metrics.
- Why not to start with **conformally flat situation** (reasonable, because compatible with the **cosmological principle/FLRW paradigm**), and (various) **perfect fluids?**

- **Question:** How to make a model of Penrose's bandaged region?
- One needs a function Ω , **vanishing on some spacelike hypersurface**, such that if $\check{g} = \Omega^2 g$ satisfies **Einstein equations** with some physically reasonable energy momentum tensor, **then $\hat{g} = \frac{1}{\Omega^2} g$ also satisfies Einstein equations** with possibly different, but still physically reasonable energy momentum tensor.
- Similar, but seems to me simpler, than a problem of **Brinkman**, who in 1925 asked a question '**when in a conformal class of metrics can be two different Einstein metrics?**'. Brinkman found all such metrics.
- Why not to start with **conformally flat situation** (reasonable, because compatible with the **cosmological principle/FLRW paradigm**), and (various) **perfect fluids?**

- **Question:** How to make a model of Penrose's bandaged region?
- One needs a function Ω , **vanishing on some spacelike hypersurface**, such that **if $\check{g} = \Omega^2 g$ satisfies Einstein equations** with some physically reasonable energy momentum tensor, **then $\hat{g} = \frac{1}{\Omega^2} g$ also satisfies Einstein equations** with possibly different, but still physically reasonable energy momentum tensor.
- Similar, but seems to me simpler, than a problem of **Brinkman**, who in 1925 asked a question '**when in a conformal class of metrics can be two different Einstein metrics?**'. Brinkman found all such metrics.
- Why not to start with **conformally flat situation** (reasonable, because compatible with the **cosmological principle/FLRW paradigm**), and (various) **perfect fluids?**

- **Question:** How to make a model of Penrose's bandaged region?
- One needs a function Ω , **vanishing on some spacelike hypersurface**, such that if $\check{g} = \Omega^2 g$ satisfies **Einstein equations** with some physically reasonable energy momentum tensor, **then $\hat{g} = \frac{1}{\Omega^2} g$ also satisfies Einstein equations** with possibly different, but still physically reasonable energy momentum tensor.
- Similar, but seems to me simpler, than a problem of **Brinkman**, who in 1925 asked a question '**when in a conformal class of metrics can be two different Einstein metrics?**'. Brinkman found all such metrics.
- Why not to start with **conformally flat situation** (reasonable, because compatible with the **cosmological principle/FLRW paradigm**), and (various) **perfect fluids?**

- **Question:** How to make a model of Penrose's bandaged region?
- One needs a function Ω , **vanishing on some spacelike hypersurface**, such that if $\check{g} = \Omega^2 g$ satisfies **Einstein equations** with some physically reasonable energy momentum tensor, then $\hat{g} = \frac{1}{\Omega^2} g$ **also satisfies Einstein equations** with possibly different, but still physically reasonable energy momentum tensor.
- Similar, but seems to me simpler, than a problem of **Brinkman**, who in 1925 asked a question '**when in a conformal class of metrics can be two different Einstein metrics?**'. Brinkman found all such metrics.
- Why not to start with **conformally flat situation** (reasonable, because compatible with the **cosmological principle/FLRW paradigm**), and (various) **perfect fluids?**

- **Question:** How to make a model of Penrose's bandaged region?
- One needs a function Ω , **vanishing on some spacelike hypersurface**, such that if $\check{g} = \Omega^2 g$ satisfies **Einstein equations** with some physically reasonable energy momentum tensor, then $\hat{g} = \frac{1}{\Omega^2} g$ **also satisfies Einstein equations** with possibly different, but still physically reasonable energy momentum tensor.
- Similar, but seems to me simpler, than a problem of **Brinkman**, who in 1925 asked a question '**when in a conformal class of metrics can be two different Einstein metrics?**'. Brinkman found all such metrics.
- Why not to start with **conformally flat situation** (reasonable, because compatible with the **cosmological principle/FLRW paradigm**), and (various) **perfect fluids?**

- **Question:** How to make a model of Penrose's bandaged region?
- One needs a function Ω , **vanishing on some spacelike hypersurface**, such that if $\check{g} = \Omega^2 g$ satisfies **Einstein equations** with some physically reasonable energy momentum tensor, then $\hat{g} = \frac{1}{\Omega^2} g$ **also satisfies Einstein equations** with possibly different, but still physically reasonable energy momentum tensor.
- Similar, but seems to me simpler, than a problem of **Brinkman**, who in 1925 asked a question '**when in a conformal class of metrics can be two different Einstein metrics?**'. Brinkman found all such metrics.
- Why not to start with **conformally flat situation** (reasonable, because compatible with the **cosmological principle/FLRW paradigm**), and (various) **perfect fluids?**

- From now on I restrict myself to FLRW metrics with $\kappa = 1$,
$$g = -dt^2 + \Omega^2(t)r_0^2 \left(d\chi^2 + \sin^2 \chi (d\theta^2 + \sin^2 \theta d\phi^2) \right).$$

- It is convenient to introduce a **conformal time** $\eta = \int \frac{dt}{a(t)}$ so that the FLRW metric looks

$$g = \Omega^2(\eta) \left(-d\eta^2 + r_0^2 (d\chi^2 + \sin^2 \chi (d\theta^2 + \sin^2 \theta d\phi^2)) \right),$$

i.e. $g = \Omega^2(\eta)g_{Einst}$.

- This parametrization is very convenient since taking $u = -\Omega(\eta)d\eta$, the most general FLRW metric g satisfying **Einstein's equations**

$$Ric - \frac{1}{2}Rg = (\mu + p)u \otimes u + pg$$

with **polytropic equation of state** $p = w\mu$, $w = const$, is given by

$$\Omega(\eta) = \Omega_0 \left(\sin^2 \frac{(1+3w)\eta}{2r_0} \right)^{\frac{1}{1+3w}} \text{ if } w \neq -\frac{1}{3},$$

and

$$\Omega(\eta) = \Omega_0 \exp(b\eta) \text{ if } w = -\frac{1}{3}.$$

- From now on I restrict myself to FLRW metrics with $\kappa = 1$,

$$g = -dt^2 + \Omega^2(t)r_0^2 \left(d\chi^2 + \sin^2 \chi (d\theta^2 + \sin^2 \theta d\phi^2) \right).$$

- It is convenient to introduce a **conformal time** $\eta = \int \frac{dt}{a(t)}$ so that the FLRW metric looks

$$g = \Omega^2(\eta) \left(-d\eta^2 + r_0^2 (d\chi^2 + \sin^2 \chi (d\theta^2 + \sin^2 \theta d\phi^2)) \right),$$

i.e. $g = \Omega^2(\eta)g_{Einst}$.

- This parametrization is very convenient since taking $u = -\Omega(\eta)d\eta$, the most general FLRW metric g satisfying **Einstein's equations**

$$Ric - \frac{1}{2}Rg = (\mu + p)u \otimes u + pg$$

with **polytropic equation of state** $p = w\mu$, $w = const$, is given by

$$\Omega(\eta) = \Omega_0 \left(\sin^2 \frac{(1+3w)\eta}{2r_0} \right)^{\frac{1}{1+3w}} \text{ if } w \neq -\frac{1}{3},$$

and

$$\Omega(\eta) = \Omega_0 \exp(b\eta) \text{ if } w = -\frac{1}{3}.$$

- From now on I restrict myself to FLRW metrics with $\kappa = 1$,

$$g = -dt^2 + \Omega^2(t)r_0^2 \left(d\chi^2 + \sin^2 \chi (d\theta^2 + \sin^2 \theta d\phi^2) \right).$$

- It is convenient to introduce a **conformal time** $\eta = \int \frac{dt}{a(t)}$ so that the FLRW metric looks

$$g = \Omega^2(\eta) \left(-d\eta^2 + r_0^2 (d\chi^2 + \sin^2 \chi (d\theta^2 + \sin^2 \theta d\phi^2)) \right),$$

i.e. $g = \Omega^2(\eta)g_{Einst}$.

- This parametrization is very convenient since taking $u = -\Omega(\eta)d\eta$, the most general FLRW metric g satisfying **Einstein's equations**

$$Ric - \frac{1}{2}Rg = (\mu + p)u \otimes u + pg$$

with **polytropic equation of state** $p = w\mu$, $w = const$, is given by

$$\Omega(\eta) = \Omega_0 \left(\sin^2 \frac{(1+3w)\eta}{2r_0} \right)^{\frac{1}{1+3w}} \text{ if } w \neq -\frac{1}{3},$$

and

$$\Omega(\eta) = \Omega_0 \exp(b\eta) \text{ if } w = -\frac{1}{3}.$$

- From now on I restrict myself to FLRW metrics with $\kappa = 1$,

$$g = -dt^2 + \Omega^2(t)r_0^2 \left(d\chi^2 + \sin^2 \chi (d\theta^2 + \sin^2 \theta d\phi^2) \right).$$

- It is convenient to introduce a **conformal time** $\eta = \int \frac{dt}{a(t)}$ so that the FLRW metric looks

$$g = \Omega^2(\eta) \left(-d\eta^2 + r_0^2 (d\chi^2 + \sin^2 \chi (d\theta^2 + \sin^2 \theta d\phi^2)) \right),$$

i.e. $g = \Omega^2(\eta)g_{Einst}$.

- This parametrization is very convenient since taking $u = -\Omega(\eta)d\eta$, the most general FLRW metric g satisfying **Einstein's equations**

$$Ric - \frac{1}{2}Rg = (\mu + p)u \otimes u + pg$$

with **polytropic equation of state** $p = w\mu$, $w = const$, is given by

$$\Omega(\eta) = \Omega_0 \left(\sin^2 \frac{(1+3w)\eta}{2r_0} \right)^{\frac{1}{1+3w}} \text{ if } w \neq -\frac{1}{3},$$

and

$$\Omega(\eta) = \Omega_0 \exp(b\eta) \text{ if } w = -\frac{1}{3}.$$

- From now on I restrict myself to FLRW metrics with $\kappa = 1$,

$$g = -dt^2 + \Omega^2(t)r_0^2 \left(d\chi^2 + \sin^2 \chi (d\theta^2 + \sin^2 \theta d\phi^2) \right).$$

- It is convenient to introduce a **conformal time** $\eta = \int \frac{dt}{a(t)}$ so that the FLRW metric looks

$$g = \Omega^2(\eta) \left(-d\eta^2 + r_0^2 (d\chi^2 + \sin^2 \chi (d\theta^2 + \sin^2 \theta d\phi^2)) \right),$$

i.e. $g = \Omega^2(\eta)g_{Einst}$.

- This parametrization is very convenient since taking $u = -\Omega(\eta)d\eta$, the most general FLRW metric g satisfying **Einstein's equations**

$$Ric - \frac{1}{2}Rg = (\mu + p)u \otimes u + pg$$

with **polytropic equation of state** $p = w\mu$, $w = const$, is given by

$$\Omega(\eta) = \Omega_0 \left(\sin^2 \frac{(1+3w)\eta}{2r_0} \right)^{\frac{1}{1+3w}} \text{ if } w \neq -\frac{1}{3},$$

and

$$\Omega(\eta) = \Omega_0 \exp(b\eta) \text{ if } w = -\frac{1}{3}.$$

- From now on I restrict myself to FLRW metrics with $\kappa = 1$,

$$g = -dt^2 + \Omega^2(t)r_0^2 \left(d\chi^2 + \sin^2 \chi (d\theta^2 + \sin^2 \theta d\phi^2) \right).$$

- It is convenient to introduce a **conformal time** $\eta = \int \frac{dt}{a(t)}$ so that the FLRW metric looks

$$g = \Omega^2(\eta) \left(-d\eta^2 + r_0^2 (d\chi^2 + \sin^2 \chi (d\theta^2 + \sin^2 \theta d\phi^2)) \right),$$

i.e. $g = \Omega^2(\eta)g_{Einst}$.

- This parametrization is very convenient since taking $u = -\Omega(\eta)d\eta$, the most general FLRW metric g satisfying **Einstein's equations**

$$Ric - \frac{1}{2}Rg = (\mu + p)u \otimes u + pg$$

with **polytropic equation of state** $p = w\mu$, $w = const$, is given by

$$\Omega(\eta) = \Omega_0 \left(\sin^2 \frac{(1+3w)\eta}{2r_0} \right)^{\frac{1}{1+3w}} \text{ if } w \neq -\frac{1}{3},$$

and

$$\Omega(\eta) = \Omega_0 \exp(b\eta) \text{ if } w = -\frac{1}{3}.$$

- From now on I restrict myself to FLRW metrics with $\kappa = 1$,

$$g = -dt^2 + \Omega^2(t)r_0^2 \left(d\chi^2 + \sin^2 \chi (d\theta^2 + \sin^2 \theta d\phi^2) \right).$$

- It is convenient to introduce a **conformal time** $\eta = \int \frac{dt}{a(t)}$ so that the FLRW metric looks

$$g = \Omega^2(\eta) \left(-d\eta^2 + r_0^2 (d\chi^2 + \sin^2 \chi (d\theta^2 + \sin^2 \theta d\phi^2)) \right),$$

i.e. $g = \Omega^2(\eta)g_{Einst}$.

- This parametrization is very convenient since taking $u = -\Omega(\eta)d\eta$, the most general FLRW metric g satisfying **Einstein's equations**

$$Ric - \frac{1}{2}Rg = (\mu + p)u \otimes u + pg$$

with **polytropic equation of state** $p = w\mu$, $w = const$, is given by

$$\Omega(\eta) = \Omega_0 \left(\sin^2 \frac{(1+3w)\eta}{2r_0} \right)^{\frac{1}{1+3w}} \text{ if } w \neq -\frac{1}{3},$$

and

$$\Omega(\eta) = \Omega_0 \exp(b\eta) \text{ if } w = -\frac{1}{3}.$$

- From now on I restrict myself to FLRW metrics with $\kappa = 1$,

$$g = -dt^2 + \Omega^2(t)r_0^2 \left(d\chi^2 + \sin^2 \chi (d\theta^2 + \sin^2 \theta d\phi^2) \right).$$

- It is convenient to introduce a **conformal time** $\eta = \int \frac{dt}{a(t)}$ so that the FLRW metric looks

$$g = \Omega^2(\eta) \left(-d\eta^2 + r_0^2 (d\chi^2 + \sin^2 \chi (d\theta^2 + \sin^2 \theta d\phi^2)) \right),$$

i.e. $g = \Omega^2(\eta)g_{Einst}$.

- This parametrization is very convenient since taking $u = -\Omega(\eta)d\eta$, the most general FLRW metric g satisfying **Einstein's equations**

$$Ric - \frac{1}{2}Rg = (\mu + p)u \otimes u + pg$$

with **polytropic equation of state** $p = w\mu$, $w = const$, is given by

$$\Omega(\eta) = \Omega_0 \left(\sin^2 \frac{(1+3w)\eta}{2r_0} \right)^{\frac{1}{1+3w}} \text{ if } w \neq -\frac{1}{3},$$

and

$$\Omega(\eta) = \Omega_0 \exp(b\eta) \text{ if } w = -\frac{1}{3}.$$

- From now on I restrict myself to FLRW metrics with $\kappa = 1$,

$$g = -dt^2 + \Omega^2(t)r_0^2 \left(d\chi^2 + \sin^2 \chi (d\theta^2 + \sin^2 \theta d\phi^2) \right).$$

- It is convenient to introduce a **conformal time** $\eta = \int \frac{dt}{a(t)}$ so that the FLRW metric looks

$$g = \Omega^2(\eta) \left(-d\eta^2 + r_0^2 (d\chi^2 + \sin^2 \chi (d\theta^2 + \sin^2 \theta d\phi^2)) \right),$$

i.e. $g = \Omega^2(\eta)g_{Einst}$.

- This parametrization is very convenient since taking $u = -\Omega(\eta)d\eta$, the most general FLRW metric g satisfying **Einstein's equations**

$$Ric - \frac{1}{2}Rg = (\mu + p)u \otimes u + pg$$

with **polytropic equation of state** $p = w\mu$, $w = const$, is given by

$$\Omega(\eta) = \Omega_0 \left(\sin^2 \frac{(1+3w)\eta}{2r_0} \right)^{\frac{1}{1+3w}} \text{ if } w \neq -\frac{1}{3},$$

and

$$\Omega(\eta) = \Omega_0 \exp(b\eta) \text{ if } w = -\frac{1}{3}.$$

- From now on I restrict myself to FLRW metrics with $\kappa = 1$,

$$g = -dt^2 + \Omega^2(t)r_0^2 \left(d\chi^2 + \sin^2 \chi (d\theta^2 + \sin^2 \theta d\phi^2) \right).$$

- It is convenient to introduce a **conformal time** $\eta = \int \frac{dt}{a(t)}$ so that the FLRW metric looks

$$g = \Omega^2(\eta) \left(-d\eta^2 + r_0^2 (d\chi^2 + \sin^2 \chi (d\theta^2 + \sin^2 \theta d\phi^2)) \right),$$

i.e. $g = \Omega^2(\eta)g_{Einst}$.

- This parametrization is very convenient since taking $u = -\Omega(\eta)d\eta$, the most general FLRW metric g satisfying **Einstein's equations**

$$Ric - \frac{1}{2}Rg = (\mu + p)u \otimes u + pg$$

with **polytropic equation of state** $p = w\mu$, $w = const$, is given by

$$\Omega(\eta) = \Omega_0 \left(\sin^2 \frac{(1+3w)\eta}{2r_0} \right)^{\frac{1}{1+3w}} \text{ if } w \neq -\frac{1}{3},$$

and

$$\Omega(\eta) = \Omega_0 \exp(b\eta) \text{ if } w = -\frac{1}{3}.$$

- From now on I restrict myself to FLRW metrics with $\kappa = 1$,

$$g = -dt^2 + \Omega^2(t)r_0^2 \left(d\chi^2 + \sin^2 \chi (d\theta^2 + \sin^2 \theta d\phi^2) \right).$$

- It is convenient to introduce a **conformal time** $\eta = \int \frac{dt}{a(t)}$ so that the FLRW metric looks

$$g = \Omega^2(\eta) \left(-d\eta^2 + r_0^2 (d\chi^2 + \sin^2 \chi (d\theta^2 + \sin^2 \theta d\phi^2)) \right),$$

i.e. $g = \Omega^2(\eta)g_{Einst}$.

- This parametrization is very convenient since taking $u = -\Omega(\eta)d\eta$, the most general FLRW metric g satisfying **Einstein's equations**

$$Ric - \frac{1}{2}Rg = (\mu + p)u \otimes u + pg$$

with **polytropic equation of state** $p = w\mu$, $w = const$, is given by

$$\Omega(\eta) = \Omega_0 \left(\sin^2 \frac{(1+3w)\eta}{2r_0} \right)^{\frac{1}{1+3w}} \text{ if } w \neq -\frac{1}{3},$$

and

$$\Omega(\eta) = \Omega_0 \exp(b\eta) \text{ if } w = -\frac{1}{3}.$$

Theorem

If $\Omega = \Omega(\eta)$ is such that $\check{g} = \Omega^2 g_{Einst}$ satisfies Einstein's equations, with $\Lambda = 0$, and with the energy momentum tensor \check{T} of a perfect fluid, whose pressure \check{p} is proportional to the energy density $\check{\mu}$, via $\check{p} = \check{w}\check{\mu}$, $\check{w} = const$, then

$\hat{g} = \frac{1}{\Omega^2} g_{Einst}$ satisfies Einstein's equations, with $\Lambda = 0$, and with the energy momentum tensor \hat{T} of a perfect fluid, whose pressure \hat{p} and the energy density $\hat{\mu}$ are related by $\hat{p} = \hat{w}\hat{\mu}$ with

$$\hat{w} = -\frac{1}{3}(2 + 3\check{w}).$$

The Ricci scalar of the metric \check{g} is

$$R = \frac{3(1-3\check{w})}{\Omega_0^2 r_0^2 \left(\sin^6 \frac{(1+3\check{w})\eta}{2r_0} \right)^{\frac{1+w}{1+3w}}} \text{ if } \check{w} \neq -1/3 \text{ and } R = \frac{6(1+b^2 r_0^2)}{\Omega_0^2 r_0^2 \exp(2b\eta)} \text{ if } \check{w} = -1/3,$$

so it is **positive** if $-1 \leq \check{w} < 1/3$ (recall the energy conditions $-1 \leq \check{w} \leq 1$).

Theorem

If $\Omega = \Omega(\eta)$ is such that $\check{g} = \Omega^2 g_{Einst}$ satisfies Einstein's equations, with $\Lambda = 0$, and with the energy momentum tensor \check{T} of a perfect fluid, whose pressure \check{p} is proportional to the energy density $\check{\mu}$, via $\check{p} = \check{w}\check{\mu}$, $\check{w} = const$, then

$\hat{g} = \frac{1}{\Omega^2} g_{Einst}$ satisfies Einstein's equations, with $\Lambda = 0$, and with the energy momentum tensor \hat{T} of a perfect fluid, whose pressure \hat{p} and the energy density $\hat{\mu}$ are related by $\hat{p} = \hat{w}\hat{\mu}$ with

$$\hat{w} = -\frac{1}{3}(2 + 3\check{w}).$$

The Ricci scalar of the metric \check{g} is

$$R = \frac{3(1-3\check{w})}{\Omega_0^2 r_0^2 \left(\sin^6 \frac{(1+3\check{w})\eta}{2r_0} \right)^{\frac{1+w}{1+3w}}} \text{ if } \check{w} \neq -1/3 \text{ and } R = \frac{6(1+b^2 r_0^2)}{\Omega_0^2 r_0^2 \exp(2b\eta)} \text{ if } \check{w} = -1/3,$$

so it is **positive** if $-1 \leq \check{w} < 1/3$ (recall the energy conditions $-1 \leq \check{w} \leq 1$).

Theorem

If $\Omega = \Omega(\eta)$ is such that $\check{g} = \Omega^2 g_{Einst}$ satisfies Einstein's equations, with $\Lambda = 0$, and with the energy momentum tensor \check{T} of a perfect fluid, whose pressure \check{p} is proportional to the energy density $\check{\mu}$, via $\check{p} = \check{w}\check{\mu}$, $\check{w} = const$, then

$\hat{g} = \frac{1}{\Omega^2} g_{Einst}$ satisfies Einstein's equations, with $\Lambda = 0$, and with the energy momentum tensor \hat{T} of a perfect fluid, whose pressure \hat{p} and the energy density $\hat{\mu}$ are related by $\hat{p} = \hat{w}\hat{\mu}$ with

$$\hat{w} = -\frac{1}{3}(2 + 3\check{w}).$$

The Ricci scalar of the metric \check{g} is

$$R = \frac{3(1-3\check{w})}{\Omega_0^2 r_0^2 \left(\sin^6 \frac{(1+3\check{w})\eta}{2r_0} \right)^{\frac{1+w}{1+3w}}} \text{ if } \check{w} \neq -1/3 \text{ and } R = \frac{6(1+b^2 r_0^2)}{\Omega_0^2 r_0^2 \exp(2b\eta)} \text{ if } \check{w} = -1/3,$$

so it is **positive** if $-1 \leq \check{w} < 1/3$ (recall the energy conditions $-1 \leq \check{w} \leq 1$).

Theorem

If $\Omega = \Omega(\eta)$ is such that $\check{g} = \Omega^2 g_{Einst}$ satisfies Einstein's equations, with $\Lambda = 0$, and with the energy momentum tensor \check{T} of a perfect fluid, whose pressure \check{p} is proportional to the energy density $\check{\mu}$, via $\check{p} = \check{w}\check{\mu}$, $\check{w} = const$, then

$\hat{g} = \frac{1}{\Omega^2} g_{Einst}$ satisfies Einstein's equations, with $\Lambda = 0$, and with the energy momentum tensor \hat{T} of a perfect fluid, whose pressure \hat{p} and the energy density $\hat{\mu}$ are related by $\hat{p} = \hat{w}\hat{\mu}$ with

$$\hat{w} = -\frac{1}{3}(2 + 3\check{w}).$$

The Ricci scalar of the metric \check{g} is

$$R = \frac{3(1-3\check{w})}{\Omega_0^2 r_0^2 \left(\sin^6 \frac{(1+3\check{w})\eta}{2r_0} \right)^{\frac{1+w}{1+3w}}} \text{ if } \check{w} \neq -1/3 \text{ and } R = \frac{6(1+b^2 r_0^2)}{\Omega_0^2 r_0^2 \exp(2b\eta)} \text{ if } \check{w} = -1/3,$$

so it is **positive** if $-1 \leq \check{w} < 1/3$ (recall the energy conditions $-1 \leq \check{w} \leq 1$).

Theorem

If $\Omega = \Omega(\eta)$ is such that $\check{g} = \Omega^2 g_{Einst}$ satisfies Einstein's equations, with $\Lambda = 0$, and with the energy momentum tensor \check{T} of a perfect fluid, whose pressure \check{p} is proportional to the energy density $\check{\mu}$, via $\check{p} = \check{w}\check{\mu}$, $\check{w} = const$, then

$\hat{g} = \frac{1}{\Omega^2} g_{Einst}$ satisfies Einstein's equations, with $\Lambda = 0$, and with the energy momentum tensor \hat{T} of a perfect fluid, whose pressure \hat{p} and the energy density $\hat{\mu}$ are related by $\hat{p} = \hat{w}\hat{\mu}$ with

$$\hat{w} = -\frac{1}{3}(2 + 3\check{w}).$$

The Ricci scalar of the metric \check{g} is

$$R = \frac{3(1-3\check{w})}{\Omega_0^2 r_0^2 \left(\sin^6 \frac{(1+3\check{w})\eta}{2r_0} \right)^{\frac{1+w}{1+3w}}} \text{ if } \check{w} \neq -1/3 \text{ and } R = \frac{6(1+b^2 r_0^2)}{\Omega_0^2 r_0^2 \exp(2b\eta)} \text{ if } \check{w} = -1/3,$$

so it is **positive** if $-1 \leq \check{w} < 1/3$ (recall the energy conditions $-1 \leq \check{w} \leq 1$).

Theorem

If $\Omega = \Omega(\eta)$ is such that $\check{g} = \Omega^2 g_{Einst}$ satisfies Einstein's equations, with $\Lambda = 0$, and with the energy momentum tensor \check{T} of a perfect fluid, whose pressure \check{p} is proportional to the energy density $\check{\mu}$, via $\check{p} = \check{w}\check{\mu}$, $\check{w} = const$, then

$\hat{g} = \frac{1}{\Omega^2} g_{Einst}$ satisfies Einstein's equations, with $\Lambda = 0$, and with the energy momentum tensor \hat{T} of a perfect fluid, whose pressure \hat{p} and the energy density $\hat{\mu}$ are related by $\hat{p} = \hat{w}\hat{\mu}$ with

$$\hat{w} = -\frac{1}{3}(2 + 3\check{w}).$$

The Ricci scalar of the metric \check{g} is

$$R = \frac{3(1-3\check{w})}{\Omega_0^2 r_0^2 \left(\sin^6 \frac{(1+3\check{w})\eta}{2r_0} \right)^{\frac{1+w}{1+3w}}} \text{ if } \check{w} \neq -1/3 \text{ and } R = \frac{6(1+b^2 r_0^2)}{\Omega_0^2 r_0^2 \exp(2b\eta)} \text{ if } \check{w} = -1/3,$$

so it is **positive** if $-1 \leq \check{w} < 1/3$ (recall the energy conditions $-1 \leq \check{w} \leq 1$).

Theorem

If $\Omega = \Omega(\eta)$ is such that $\check{g} = \Omega^2 g_{Einst}$ satisfies Einstein's equations, with $\Lambda = 0$, and with the energy momentum tensor \check{T} of a perfect fluid, whose pressure \check{p} is proportional to the energy density $\check{\mu}$, via $\check{p} = \check{w}\check{\mu}$, $\check{w} = const$, then

$\hat{g} = \frac{1}{\Omega^2} g_{Einst}$ satisfies Einstein's equations, with $\Lambda = 0$, and with the energy momentum tensor \hat{T} of a perfect fluid, whose pressure \hat{p} and the energy density $\hat{\mu}$ are related by $\hat{p} = \hat{w}\hat{\mu}$ with

$$\hat{w} = -\frac{1}{3}(2 + 3\check{w}).$$

The Ricci scalar of the metric \check{g} is

$$R = \frac{3(1-3\check{w})}{\Omega_0^2 r_0^2 \left(\sin^6 \frac{(1+3\check{w})\eta}{2r_0} \right)^{\frac{1+w}{1+3w}}} \text{ if } \check{w} \neq -1/3 \text{ and } R = \frac{6(1+b^2 r_0^2)}{\Omega_0^2 r_0^2 \exp(2b\eta)} \text{ if } \check{w} = -1/3,$$

so it is **positive** if $-1 \leq \check{w} < 1/3$ (recall the energy conditions $-1 \leq \check{w} \leq 1$).

Theorem

If $\Omega = \Omega(\eta)$ is such that $\check{g} = \Omega^2 g_{Einst}$ satisfies Einstein's equations, with $\Lambda = 0$, and with the energy momentum tensor \check{T} of a perfect fluid, whose pressure \check{p} is proportional to the energy density $\check{\mu}$, via $\check{p} = \check{w}\check{\mu}$, $\check{w} = const$, then

$\hat{g} = \frac{1}{\Omega^2} g_{Einst}$ satisfies Einstein's equations, with $\Lambda = 0$, and with the energy momentum tensor \hat{T} of a perfect fluid, whose pressure \hat{p} and the energy density $\hat{\mu}$ are related by $\hat{p} = \hat{w}\hat{\mu}$ with

$$\hat{w} = -\frac{1}{3}(2 + 3\check{w}).$$

The Ricci scalar of the metric \check{g} is

$$R = \frac{3(1-3\check{w})}{\Omega_0^2 r_0^2 \left(\sin^6 \frac{(1+3\check{w})\eta}{2r_0} \right)^{\frac{1+w}{1+3w}}} \text{ if } \check{w} \neq -1/3 \text{ and } R = \frac{6(1+b^2 r_0^2)}{\Omega_0^2 r_0^2 \exp(2b\eta)} \text{ if } \check{w} = -1/3,$$

so it is **positive** if $-1 \leq \check{w} < 1/3$ (recall the energy conditions $-1 \leq \check{w} \leq 1$).

Theorem

If $\Omega = \Omega(\eta)$ is such that $\check{g} = \Omega^2 g_{Einst}$ satisfies Einstein's equations, with $\Lambda = 0$, and with the energy momentum tensor \check{T} of a perfect fluid, whose pressure \check{p} is proportional to the energy density $\check{\mu}$, via $\check{p} = \check{w}\check{\mu}$, $\check{w} = const$, then

$\hat{g} = \frac{1}{\Omega^2} g_{Einst}$ satisfies Einstein's equations, with $\Lambda = 0$, and with the energy momentum tensor \hat{T} of a perfect fluid, whose pressure \hat{p} and the energy density $\hat{\mu}$ are related by $\hat{p} = \hat{w}\hat{\mu}$ with

$$\hat{w} = -\frac{1}{3}(2 + 3\check{w}).$$

The Ricci scalar of the metric \check{g} is

$$R = \frac{3(1-3\check{w})}{\Omega_0^2 r_0^2 \left(\sin^6 \frac{(1+3\check{w})\eta}{2r_0} \right)^{\frac{1+w}{1+3w}}} \text{ if } \check{w} \neq -1/3 \text{ and } R = \frac{6(1+b^2 r_0^2)}{\Omega_0^2 r_0^2 \exp(2b\eta)} \text{ if } \check{w} = -1/3,$$

so it is **positive** if $-1 \leq \check{w} < 1/3$ (recall the energy conditions $-1 \leq \check{w} \leq 1$).

Theorem

If $\Omega = \Omega(\eta)$ is such that $\check{g} = \Omega^2 g_{Einst}$ satisfies Einstein's equations, with $\Lambda = 0$, and with the energy momentum tensor \check{T} of a perfect fluid, whose pressure \check{p} is proportional to the energy density $\check{\mu}$, via $\check{p} = \check{w}\check{\mu}$, $\check{w} = const$, then

$\hat{g} = \frac{1}{\Omega^2} g_{Einst}$ satisfies Einstein's equations, with $\Lambda = 0$, and with the energy momentum tensor \hat{T} of a perfect fluid, whose pressure \hat{p} and the energy density $\hat{\mu}$ are related by $\hat{p} = \hat{w}\hat{\mu}$ with

$$\hat{w} = -\frac{1}{3}(2 + 3\check{w}).$$

The Ricci scalar of the metric \check{g} is

$$R = \frac{3(1-3\check{w})}{\Omega_0^2 r_0^2 \left(\sin^6 \frac{(1+3\check{w})\eta}{2r_0} \right)^{\frac{1+w}{1+3w}}} \text{ if } \check{w} \neq -1/3 \text{ and } R = \frac{6(1+b^2 r_0^2)}{\Omega_0^2 r_0^2 \exp(2b\eta)} \text{ if } \check{w} = -1/3,$$

so it is **positive** if $-1 \leq \check{w} < 1/3$ (recall the energy conditions $-1 \leq \check{w} \leq 1$).

Theorem

If $\Omega = \Omega(\eta)$ is such that $\check{g} = \Omega^2 g_{Einst}$ satisfies Einstein's equations, with $\Lambda = 0$, and with the energy momentum tensor \check{T} of a perfect fluid, whose pressure \check{p} is proportional to the energy density $\check{\mu}$, via $\check{p} = \check{w}\check{\mu}$, $\check{w} = const$, then

$\hat{g} = \frac{1}{\Omega^2} g_{Einst}$ satisfies Einstein's equations, with $\Lambda = 0$, and with the energy momentum tensor \hat{T} of a perfect fluid, whose pressure \hat{p} and the energy density $\hat{\mu}$ are related by $\hat{p} = \hat{w}\hat{\mu}$ with

$$\hat{w} = -\frac{1}{3}(2 + 3\check{w}).$$

The Ricci scalar of the metric \check{g} is

$$R = \frac{3(1-3\check{w})}{\Omega_0^2 r_0^2 \left(\sin^6 \frac{(1+3\check{w})\eta}{2r_0} \right)^{\frac{1+w}{1+3w}}} \text{ if } \check{w} \neq -1/3 \text{ and } R = \frac{6(1+b^2 r_0^2)}{\Omega_0^2 r_0^2 \exp(2b\eta)} \text{ if } \check{w} = -1/3,$$

so it is **positive** if $-1 \leq \check{w} < 1/3$ (recall the energy conditions $-1 \leq \check{w} \leq 1$).

Theorem

If $\Omega = \Omega(\eta)$ is such that $\check{g} = \Omega^2 g_{Einst}$ satisfies Einstein's equations, with $\Lambda = 0$, and with the energy momentum tensor \check{T} of a perfect fluid, whose pressure \check{p} is proportional to the energy density $\check{\mu}$, via $\check{p} = \check{w}\check{\mu}$, $\check{w} = const$, then

$\hat{g} = \frac{1}{\Omega^2} g_{Einst}$ satisfies Einstein's equations, with $\Lambda = 0$, and with the energy momentum tensor \hat{T} of a perfect fluid, whose pressure \hat{p} and the energy density $\hat{\mu}$ are related by $\hat{p} = \hat{w}\hat{\mu}$ with

$$\hat{w} = -\frac{1}{3}(2 + 3\check{w}).$$

The Ricci scalar of the metric \check{g} is

$$R = \frac{3(1-3\check{w})}{\Omega_0^2 r_0^2 \left(\sin^6 \frac{(1+3\check{w})\eta}{2r_0} \right)^{\frac{1+w}{1+3w}}} \text{ if } \check{w} \neq -1/3 \text{ and } R = \frac{6(1+b^2 r_0^2)}{\Omega_0^2 r_0^2 \exp(2b\eta)} \text{ if } \check{w} = -1/3,$$

so it is **positive** if $-1 \leq \check{w} < 1/3$ (recall the energy conditions $-1 \leq \check{w} \leq 1$).

Theorem

If $\Omega = \Omega(\eta)$ is such that $\check{g} = \Omega^2 g_{Einst}$ satisfies Einstein's equations, with $\Lambda = 0$, and with the energy momentum tensor \check{T} of a perfect fluid, whose pressure \check{p} is proportional to the energy density $\check{\mu}$, via $\check{p} = \check{w}\check{\mu}$, $\check{w} = const$, then

$\hat{g} = \frac{1}{\Omega^2} g_{Einst}$ satisfies Einstein's equations, with $\Lambda = 0$, and with the energy momentum tensor \hat{T} of a perfect fluid, whose pressure \hat{p} and the energy density $\hat{\mu}$ are related by $\hat{p} = \hat{w}\hat{\mu}$ with

$$\hat{w} = -\frac{1}{3}(2 + 3\check{w}).$$

The Ricci scalar of the metric \check{g} is

$$R = \frac{3(1-3\check{w})}{\Omega_0^2 r_0^2 \left(\sin^6 \frac{(1+3\check{w})\eta}{2r_0} \right)^{\frac{1+w}{1+3w}}} \text{ if } \check{w} \neq -1/3 \text{ and } R = \frac{6(1+b^2 r_0^2)}{\Omega_0^2 r_0^2 \exp(2b\eta)} \text{ if } \check{w} = -1/3,$$

so it is positive if $-1 \leq \check{w} < 1/3$ (recall the energy conditions $-1 \leq \check{w} \leq 1$).

Theorem

If $\Omega = \Omega(\eta)$ is such that $\check{g} = \Omega^2 g_{Einst}$ satisfies Einstein's equations, with $\Lambda = 0$, and with the energy momentum tensor \check{T} of a perfect fluid, whose pressure \check{p} is proportional to the energy density $\check{\mu}$, via $\check{p} = \check{w}\check{\mu}$, $\check{w} = const$, then

$\hat{g} = \frac{1}{\Omega^2} g_{Einst}$ satisfies Einstein's equations, with $\Lambda = 0$, and with the energy momentum tensor \hat{T} of a perfect fluid, whose pressure \hat{p} and the energy density $\hat{\mu}$ are related by $\hat{p} = \hat{w}\hat{\mu}$ with

$$\hat{w} = -\frac{1}{3}(2 + 3\check{w}).$$

The Ricci scalar of the metric \check{g} is

$$R = \frac{3(1-3\check{w})}{\Omega_0^2 r_0^2 \left(\sin^6 \frac{(1+3\check{w})\eta}{2r_0} \right)^{\frac{1+w}{1+3w}}} \text{ if } \check{w} \neq -1/3 \text{ and } R = \frac{6(1+b^2 r_0^2)}{\Omega_0^2 r_0^2 \exp(2b\eta)} \text{ if } \check{w} = -1/3,$$

so it is positive if $-1 \leq \check{w} < 1/3$ (recall the energy conditions $-1 \leq \check{w} \leq 1$).

Theorem

If $\Omega = \Omega(\eta)$ is such that $\check{g} = \Omega^2 g_{Einst}$ satisfies Einstein's equations, with $\Lambda = 0$, and with the energy momentum tensor \check{T} of a perfect fluid, whose pressure \check{p} is proportional to the energy density $\check{\mu}$, via $\check{p} = \check{w}\check{\mu}$, $\check{w} = const$, then

$\hat{g} = \frac{1}{\Omega^2} g_{Einst}$ satisfies Einstein's equations, with $\Lambda = 0$, and with the energy momentum tensor \hat{T} of a perfect fluid, whose pressure \hat{p} and the energy density $\hat{\mu}$ are related by $\hat{p} = \hat{w}\hat{\mu}$ with

$$\hat{w} = -\frac{1}{3}(2 + 3\check{w}).$$

The Ricci scalar of the metric \check{g} is

$$R = \frac{3(1-3\check{w})}{\Omega_0^2 r_0^2 \left(\sin^6 \frac{(1+3\check{w})\eta}{2r_0} \right)^{\frac{1+w}{1+3w}}} \text{ if } \check{w} \neq -1/3 \text{ and } R = \frac{6(1+b^2 r_0^2)}{\Omega_0^2 r_0^2 \exp(2b\eta)} \text{ if } \check{w} = -1/3,$$

so it is **positive** if $-1 \leq \check{w} < 1/3$ (recall the energy conditions $-1 \leq \check{w} \leq 1$).

Theorem

If $\Omega = \Omega(\eta)$ is such that $\check{g} = \Omega^2 g_{Einst}$ satisfies Einstein's equations, with $\Lambda = 0$, and with the energy momentum tensor \check{T} of a perfect fluid, whose pressure \check{p} is proportional to the energy density $\check{\mu}$, via $\check{p} = \check{w}\check{\mu}$, $\check{w} = const$, then

$\hat{g} = \frac{1}{\Omega^2} g_{Einst}$ satisfies Einstein's equations, with $\Lambda = 0$, and with the energy momentum tensor \hat{T} of a perfect fluid, whose pressure \hat{p} and the energy density $\hat{\mu}$ are related by $\hat{p} = \hat{w}\hat{\mu}$ with

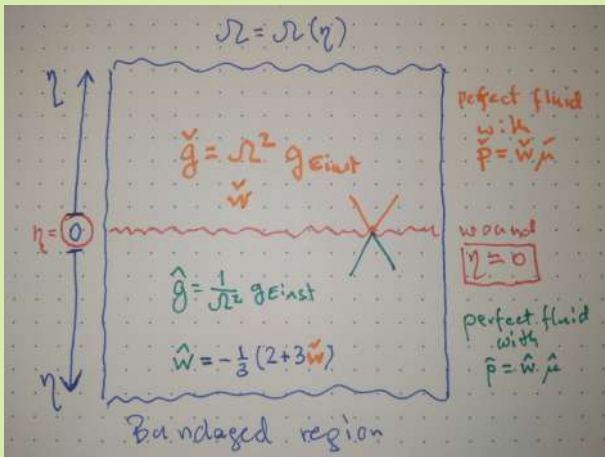
$$\hat{w} = -\frac{1}{3}(2 + 3\check{w}).$$

The Ricci scalar of the metric \check{g} is

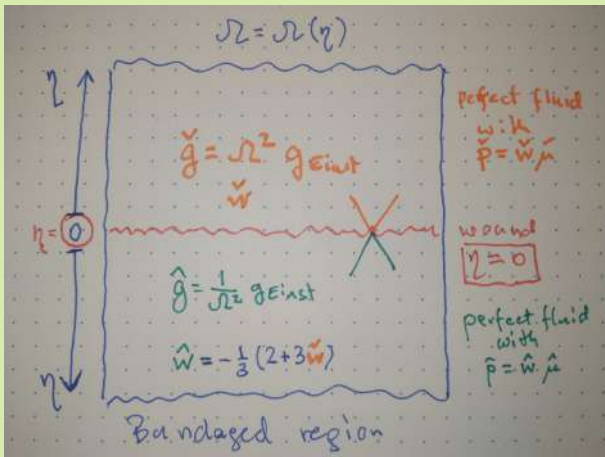
$$R = \frac{3(1-3\check{w})}{\Omega_0^2 r_0^2 \left(\sin^6 \frac{(1+3\check{w})\eta}{2r_0} \right)^{\frac{1+w}{1+3w}}} \text{ if } \check{w} \neq -1/3 \text{ and } R = \frac{6(1+b^2 r_0^2)}{\Omega_0^2 r_0^2 \exp(2b\eta)} \text{ if } \check{w} = -1/3,$$

so it is **positive** if $-1 \leq \check{w} < 1/3$ (recall the energy conditions $-1 \leq \check{w} \leq 1$).

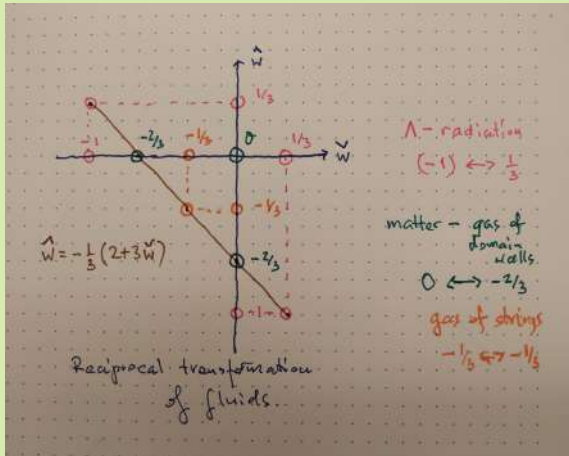
Transformation $\check{g} \rightarrow \hat{g} \rightarrow \check{g} \rightarrow \hat{g} \rightarrow \dots$ of fluids



Transformation $\check{g} \rightarrow \hat{g} \rightarrow \check{g} \rightarrow \hat{g} \rightarrow \dots$ of fluids

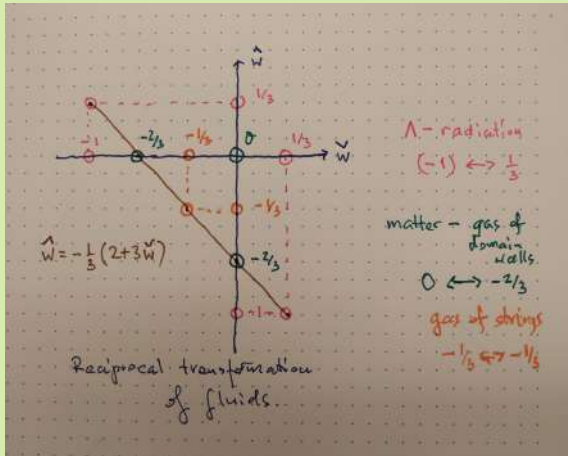


Transformation $\check{g} \rightarrow \hat{g} \rightarrow \check{g} \rightarrow \hat{g} \rightarrow \dots$ of fluids



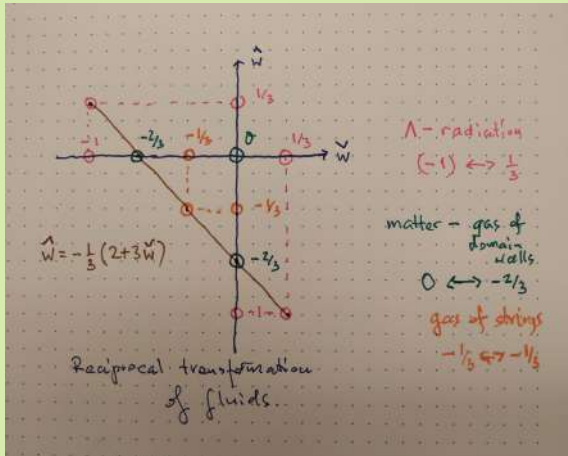
Suspicious points: $\check{w} = -1, 1/3$ (cosmological constant - radiation), since the scalar curvature $R = 0$, when $\check{w} = 1/3$; and $\check{w} = -1/3$ (gas of strings), when $\Omega \neq 0$ on \mathcal{I} .

Transformation $\check{g} \rightarrow \hat{g} \rightarrow \check{g} \rightarrow \hat{g} \rightarrow \dots$ of fluids



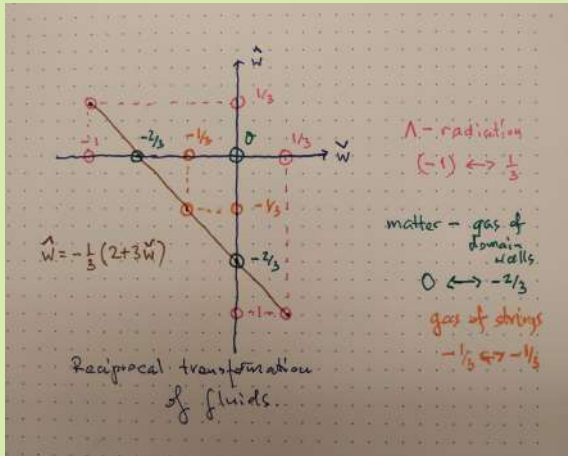
Suspicious points: $\check{w} = -1, 1/3$ (cosmological constant - radiation), since the scalar curvature $R = 0$, when $\check{w} = 1/3$; and $\check{w} = -1/3$ (gas of strings), when $\Omega \neq 0$ on \mathcal{I} .

Transformation $\check{g} \rightarrow \hat{g} \rightarrow \check{g} \rightarrow \hat{g} \rightarrow \dots$ of fluids

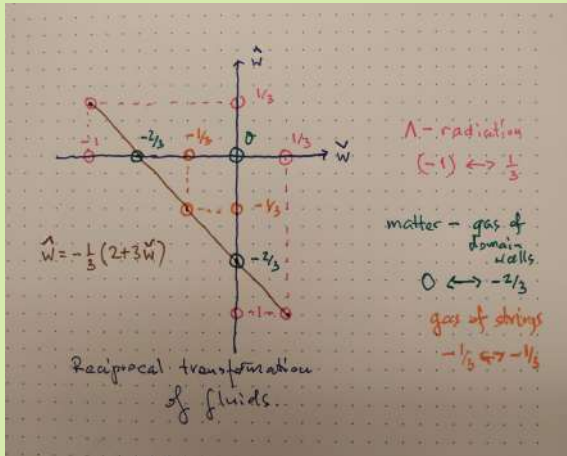


Suspicious points: $\check{w} = -1, 1/3$ (cosmological constant - radiation), since the scalar curvature $R = 0$, when $\check{w} = 1/3$; and $\check{w} = -1/3$ (gas of strings), when $\Omega \neq 0$ on \mathcal{I} .

Transformation $\check{g} \rightarrow \hat{g} \rightarrow \check{g} \rightarrow \hat{g} \rightarrow \dots$ of fluids



Suspicious points: $\check{w} = -1, 1/3$ (cosmological constant - radiation), since the scalar curvature $R = 0$, when $\check{w} = 1/3$; and $\check{w} = -1/3$ (gas of strings), when $\Omega \neq 0$ on \mathcal{I} .



Suspicious points: $\check{w} = -1, 1/3$ (cosmological constant - radiation), since the scalar curvature $R = 0$, when $\check{w} = 1/3$; and $\check{w} = -1/3$ (gas of strings), when $\Omega \neq 0$ on \mathcal{I} .

Transformation $\check{g} \rightarrow \hat{g} \rightarrow \check{g} \rightarrow \hat{g} \rightarrow \dots$ of fluids: more careful approach

- We come back to the FLRW metric $\check{g} = -dt^2 + \Omega^2(t)r_0^2 g_{S^3}$.
- We write it as $\check{g} = \Omega^2(t) \left(-\frac{dt^2}{\Omega^2(t)} + r_0^2 g_{S^3} \right)$, so that it is clear that $\check{g} = \Omega^2(t) g_{Einst}$.
- Then the condition that \check{g} satisfies perfect fluid Einstein's equations with $\check{u} = -dt$, $\check{\rho} = \check{w}\check{\mu}$, and the cosmological constant $\check{\Lambda}$, is equivalent to the following ODE for Ω :

$$2r_0^2 \Omega \Omega'' = -(1 + 3\check{w})(1 + r_0^2 \Omega'^2) + (1 + \check{w})\check{\Lambda} r_0^2 \Omega^2.$$

- We want that $\check{w} = \text{const}$ and that $\hat{g} = \frac{1}{\Omega^2} g_{Einst}$ satisfies perfect fluid Einstein's equations with $\hat{u} = -\frac{dt}{\Omega^2}$, $\hat{\rho} = \hat{w}\hat{\mu}$, the cosmological constant $\hat{\Lambda}$, and $\hat{w} = \text{const}$.
- From the Einstein's equations for \hat{g} we easily calculate \hat{w} , and forcing it to be constant, because of the above ODE satisfied by Ω , we find that it is possible provided that:

$$\check{\Lambda} \hat{\Lambda} (1 + \check{w})(1 - 3\check{w}) = 0.$$

- Thus, a **necessary condition** for both Ω and Ω^{-1} to describe the polytropes, is that **either** one of the Λ s is zero, **or** \check{w} is of the 'radiation- Λ ' type.

- We come back to the FLRW metric $\check{g} = -dt^2 + \Omega^2(t)r_0^2 g_{\mathbb{S}^3}$.
- We write it as $\check{g} = \Omega^2(t) \left(-\frac{dt^2}{\Omega^2(t)} + r_0^2 g_{\mathbb{S}^3} \right)$, so that it is clear that $\check{g} = \Omega^2(t) g_{Einst}$.
- Then the condition that \check{g} satisfies perfect fluid Einstein's equations with $\check{u} = -dt$, $\check{p} = \check{w}\check{\mu}$, and the cosmological constant $\check{\Lambda}$, is equivalent to the following ODE for Ω :

$$2r_0^2 \Omega \Omega'' = -(1 + 3\check{w})(1 + r_0^2 \Omega'^2) + (1 + \check{w})\check{\Lambda} r_0^2 \Omega^2.$$

- We want that $\check{w} = \text{const}$ and that $\hat{g} = \frac{1}{\Omega^2} g_{Einst}$ satisfies perfect fluid Einstein's equations with $\hat{u} = -\frac{dt}{\Omega^2}$, $\hat{p} = \hat{w}\hat{\mu}$, the cosmological constant $\hat{\Lambda}$, and $\hat{w} = \text{const}$.
- From the Einstein's equations for \hat{g} we easily calculate \hat{w} , and forcing it to be constant, because of the above ODE satisfied by Ω , we find that it is possible provided that:

$$\check{\Lambda} \hat{\Lambda} (1 + \check{w})(1 - 3\check{w}) = 0.$$

- Thus, a **necessary condition** for both Ω and Ω^{-1} to describe the polytropes, is that **either** one of the Λ s is zero, **or** \check{w} is of the 'radiation- Λ ' type.

Transformation $\check{g} \rightarrow \hat{g} \rightarrow \check{g} \rightarrow \hat{g} \rightarrow \dots$ of fluids: more careful approach

- We come back to the FLRW metric $\check{g} = -dt^2 + \Omega^2(t)r_0^2 g_{\mathbb{S}^3}$.
- We write it as $\check{g} = \Omega^2(t) \left(-\frac{dt^2}{\Omega^2(t)} + r_0^2 g_{\mathbb{S}^3} \right)$, so that it is clear that $\check{g} = \Omega^2(t) g_{Einst}$.
- Then the condition that \check{g} satisfies perfect fluid Einstein's equations with $\check{u} = -dt$, $\check{p} = \check{w}\check{\mu}$, and the cosmological constant $\check{\Lambda}$, is equivalent to the following ODE for Ω :

$$2r_0^2 \Omega \Omega'' = -(1 + 3\check{w})(1 + r_0^2 \Omega'^2) + (1 + \check{w})\check{\Lambda} r_0^2 \Omega^2.$$

- We want that $\check{w} = const$ and that $\hat{g} = \frac{1}{\Omega^2} g_{Einst}$ satisfies perfect fluid Einstein's equations with $\hat{u} = -\frac{dt}{\Omega^2}$, $\hat{p} = \hat{w}\hat{\mu}$, the cosmological constant $\hat{\Lambda}$, and $\hat{w} = const$.
- From the Einstein's equations for \hat{g} we easily calculate \hat{w} , and forcing it to be constant, because of the above ODE satisfied by Ω , we find that it is possible provided that:

$$\check{\Lambda} \hat{\Lambda} (1 + \check{w})(1 - 3\check{w}) = 0.$$

- Thus, a **necessary condition** for both Ω and Ω^{-1} to describe the polytropes, is that **either** one of the Λ s is zero, **or** \check{w} is of the 'radiation- Λ ' type.

Transformation $\check{g} \rightarrow \hat{g} \rightarrow \check{g} \rightarrow \hat{g} \rightarrow \dots$ of fluids: more careful approach

- We come back to the FLRW metric $\check{g} = -dt^2 + \Omega^2(t)r_0^2 g_{\mathbb{S}^3}$.
- We write it as $\check{g} = \Omega^2(t) \left(-\frac{dt^2}{\Omega^2(t)} + r_0^2 g_{\mathbb{S}^3} \right)$, so that it is clear that $\check{g} = \Omega^2(t) g_{Einst}$.
- Then the condition that \check{g} satisfies perfect fluid Einstein's equations with $\check{u} = -dt$, $\check{p} = \check{w}\check{\mu}$, and the cosmological constant $\check{\Lambda}$, is equivalent to the following ODE for Ω :

$$2r_0^2 \Omega \Omega'' = -(1 + 3\check{w})(1 + r_0^2 \Omega'^2) + (1 + \check{w})\check{\Lambda} r_0^2 \Omega^2.$$

- We want that $\check{w} = const$ and that $\hat{g} = \frac{1}{\Omega^2} g_{Einst}$ satisfies perfect fluid Einstein's equations with $\hat{u} = -\frac{dt}{\Omega^2}$, $\hat{p} = \hat{w}\hat{\mu}$, the cosmological constant $\hat{\Lambda}$, and $\hat{w} = const$.
- From the Einstein's equations for \hat{g} we easily calculate \hat{w} , and forcing it to be constant, because of the above ODE satisfied by Ω , we find that it is possible provided that:

$$\check{\Lambda} \hat{\Lambda} (1 + \check{w})(1 - 3\check{w}) = 0.$$

- Thus, a **necessary condition** for both Ω and Ω^{-1} to describe the polytropes, is that **either** one of the Λ s is zero, **or** \check{w} is of the 'radiation- Λ ' type.

- We come back to the FLRW metric $\check{g} = -dt^2 + \Omega^2(t)r_0^2 g_{\mathbb{S}^3}$.
- We write it as $\check{g} = \Omega^2(t) \left(-\frac{dt^2}{\Omega^2(t)} + r_0^2 g_{\mathbb{S}^3} \right)$, so that it is clear that $\check{g} = \Omega^2(t) g_{Einst}$.
- Then the condition that \check{g} satisfies perfect fluid Einstein's equations with $\check{u} = -dt$, $\check{p} = \check{w}\check{\mu}$, and the cosmological constant $\check{\Lambda}$, is equivalent to the following ODE for Ω :

$$2r_0^2 \Omega \Omega'' = -(1 + 3\check{w})(1 + r_0^2 \Omega'^2) + (1 + \check{w})\check{\Lambda} r_0^2 \Omega^2.$$

- We want that $\check{w} = const$ and that $\hat{g} = \frac{1}{\Omega^2} g_{Einst}$ satisfies perfect fluid Einstein's equations with $\hat{u} = -\frac{dt}{\Omega^2}$, $\hat{p} = \hat{w}\hat{\mu}$, the cosmological constant $\hat{\Lambda}$, and $\hat{w} = const$.
- From the Einstein's equations for \hat{g} we easily calculate \hat{w} , and forcing it to be constant, because of the above ODE satisfied by Ω , we find that it is possible provided that:

$$\check{\Lambda} \hat{\Lambda} (1 + \check{w})(1 - 3\check{w}) = 0.$$

- Thus, a **necessary condition** for both Ω and Ω^{-1} to describe the polytropes, is that **either** one of the Λ s is zero, **or** \check{w} is of the 'radiation- Λ ' type.

- We come back to the FLRW metric $\check{g} = -dt^2 + \Omega^2(t)r_0^2 g_{\mathbb{S}^3}$.
- We write it as $\check{g} = \Omega^2(t) \left(-\frac{dt^2}{\Omega^2(t)} + r_0^2 g_{\mathbb{S}^3} \right)$, so that it is clear that $\check{g} = \Omega^2(t) g_{Einst}$.
- Then the condition that \check{g} satisfies perfect fluid Einstein's equations with $\check{u} = -dt$, $\check{p} = \check{w}\check{\mu}$, and the cosmological constant $\check{\Lambda}$, is equivalent to the following ODE for Ω :

$$2r_0^2 \Omega \Omega'' = -(1 + 3\check{w})(1 + r_0^2 \Omega'^2) + (1 + \check{w})\check{\Lambda} r_0^2 \Omega^2.$$

- We want that $\check{w} = const$ and that $\hat{g} = \frac{1}{\Omega^2} g_{Einst}$ satisfies perfect fluid Einstein's equations with $\hat{u} = -\frac{dt}{\Omega^2}$, $\hat{p} = \hat{w}\hat{\mu}$, the cosmological constant $\hat{\Lambda}$, and $\hat{w} = const$.
- From the Einstein's equations for \hat{g} we easily calculate \hat{w} , and forcing it to be constant, because of the above ODE satisfied by Ω , we find that it is possible provided that:

$$\check{\Lambda} \hat{\Lambda} (1 + \check{w})(1 - 3\check{w}) = 0.$$

- Thus, a **necessary condition** for both Ω and Ω^{-1} to describe the polytropes, is that **either** one of the Λ s is zero, **or** \check{w} is of the 'radiation- Λ ' type.

- We come back to the FLRW metric $\check{g} = -dt^2 + \Omega^2(t)r_0^2 g_{\mathbb{S}^3}$.
- We write it as $\check{g} = \Omega^2(t) \left(-\frac{dt^2}{\Omega^2(t)} + r_0^2 g_{\mathbb{S}^3} \right)$, so that it is clear that $\check{g} = \Omega^2(t) g_{Einst}$.
- Then the condition that \check{g} satisfies perfect fluid Einstein's equations with $\check{u} = -dt$, $\check{p} = \check{w}\check{\mu}$, and the cosmological constant $\check{\Lambda}$, is equivalent to the following ODE for Ω :

$$2r_0^2 \Omega \Omega'' = -(1 + 3\check{w})(1 + r_0^2 \Omega'^2) + (1 + \check{w})\check{\Lambda} r_0^2 \Omega^2.$$

- We want that $\check{w} = \text{const}$ and that $\hat{g} = \frac{1}{\Omega^2} g_{Einst}$ satisfies perfect fluid Einstein's equations with $\hat{u} = -\frac{dt}{\Omega^2}$, $\hat{p} = \hat{w}\hat{\mu}$, the cosmological constant $\hat{\Lambda}$, and $\hat{w} = \text{const}$.
- From the Einstein's equations for \hat{g} we easily calculate \hat{w} , and forcing it to be constant, because of the above ODE satisfied by Ω , we find that it is possible provided that:

$$\check{\Lambda} \hat{\Lambda} (1 + \check{w})(1 - 3\check{w}) = 0.$$

- Thus, a **necessary condition** for both Ω and Ω^{-1} to describe the polytropes, is that **either** one of the Λ s is zero, **or** \check{w} is of the 'radiation- Λ ' type.

- We come back to the FLRW metric $\check{g} = -dt^2 + \Omega^2(t)r_0^2 g_{\mathbb{S}^3}$.
- We write it as $\check{g} = \Omega^2(t) \left(-\frac{dt^2}{\Omega^2(t)} + r_0^2 g_{\mathbb{S}^3} \right)$, so that it is clear that $\check{g} = \Omega^2(t) g_{Einst}$.
- Then the condition that \check{g} satisfies perfect fluid Einstein's equations with $\check{u} = -dt$, $\check{p} = \check{w}\check{\mu}$, and the cosmological constant $\check{\Lambda}$, is equivalent to the following ODE for Ω :

$$2r_0^2 \Omega \Omega'' = -(1 + 3\check{w})(1 + r_0^2 \Omega'^2) + (1 + \check{w})\check{\Lambda} r_0^2 \Omega^2.$$

- We want that $\check{w} = \text{const}$ and that $\hat{g} = \frac{1}{\Omega^2} g_{Einst}$ satisfies perfect fluid Einstein's equations with $\hat{u} = -\frac{dt}{\Omega^2}$, $\hat{p} = \hat{w}\hat{\mu}$, the cosmological constant $\hat{\Lambda}$, and $\hat{w} = \text{const}$.
- From the Einstein's equations for \hat{g} we easily calculate \hat{w} , and forcing it to be constant, because of the above ODE satisfied by Ω , we find that it is possible provided that:

$$\check{\Lambda} \hat{\Lambda} (1 + \check{w})(1 - 3\check{w}) = 0.$$

- Thus, a **necessary condition** for both Ω and Ω^{-1} to describe the polytropes, is that **either** one of the Λ s is zero, **or** \check{w} is of the 'radiation- Λ ' type.

- We come back to the FLRW metric $\check{g} = -dt^2 + \Omega^2(t)r_0^2 g_{\mathbb{S}^3}$.
- We write it as $\check{g} = \Omega^2(t) \left(-\frac{dt^2}{\Omega^2(t)} + r_0^2 g_{\mathbb{S}^3} \right)$, so that it is clear that $\check{g} = \Omega^2(t) g_{Einst}$.
- Then the condition that \check{g} satisfies perfect fluid Einstein's equations with $\check{u} = -dt$, $\check{p} = \check{w}\check{\mu}$, and the cosmological constant $\check{\Lambda}$, is equivalent to the following ODE for Ω :

$$2r_0^2 \Omega \Omega'' = -(1 + 3\check{w})(1 + r_0^2 \Omega'^2) + (1 + \check{w})\check{\Lambda} r_0^2 \Omega^2.$$

- We want that $\check{w} = \text{const}$ and that $\hat{g} = \frac{1}{\Omega^2} g_{Einst}$ satisfies perfect fluid Einstein's equations with $\hat{u} = -\frac{dt}{\Omega^2}$, $\hat{p} = \hat{w}\hat{\mu}$, the cosmological constant $\hat{\Lambda}$, and $\hat{w} = \text{const}$.
- From the Einstein's equations for \hat{g} we easily calculate \hat{w} , and forcing it to be constant, because of the above ODE satisfied by Ω , we find that it is possible provided that:

$$\check{\Lambda} \hat{\Lambda} (1 + \check{w})(1 - 3\check{w}) = 0.$$

- Thus, a **necessary condition** for both Ω and Ω^{-1} to describe the polytropes, is that **either** one of the Λ s is zero, **or** \check{w} is of the 'radiation- Λ ' type.

- We come back to the FLRW metric $\check{g} = -dt^2 + \Omega^2(t)r_0^2 g_{\mathbb{S}^3}$.
- We write it as $\check{g} = \Omega^2(t) \left(-\frac{dt^2}{\Omega^2(t)} + r_0^2 g_{\mathbb{S}^3} \right)$, so that it is clear that $\check{g} = \Omega^2(t) g_{Einst}$.
- Then the condition that \check{g} satisfies perfect fluid Einstein's equations with $\check{u} = -dt$, $\check{p} = \check{w}\check{\mu}$, and the cosmological constant $\check{\Lambda}$, is equivalent to the following ODE for Ω :

$$2r_0^2 \Omega \Omega'' = -(1 + 3\check{w})(1 + r_0^2 \Omega'^2) + (1 + \check{w})\check{\Lambda} r_0^2 \Omega^2.$$

- We want that $\check{w} = \text{const}$ and that $\hat{g} = \frac{1}{\Omega^2} g_{Einst}$ satisfies perfect fluid Einstein's equations with $\hat{u} = -\frac{dt}{\Omega^2}$, $\hat{p} = \hat{w}\hat{\mu}$, the cosmological constant $\hat{\Lambda}$, and $\hat{w} = \text{const}$.
- From the Einstein's equations for \hat{g} we easily calculate \hat{w} , and forcing it to be constant, because of the above ODE satisfied by Ω , we find that it is possible provided that:

$$\check{\Lambda} \hat{\Lambda} (1 + \check{w})(1 - 3\check{w}) = 0.$$

- Thus, a **necessary condition** for both Ω and Ω^{-1} to describe the polytropes, is that **either** one of the Λ s is zero, **or** \check{w} is of the 'radiation- Λ ' type.

- We come back to the FLRW metric $\check{g} = -dt^2 + \Omega^2(t)r_0^2 g_{\mathbb{S}^3}$.
- We write it as $\check{g} = \Omega^2(t) \left(-\frac{dt^2}{\Omega^2(t)} + r_0^2 g_{\mathbb{S}^3} \right)$, so that it is clear that $\check{g} = \Omega^2(t) g_{Einst}$.
- Then the condition that \check{g} satisfies perfect fluid Einstein's equations with $\check{u} = -dt$, $\check{p} = \check{w}\check{\mu}$, and the cosmological constant $\check{\Lambda}$, is equivalent to the following ODE for Ω :

$$2r_0^2 \Omega \Omega'' = -(1 + 3\check{w})(1 + r_0^2 \Omega'^2) + (1 + \check{w})\check{\Lambda} r_0^2 \Omega^2.$$

- We want that $\check{w} = \text{const}$ and that $\hat{g} = \frac{1}{\Omega^2} g_{Einst}$ satisfies perfect fluid Einstein's equations with $\hat{u} = -\frac{dt}{\Omega^2}$, $\hat{p} = \hat{w}\hat{\mu}$, the cosmological constant $\hat{\Lambda}$, and $\hat{w} = \text{const}$.
- From the Einstein's equations for \hat{g} we easily calculate \hat{w} , and forcing it to be constant, because of the above ODE satisfied by Ω , we find that it is possible provided that:

$$\check{\Lambda} \hat{\Lambda} (1 + \check{w})(1 - 3\check{w}) = 0.$$

- Thus, a **necessary condition** for both Ω and Ω^{-1} to describe the polytropes, is that **either** one of the Λ s is zero, **or** \check{w} is of the 'radiation- Λ ' type.

- We come back to the FLRW metric $\check{g} = -dt^2 + \Omega^2(t)r_0^2 g_{\mathbb{S}^3}$.
- We write it as $\check{g} = \Omega^2(t) \left(-\frac{dt^2}{\Omega^2(t)} + r_0^2 g_{\mathbb{S}^3} \right)$, so that it is clear that $\check{g} = \Omega^2(t) g_{Einst}$.
- Then the condition that \check{g} satisfies perfect fluid Einstein's equations with $\check{u} = -dt$, $\check{p} = \check{w}\check{\mu}$, and the cosmological constant $\check{\Lambda}$, is equivalent to the following ODE for Ω :

$$2r_0^2 \Omega \Omega'' = -(1 + 3\check{w})(1 + r_0^2 \Omega'^2) + (1 + \check{w})\check{\Lambda} r_0^2 \Omega^2.$$

- We want that $\check{w} = \text{const}$ and that $\hat{g} = \frac{1}{\Omega^2} g_{Einst}$ satisfies perfect fluid Einstein's equations with $\hat{u} = -\frac{dt}{\Omega^2}$, $\hat{p} = \hat{w}\hat{\mu}$, the cosmological constant $\hat{\Lambda}$, and $\hat{w} = \text{const}$.
- From the Einstein's equations for \hat{g} we easily calculate \hat{w} , and forcing it to be constant, because of the above ODE satisfied by Ω , we find that it is possible provided that:

$$\check{\Lambda} \hat{\Lambda} (1 + \check{w})(1 - 3\check{w}) = 0.$$

- Thus, a **necessary condition** for both Ω and Ω^{-1} to describe the polytropes, is that **either** one of the Λ s is zero, **or** \check{w} is of the 'radiation- Λ ' type.

- We come back to the FLRW metric $\check{g} = -dt^2 + \Omega^2(t)r_0^2 g_{\mathbb{S}^3}$.
- We write it as $\check{g} = \Omega^2(t) \left(-\frac{dt^2}{\Omega^2(t)} + r_0^2 g_{\mathbb{S}^3} \right)$, so that it is clear that $\check{g} = \Omega^2(t) g_{Einst}$.
- Then the condition that \check{g} satisfies perfect fluid Einstein's equations with $\check{u} = -dt$, $\check{p} = \check{w}\check{\mu}$, and the cosmological constant $\check{\Lambda}$, is equivalent to the following ODE for Ω :

$$2r_0^2 \Omega \Omega'' = -(1 + 3\check{w})(1 + r_0^2 \Omega'^2) + (1 + \check{w})\check{\Lambda} r_0^2 \Omega^2.$$

- We want that $\check{w} = \text{const}$ and that $\hat{g} = \frac{1}{\Omega^2} g_{Einst}$ satisfies perfect fluid Einstein's equations with $\hat{u} = -\frac{dt}{\Omega^2}$, $\hat{p} = \hat{w}\hat{\mu}$, the cosmological constant $\hat{\Lambda}$, and $\hat{w} = \text{const}$.
- From the Einstein's equations for \hat{g} we easily calculate \hat{w} , and forcing it to be constant, because of the above ODE satisfied by Ω , we find that it is possible provided that:

$$\check{\Lambda} \hat{\Lambda} (1 + \check{w})(1 - 3\check{w}) = 0.$$

- Thus, a **necessary condition** for both Ω and Ω^{-1} to describe the polytropes, is that **either** one of the Λ s is zero, **or** \check{w} is of the 'radiation- Λ ' type.

- Considering the case $\check{w} = 1/3$, one shows that **remarkably** $\hat{w} = 1/3$ (generalization of the result of **Paul Tod**). More explicitly this case can be integrated to the very end.

- Theorem.** The function $\Omega = \Omega(t)$ given by:

$$\Omega^2 = \frac{3 - 3 \cosh(2\sqrt{\frac{\check{\lambda}}{3}}t) - 2r_0^2 \sqrt{\check{\lambda}\hat{\lambda}} \sinh(2\sqrt{\frac{\check{\lambda}}{3}}t)}{\check{\lambda}r_0^2}$$

has the property that both $\check{g} = \Omega^2 g_{Einst}$ and $\hat{g} = \Omega^{-2} g_{Einst}$ satisfy Einstein's equations with polytropic perfect fluid equation of state, for which $\hat{w} = \check{w} = 1/3$ (radiation), and with the corresponding cosmological constants $\check{\lambda}$ and $\hat{\lambda}$. Here $g_{Einst} = -\Omega^{-2} dt^2 + r_0^2 g_{\mathbb{S}^3}$.

- Colloquially speaking **incoherent radiation passes happily through the wound**. However, cosmological constants can change from any positive value to any other one. Ha...

- Considering the case $\check{w} = 1/3$, one shows that **remarkably** $\hat{w} = 1/3$ (generalization of the result of **Paul Tod**). More explicitly this case can be integrated to the very end.

- Theorem.** The function $\Omega = \Omega(t)$ given by:

$$\Omega^2 = \frac{3 - 3 \cosh(2\sqrt{\frac{\check{\lambda}}{3}}t) - 2r_0^2 \sqrt{\check{\lambda}\hat{\lambda}} \sinh(2\sqrt{\frac{\check{\lambda}}{3}}t)}{\check{\lambda}r_0^2}$$

has the property that both $\check{g} = \Omega^2 g_{Einst}$ and $\hat{g} = \Omega^{-2} g_{Einst}$ satisfy Einstein's equations with polytropic perfect fluid equation of state, for which $\hat{w} = \check{w} = 1/3$ (radiation), and with the corresponding cosmological constants $\check{\lambda}$ and $\hat{\lambda}$. Here $g_{Einst} = -\Omega^{-2} dt^2 + r_0^2 g_{\mathbb{S}^3}$.

- Colloquially speaking **incoherent radiation passes happily through the wound**. However, cosmological constants can change from any positive value to any other one. Ha...

- Considering the case $\check{w} = 1/3$, one shows that **remarkably** $\hat{w} = 1/3$ (generalization of the result of **Paul Tod**). More explicitly this case can be integrated to the very end.

- Theorem.** The function $\Omega = \Omega(t)$ given by:

$$\Omega^2 = \frac{3 - 3 \cosh(2\sqrt{\frac{\check{\lambda}}{3}}t) - 2r_0^2 \sqrt{\check{\lambda}\hat{\lambda}} \sinh(2\sqrt{\frac{\check{\lambda}}{3}}t)}{\check{\lambda}r_0^2}$$

has the property that both $\check{g} = \Omega^2 g_{Einst}$ and $\hat{g} = \Omega^{-2} g_{Einst}$ satisfy Einstein's equations with polytropic perfect fluid equation of state, for which $\hat{w} = \check{w} = 1/3$ (radiation), and with the corresponding cosmological constants $\check{\lambda}$ and $\hat{\lambda}$. Here $g_{Einst} = -\Omega^{-2} dt^2 + r_0^2 g_{\mathbb{S}^3}$.

- Colloquially speaking **incoherent radiation passes happily through the wound**. However, cosmological constants can change from any positive value to any other one. Ha...

- Considering the case $\check{w} = 1/3$, one shows that **remarkably** $\hat{w} = 1/3$ (generalization of the result of **Paul Tod**). More explicitly this case can be integrated to the very end.

- Theorem.** The function $\Omega = \Omega(t)$ given by:

$$\Omega^2 = \frac{3 - 3 \cosh(2\sqrt{\frac{\check{\lambda}}{3}}t) - 2r_0^2 \sqrt{\check{\lambda}\hat{\lambda}} \sinh(2\sqrt{\frac{\check{\lambda}}{3}}t)}{\check{\lambda}r_0^2}$$

has the property that both $\check{g} = \Omega^2 g_{Einst}$ and $\hat{g} = \Omega^{-2} g_{Einst}$ satisfy Einstein's equations with polytropic perfect fluid equation of state, for which $\hat{w} = \check{w} = 1/3$ (radiation), and with the corresponding cosmological constants $\check{\lambda}$ and $\hat{\lambda}$. Here $g_{Einst} = -\Omega^{-2} dt^2 + r_0^2 g_{\mathbb{S}^3}$.

- Colloquially speaking **incoherent radiation passes happily through the wound**. However, cosmological constants can change from any positive value to any other one. Ha...

- Considering the case $\check{w} = 1/3$, one shows that **remarkably** $\hat{w} = 1/3$ (generalization of the result of **Paul Tod**). More explicitly this case can be integrated to the very end.

- Theorem.** The function $\Omega = \Omega(t)$ given by:

$$\Omega^2 = \frac{3 - 3 \cosh(2\sqrt{\frac{\check{\lambda}}{3}}t) - 2r_0^2 \sqrt{\check{\lambda}\hat{\lambda}} \sinh(2\sqrt{\frac{\check{\lambda}}{3}}t)}{\check{\lambda}r_0^2}$$

has the property that both $\check{g} = \Omega^2 g_{Einst}$ and $\hat{g} = \Omega^{-2} g_{Einst}$ satisfy Einstein's equations with polytropic perfect fluid equation of state, for which $\hat{w} = \check{w} = 1/3$ (radiation), and with the corresponding cosmological constants $\check{\lambda}$ and $\hat{\lambda}$. Here $g_{Einst} = -\Omega^{-2} dt^2 + r_0^2 g_{\mathbb{S}^3}$.

- Colloquially speaking **incoherent radiation passes happily through the wound**. However, cosmological constants can change from any positive value to any other one. Ha...

- Considering the case $\check{w} = 1/3$, one shows that **remarkably** $\hat{w} = 1/3$ (generalization of the result of **Paul Tod**). More explicitly this case can be integrated to the very end.
- Theorem.** The function $\Omega = \Omega(t)$ given by:

$$\Omega^2 = \frac{3 - 3 \cosh(2\sqrt{\frac{\check{\lambda}}{3}}t) - 2r_0^2 \sqrt{\check{\lambda}\hat{\lambda}} \sinh(2\sqrt{\frac{\check{\lambda}}{3}}t)}{\check{\lambda}r_0^2}$$

has the property that both $\check{g} = \Omega^2 g_{Einst}$ and $\hat{g} = \Omega^{-2} g_{Einst}$ satisfy Einstein's equations with polytropic perfect fluid equation of state, for which $\hat{w} = \check{w} = 1/3$ (radiation), and with the corresponding cosmological constants $\check{\lambda}$ and $\hat{\lambda}$. Here $g_{Einst} = -\Omega^{-2} dt^2 + r_0^2 g_{S^3}$.

- Colloquially speaking **incoherent radiation passes happily through the wound**. However, cosmological constants can change from any positive value to any other one. Ha...

- Considering the case $\check{w} = 1/3$, one shows that **remarkably** $\hat{w} = 1/3$ (generalization of the result of **Paul Tod**). More explicitly this case can be integrated to the very end.

- Theorem.** The function $\Omega = \Omega(t)$ given by:

$$\Omega^2 = \frac{3 - 3 \cosh(2\sqrt{\frac{\check{\lambda}}{3}}t) - 2r_0^2 \sqrt{\check{\lambda}\hat{\lambda}} \sinh(2\sqrt{\frac{\check{\lambda}}{3}}t)}{\check{\lambda}r_0^2}$$

has the property that both $\check{g} = \Omega^2 g_{Einst}$ and $\hat{g} = \Omega^{-2} g_{Einst}$ satisfy Einstein's equations with polytropic perfect fluid equation of state, for which $\hat{w} = \check{w} = 1/3$ (radiation), and with the corresponding cosmological constants $\check{\lambda}$ and $\hat{\lambda}$. Here $g_{Einst} = -\Omega^{-2} dt^2 + r_0^2 g_{\mathbb{S}^3}$.

- Colloquially speaking **incoherent radiation passes happily through the wound**. However, cosmological constants can change from any positive value to any other one. Ha...

- Considering the case $\check{w} = 1/3$, one shows that **remarkably** $\hat{w} = 1/3$ (generalization of the result of **Paul Tod**). More explicitly this case can be integrated to the very end.

- Theorem.** The function $\Omega = \Omega(t)$ given by:

$$\Omega^2 = \frac{3 - 3 \cosh(2\sqrt{\frac{\check{\lambda}}{3}}t) - 2r_0^2 \sqrt{\check{\lambda}\hat{\lambda}} \sinh(2\sqrt{\frac{\check{\lambda}}{3}}t)}{\check{\lambda}r_0^2}$$

has the property that both $\check{g} = \Omega^2 g_{Einst}$ and $\hat{g} = \Omega^{-2} g_{Einst}$ satisfy Einstein's equations with polytropic perfect fluid equation of state, for which $\hat{w} = \check{w} = 1/3$ (radiation), and with the corresponding cosmological constants $\check{\lambda}$ and $\hat{\lambda}$. Here $g_{Einst} = -\Omega^{-2} dt^2 + r_0^2 g_{\mathbb{S}^3}$.

- Colloquially speaking **incoherent radiation passes happily through the wound**. However, cosmological constants can change from any positive value to any other one. Ha...

- Considering the case $\check{w} = 1/3$, one shows that **remarkably** $\hat{w} = 1/3$ (generalization of the result of **Paul Tod**). More explicitly this case can be integrated to the very end.

- Theorem.** The function $\Omega = \Omega(t)$ given by:

$$\Omega^2 = \frac{3 - 3 \cosh(2\sqrt{\frac{\check{\lambda}}{3}}t) - 2r_0^2 \sqrt{\check{\lambda}\hat{\lambda}} \sinh(2\sqrt{\frac{\check{\lambda}}{3}}t)}{\check{\lambda}r_0^2}$$

has the property that both $\check{g} = \Omega^2 g_{Einst}$ and $\hat{g} = \Omega^{-2} g_{Einst}$ satisfy Einstein's equations with polytropic perfect fluid equation of state, for which $\hat{w} = \check{w} = 1/3$ (radiation), and with the corresponding cosmological constants $\check{\lambda}$ and $\hat{\lambda}$. Here $g_{Einst} = -\Omega^{-2} dt^2 + r_0^2 g_{\mathbb{S}^3}$.

- Colloquially speaking **incoherent radiation passes happily through the wound**. However, cosmological constants can change from any positive value to any other one. Ha...

- Considering the case $\check{w} = 1/3$, one shows that **remarkably** $\hat{w} = 1/3$ (generalization of the result of **Paul Tod**). More explicitly this case can be integrated to the very end.

- Theorem.** The function $\Omega = \Omega(t)$ given by:

$$\Omega^2 = \frac{3 - 3 \cosh(2\sqrt{\frac{\check{\lambda}}{3}}t) - 2r_0^2 \sqrt{\check{\lambda}\hat{\lambda}} \sinh(2\sqrt{\frac{\check{\lambda}}{3}}t)}{\check{\lambda}r_0^2}$$

has the property that both $\check{g} = \Omega^2 g_{Einst}$ and $\hat{g} = \Omega^{-2} g_{Einst}$ satisfy Einstein's equations with polytropic perfect fluid equation of state, for which $\hat{w} = \check{w} = 1/3$ (radiation), and with the corresponding cosmological constants $\check{\lambda}$ and $\hat{\lambda}$. Here $g_{Einst} = -\Omega^{-2} dt^2 + r_0^2 g_{S^3}$.

- Colloquially speaking **incoherent radiation passes happily through the wound**. However, cosmological constants can change from any positive value to any other one. Ha...

- Considering the case $\check{w} = 1/3$, one shows that **remarkably** $\hat{w} = 1/3$ (generalization of the result of **Paul Tod**). More explicitly this case can be integrated to the very end.

- Theorem.** The function $\Omega = \Omega(t)$ given by:

$$\Omega^2 = \frac{3 - 3 \cosh(2\sqrt{\frac{\check{\lambda}}{3}}t) - 2r_0^2 \sqrt{\check{\lambda}\hat{\lambda}} \sinh(2\sqrt{\frac{\check{\lambda}}{3}}t)}{\check{\lambda}r_0^2}$$

has the property that both $\check{g} = \Omega^2 g_{Einst}$ and $\hat{g} = \Omega^{-2} g_{Einst}$ satisfy Einstein's equations with polytropic perfect fluid equation of state, for which $\hat{w} = \check{w} = 1/3$ (radiation), and with the corresponding cosmological constants $\check{\lambda}$ and $\hat{\lambda}$. Here $g_{Einst} = -\Omega^{-2} dt^2 + r_0^2 g_{S^3}$.

- Colloquially speaking **incoherent radiation passes happily through the wound**. However, cosmological constants can change from any positive value to any other one. Ha...

- Considering the case $\check{w} = 1/3$, one shows that **remarkably** $\hat{w} = 1/3$ (generalization of the result of **Paul Tod**). More explicitly this case can be integrated to the very end.

- Theorem.** The function $\Omega = \Omega(t)$ given by:

$$\Omega^2 = \frac{3 - 3 \cosh(2\sqrt{\frac{\check{\lambda}}{3}}t) - 2r_0^2 \sqrt{\check{\lambda}\hat{\lambda}} \sinh(2\sqrt{\frac{\check{\lambda}}{3}}t)}{\check{\lambda}r_0^2}$$

has the property that both $\check{g} = \Omega^2 g_{Einst}$ and $\hat{g} = \Omega^{-2} g_{Einst}$ satisfy Einstein's equations with polytropic perfect fluid equation of state, for which $\hat{w} = \check{w} = 1/3$ (radiation), and with the corresponding cosmological constants $\check{\lambda}$ and $\hat{\lambda}$.

Here $g_{Einst} = -\Omega^{-2} dt^2 + r_0^2 g_{S^3}$.

- Colloquially speaking **incoherent radiation passes happily through the wound**. However, cosmological constants can change from any positive value to any other one. Ha...

- Considering the case $\check{w} = 1/3$, one shows that **remarkably** $\hat{w} = 1/3$ (generalization of the result of **Paul Tod**). More explicitly this case can be integrated to the very end.

- Theorem.** The function $\Omega = \Omega(t)$ given by:

$$\Omega^2 = \frac{3 - 3 \cosh(2\sqrt{\frac{\check{\Lambda}}{3}}t) - 2r_0^2 \sqrt{\check{\Lambda}\hat{\Lambda}} \sinh(2\sqrt{\frac{\check{\Lambda}}{3}}t)}{\check{\Lambda}r_0^2}$$

has the property that both $\check{g} = \Omega^2 g_{Einst}$ and $\hat{g} = \Omega^{-2} g_{Einst}$ satisfy Einstein's equations with polytropic perfect fluid equation of state, for which $\hat{w} = \check{w} = 1/3$ (radiation), and with the corresponding cosmological constants $\check{\Lambda}$ and $\hat{\Lambda}$.

Here $g_{Einst} = -\Omega^{-2} dt^2 + r_0^2 g_{S^3}$.

- Colloquially speaking **incoherent radiation passes happily through the wound**. However, cosmological constants can change from any positive value to any other one. Ha...

- Considering the case $\check{w} = 1/3$, one shows that **remarkably** $\hat{w} = 1/3$ (generalization of the result of **Paul Tod**). More explicitly this case can be integrated to the very end.

- Theorem.** The function $\Omega = \Omega(t)$ given by:

$$\Omega^2 = \frac{3 - 3 \cosh(2\sqrt{\frac{\check{\Lambda}}{3}}t) - 2r_0^2 \sqrt{\check{\Lambda}\hat{\Lambda}} \sinh(2\sqrt{\frac{\check{\Lambda}}{3}}t)}{\check{\Lambda}r_0^2}$$

has the property that both $\check{g} = \Omega^2 g_{Einst}$ and $\hat{g} = \Omega^{-2} g_{Einst}$ satisfy Einstein's equations with polytropic perfect fluid equation of state, for which $\hat{w} = \check{w} = 1/3$ (radiation), and with the corresponding cosmological constants $\check{\Lambda}$ and $\hat{\Lambda}$. Here $g_{Einst} = -\Omega^{-2} dt^2 + r_0^2 g_{\mathbb{S}^3}$.

- Colloquially speaking **incoherent radiation passes happily through the wound**. However, cosmological constants can change from any positive value to any other one. Ha...

- Considering the case $\check{w} = 1/3$, one shows that **remarkably** $\hat{w} = 1/3$ (generalization of the result of **Paul Tod**). More explicitly this case can be integrated to the very end.

- Theorem.** The function $\Omega = \Omega(t)$ given by:

$$\Omega^2 = \frac{3 - 3 \cosh(2\sqrt{\frac{\check{\Lambda}}{3}}t) - 2r_0^2 \sqrt{\check{\Lambda}\hat{\Lambda}} \sinh(2\sqrt{\frac{\check{\Lambda}}{3}}t)}{\check{\Lambda}r_0^2}$$

has the property that both $\check{g} = \Omega^2 g_{Einst}$ and $\hat{g} = \Omega^{-2} g_{Einst}$ satisfy Einstein's equations with polytropic perfect fluid equation of state, for which $\hat{w} = \check{w} = 1/3$ (radiation), and with the corresponding cosmological constants $\check{\Lambda}$ and $\hat{\Lambda}$. Here $g_{Einst} = -\Omega^{-2} dt^2 + r_0^2 g_{\mathbb{S}^3}$.

- Colloquially speaking **incoherent radiation passes happily through the wound**. However, cosmological constants can change from any positive value to any other one. Ha...

- Considering the case $\check{w} = 1/3$, one shows that **remarkably** $\hat{w} = 1/3$ (generalization of the result of **Paul Tod**). More explicitly this case can be integrated to the very end.

- Theorem.** The function $\Omega = \Omega(t)$ given by:

$$\Omega^2 = \frac{3 - 3 \cosh(2\sqrt{\frac{\check{\Lambda}}{3}}t) - 2r_0^2 \sqrt{\check{\Lambda}\hat{\Lambda}} \sinh(2\sqrt{\frac{\check{\Lambda}}{3}}t)}{\check{\Lambda}r_0^2}$$

has the property that both $\check{g} = \Omega^2 g_{Einst}$ and $\hat{g} = \Omega^{-2} g_{Einst}$ satisfy Einstein's equations with polytropic perfect fluid equation of state, for which $\hat{w} = \check{w} = 1/3$ (radiation), and with the corresponding cosmological constants $\check{\Lambda}$ and $\hat{\Lambda}$. Here $g_{Einst} = -\Omega^{-2} dt^2 + r_0^2 g_{\mathbb{S}^3}$.

- Colloquially speaking **incoherent radiation passes happily through the wound**. However, cosmological constants can change from any positive value to any other one. Ha...

- Considering the case $\check{w} = 1/3$, one shows that **remarkably** $\hat{w} = 1/3$ (generalization of the result of **Paul Tod**). More explicitly this case can be integrated to the very end.

- Theorem.** The function $\Omega = \Omega(t)$ given by:

$$\Omega^2 = \frac{3 - 3 \cosh(2\sqrt{\frac{\check{\Lambda}}{3}}t) - 2r_0^2 \sqrt{\check{\Lambda}\hat{\Lambda}} \sinh(2\sqrt{\frac{\check{\Lambda}}{3}}t)}{\check{\Lambda}r_0^2}$$

has the property that both $\check{g} = \Omega^2 g_{Einst}$ and $\hat{g} = \Omega^{-2} g_{Einst}$ satisfy Einstein's equations with polytropic perfect fluid equation of state, for which $\hat{w} = \check{w} = 1/3$ (radiation), and with the corresponding cosmological constants $\check{\Lambda}$ and $\hat{\Lambda}$. Here $g_{Einst} = -\Omega^{-2} dt^2 + r_0^2 g_{\mathbb{S}^3}$.

- Colloquially speaking **incoherent radiation passes happily through the wound**. However, cosmological constants can change from any positive value to any other one. Ha...

- Consider (special) **Bianchi models**: $\check{g} = -dt^2 + \Omega^2(t)h_{ij}\sigma^i\sigma^j$, where σ^i , $i, j = 1, 2, 3$, are left invariant forms on a 3-dimensional Lie group G , and (h_{ij}) is a symmetric positive definite matrix. Here the Universe manifold is $M = \mathbb{R} \times G$, and the time variable t is along the \mathbb{R} factor. For each **Bianchi type** of G , decide which metric should play the role of g_{Einst} . In other words: find a preferred basis σ^i of the left invariant forms on G , so that the counterpart of g_{Einst} is $g_E = -\Omega^{-2}(t)dt^2 + (\sigma^1)^2 + (\sigma^2)^2 + (\sigma^3)^2$; and then play the game with $\check{g} = \Omega^2 g_E$ and $\hat{g} = \Omega^{-2} g_E$, similar to this I was describing in this talk. My game was **Bianchi IX**, i.e. I took $G = SU(2) = \mathbb{S}^3$.
- More generally, take as g_E the metric $g_E = -\Omega^{-2}(t)dt^2 + g_S$, where $M = \mathbb{R} \times S$, and (S, g_S) is a 3D Riemannian manifold (possibly **satisfying some interesting equations**, or as in the previous case **symmetry conditions**). Again play the game with polytropic perfect fluids for $\check{g} = \Omega^2 g_E$ and $\hat{g} = \Omega^{-2} g_E$

- Consider (special) **Bianchi models**: $\check{g} = -dt^2 + \Omega^2(t)h_{ij}\sigma^i\sigma^j$, where σ^i , $i, j = 1, 2, 3$, are left invariant forms on a 3-dimensional Lie group G , and (h_{ij}) is a symmetric positive definite matrix. Here the Universe manifold is $M = \mathbb{R} \times G$, and the time variable t is along the \mathbb{R} factor. For each **Bianchi type** of G , decide which metric should play the role of g_{Einst} . In other words: find a preferred basis σ^i of the left invariant forms on G , so that the counterpart of g_{Einst} is $g_E = -\Omega^{-2}(t)dt^2 + (\sigma^1)^2 + (\sigma^2)^2 + (\sigma^3)^2$; and then play the game with $\check{g} = \Omega^2 g_E$ and $\hat{g} = \Omega^{-2} g_E$, similar to this I was describing in this talk. My game was **Bianchi IX**, i.e. I took $G = SU(2) = \mathbb{S}^3$.
- More generally, take as g_E the metric $g_E = -\Omega^{-2}(t)dt^2 + g_S$, where $M = \mathbb{R} \times S$, and (S, g_S) is a 3D Riemannian manifold (possibly **satisfying some interesting equations**, or as in the previous case **symmetry conditions**). Again play the game with polytropic perfect fluids for $\check{g} = \Omega^2 g_E$ and $\hat{g} = \Omega^{-2} g_E$

- Consider (special) **Bianchi models**: $\check{g} = -dt^2 + \Omega^2(t)h_{ij}\sigma^i\sigma^j$, where σ^i , $i, j = 1, 2, 3$, are left invariant forms on a 3-dimensional Lie group G , and (h_{ij}) is a symmetric positive definite matrix. Here the Universe manifold is $M = \mathbb{R} \times G$, and the time variable t is along the \mathbb{R} factor. For each **Bianchi type** of G , decide which metric should play the role of g_{Einst} . In other words: find a preferred basis σ^i of the left invariant forms on G , so that the counterpart of g_{Einst} is $g_E = -\Omega^{-2}(t)dt^2 + (\sigma^1)^2 + (\sigma^2)^2 + (\sigma^3)^2$; and then play the game with $\check{g} = \Omega^2 g_E$ and $\hat{g} = \Omega^{-2} g_E$, similar to this I was describing in this talk. My game was **Bianchi IX**, i.e. I took $G = SU(2) = \mathbb{S}^3$.
- More generally, take as g_E the metric $g_E = -\Omega^{-2}(t)dt^2 + g_S$, where $M = \mathbb{R} \times S$, and (S, g_S) is a 3D Riemannian manifold (possibly **satisfying some interesting equations**, or as in the previous case **symmetry conditions**). Again play the game with polytropic perfect fluids for $\check{g} = \Omega^2 g_E$ and $\hat{g} = \Omega^{-2} g_E$

- Consider (special) **Bianchi models**: $\check{g} = -dt^2 + \Omega^2(t)h_{ij}\sigma^i\sigma^j$, where σ^i , $i, j = 1, 2, 3$, are left invariant forms on a 3-dimensional Lie group G , and (h_{ij}) is a symmetric positive definite matrix. Here the Universe manifold is $M = \mathbb{R} \times G$, and the time variable t is along the \mathbb{R} factor. For each **Bianchi type** of G , decide which metric should play the role of g_{Einst} . In other words: find a preferred basis σ^i of the left invariant forms on G , so that the counterpart of g_{Einst} is $g_E = -\Omega^{-2}(t)dt^2 + (\sigma^1)^2 + (\sigma^2)^2 + (\sigma^3)^2$; and then play the game with $\check{g} = \Omega^2 g_E$ and $\hat{g} = \Omega^{-2} g_E$, similar to this I was describing in this talk. My game was **Bianchi IX**, i.e. I took $G = SU(2) = \mathbb{S}^3$.
- More generally, take as g_E the metric $g_E = -\Omega^{-2}(t)dt^2 + g_S$, where $M = \mathbb{R} \times S$, and (S, g_S) is a 3D Riemannian manifold (possibly **satisfying some interesting equations**, or as in the previous case **symmetry conditions**). Again play the game with polytropic perfect fluids for $\check{g} = \Omega^2 g_E$ and $\hat{g} = \Omega^{-2} g_E$

- Consider (special) **Bianchi models**: $\check{g} = -dt^2 + \Omega^2(t)h_{ij}\sigma^i\sigma^j$, where σ^i , $i, j = 1, 2, 3$, are left invariant forms on a 3-dimensional Lie group G , and (h_{ij}) is a symmetric positive definite matrix. Here the Universe manifold is $M = \mathbb{R} \times G$, and the time variable t is along the \mathbb{R} factor. For each **Bianchi type** of G , decide which metric should play the role of g_{Einst} . In other words: find a preferred basis σ^i of the left invariant forms on G , so that the counterpart of g_{Einst} is $g_E = -\Omega^{-2}(t)dt^2 + (\sigma^1)^2 + (\sigma^2)^2 + (\sigma^3)^2$; and then play the game with $\check{g} = \Omega^2 g_E$ and $\hat{g} = \Omega^{-2} g_E$, similar to this I was describing in this talk. My game was **Bianchi IX**, i.e. I took $G = SU(2) = \mathbb{S}^3$.
- More generally, take as g_E the metric $g_E = -\Omega^{-2}(t)dt^2 + g_S$, where $M = \mathbb{R} \times S$, and (S, g_S) is a 3D Riemannian manifold (possibly **satisfying some interesting equations**, or as in the previous case **symmetry conditions**). Again play the game with polytropic perfect fluids for $\check{g} = \Omega^2 g_E$ and $\hat{g} = \Omega^{-2} g_E$

- Consider (special) **Bianchi models**: $\check{g} = -dt^2 + \Omega^2(t)h_{ij}\sigma^i\sigma^j$, where σ^i , $i, j = 1, 2, 3$, are left invariant forms on a 3-dimensional Lie group G , and (h_{ij}) is a symmetric positive definite matrix. Here the Universe manifold is $M = \mathbb{R} \times G$, and the time variable t is along the \mathbb{R} factor. For each **Bianchi type** of G , decide which metric should play the role of g_{Einst} . In other words: find a preferred basis σ^i of the left invariant forms on G , so that the counterpart of g_{Einst} is $g_E = -\Omega^{-2}(t)dt^2 + (\sigma^1)^2 + (\sigma^2)^2 + (\sigma^3)^2$; and then play the game with $\check{g} = \Omega^2 g_E$ and $\hat{g} = \Omega^{-2} g_E$, similar to this I was describing in this talk. My game was **Bianchi IX**, i.e. I took $G = SU(2) = \mathbb{S}^3$.
- More generally, take as g_E the metric $g_E = -\Omega^{-2}(t)dt^2 + g_S$, where $M = \mathbb{R} \times S$, and (S, g_S) is a 3D Riemannian manifold (possibly **satisfying some interesting equations**, or as in the previous case **symmetry conditions**). Again play the game with polytropic perfect fluids for $\check{g} = \Omega^2 g_E$ and $\hat{g} = \Omega^{-2} g_E$

- Consider (special) **Bianchi models**: $\check{g} = -dt^2 + \Omega^2(t)h_{ij}\sigma^i\sigma^j$, where σ^i , $i, j = 1, 2, 3$, are left invariant forms on a 3-dimensional Lie group G , and (h_{ij}) is a symmetric positive definite matrix. Here the Universe manifold is $M = \mathbb{R} \times G$, and the time variable t is along the \mathbb{R} factor. For each **Bianchi type** of G , decide which metric should play the role of g_{Einst} . In other words: find a preferred basis σ^i of the left invariant forms on G , so that the counterpart of g_{Einst} is $g_E = -\Omega^{-2}(t)dt^2 + (\sigma^1)^2 + (\sigma^2)^2 + (\sigma^3)^2$; and then play the game with $\check{g} = \Omega^2 g_E$ and $\hat{g} = \Omega^{-2} g_E$, similar to this I was describing in this talk. My game was **Bianchi IX**, i.e. I took $G = SU(2) = \mathbb{S}^3$.
- More generally, take as g_E the metric $g_E = -\Omega^{-2}(t)dt^2 + g_S$, where $M = R \times S$, and (S, g_S) is a 3D Riemannian manifold (possibly **satisfying some interesting equations**, or as in the previous case **symmetry conditions**). Again play the game with polytropic perfect fluids for $\check{g} = \Omega^2 g_E$ and $\hat{g} = \Omega^{-2} g_E$

- Consider (special) **Bianchi models**: $\check{g} = -dt^2 + \Omega^2(t)h_{ij}\sigma^i\sigma^j$, where σ^i , $i, j = 1, 2, 3$, are left invariant forms on a 3-dimensional Lie group G , and (h_{ij}) is a symmetric positive definite matrix. Here the Universe manifold is $M = \mathbb{R} \times G$, and the time variable t is along the \mathbb{R} factor. For each **Bianchi type** of G , decide which metric should play the role of g_{Einst} . In other words: find a preferred basis σ^i of the left invariant forms on G , so that the counterpart of g_{Einst} is $g_E = -\Omega^{-2}(t)dt^2 + (\sigma^1)^2 + (\sigma^2)^2 + (\sigma^3)^2$; and then play the game with $\check{g} = \Omega^2 g_E$ and $\hat{g} = \Omega^{-2} g_E$, similar to this I was describing in this talk. My game was **Bianchi IX**, i.e. I took $G = SU(2) = \mathbb{S}^3$.
- More generally, take as g_E the metric $g_E = -\Omega^{-2}(t)dt^2 + g_S$, where $M = R \times S$, and (S, g_S) is a 3D Riemannian manifold (possibly **satisfying some interesting equations**, or as in the previous case **symmetry conditions**). Again play the game with polytropic perfect fluids for $\check{g} = \Omega^2 g_E$ and $\hat{g} = \Omega^{-2} g_E$

- Consider (special) **Bianchi models**: $\check{g} = -dt^2 + \Omega^2(t)h_{ij}\sigma^i\sigma^j$, where σ^i , $i, j = 1, 2, 3$, are left invariant forms on a 3-dimensional Lie group G , and (h_{ij}) is a symmetric positive definite matrix. Here the Universe manifold is $M = \mathbb{R} \times G$, and the time variable t is along the \mathbb{R} factor. For each **Bianchi type** of G , decide which metric should play the role of g_{Einst} . In other words: find a preferred basis σ^i of the left invariant forms on G , so that the counterpart of g_{Einst} is $g_E = -\Omega^{-2}(t)dt^2 + (\sigma^1)^2 + (\sigma^2)^2 + (\sigma^3)^2$; and then play the game with $\check{g} = \Omega^2 g_E$ and $\hat{g} = \Omega^{-2} g_E$, similar to this I was describing in this talk. My game was **Bianchi IX**, i.e. I took $G = SU(2) = \mathbb{S}^3$.
- More generally, take as g_E the metric $g_E = -\Omega^{-2}(t)dt^2 + g_S$, where $M = R \times S$, and (S, g_S) is a 3D Riemannian manifold (possibly **satisfying some interesting equations**, or as in the previous case **symmetry conditions**). Again play the game with polytropic perfect fluids for $\check{g} = \Omega^2 g_E$ and $\hat{g} = \Omega^{-2} g_E$

- Consider (special) **Bianchi models**: $\check{g} = -dt^2 + \Omega^2(t)h_{ij}\sigma^i\sigma^j$, where σ^i , $i, j = 1, 2, 3$, are left invariant forms on a 3-dimensional Lie group G , and (h_{ij}) is a symmetric positive definite matrix. Here the Universe manifold is $M = \mathbb{R} \times G$, and the time variable t is along the \mathbb{R} factor. For each **Bianchi type** of G , decide which metric should play the role of g_{Einst} . In other words: find a preferred basis σ^i of the left invariant forms on G , so that the counterpart of g_{Einst} is $g_E = -\Omega^{-2}(t)dt^2 + (\sigma^1)^2 + (\sigma^2)^2 + (\sigma^3)^2$; and then play the game with $\check{g} = \Omega^2 g_E$ and $\hat{g} = \Omega^{-2} g_E$, similar to this I was describing in this talk. My game was **Bianchi IX**, i.e. I took $G = SU(2) = \mathbb{S}^3$.
- More generally, take as g_E the metric $g_E = -\Omega^{-2}(t)dt^2 + g_S$, where $M = R \times S$, and (S, g_S) is a 3D Riemannian manifold (possibly **satisfying some interesting equations**, or as in the previous case **symmetry conditions**). Again play the game with polytropic perfect fluids for $\check{g} = \Omega^2 g_E$ and $\hat{g} = \Omega^{-2} g_E$

- Consider (special) **Bianchi models**: $\check{g} = -dt^2 + \Omega^2(t)h_{ij}\sigma^i\sigma^j$, where σ^i , $i, j = 1, 2, 3$, are left invariant forms on a 3-dimensional Lie group G , and (h_{ij}) is a symmetric positive definite matrix. Here the Universe manifold is $M = \mathbb{R} \times G$, and the time variable t is along the \mathbb{R} factor. For each **Bianchi type** of G , decide which metric should play the role of g_{Einst} . In other words: find a preferred basis σ^i of the left invariant forms on G , so that the counterpart of g_{Einst} is $g_E = -\Omega^{-2}(t)dt^2 + (\sigma^1)^2 + (\sigma^2)^2 + (\sigma^3)^2$; and then play the game with $\check{g} = \Omega^2 g_E$ and $\hat{g} = \Omega^{-2} g_E$, similar to this I was describing in this talk. My game was **Bianchi IX**, i.e. I took $G = SU(2) = \mathbb{S}^3$.
- More generally, take as g_E the metric $g_E = -\Omega^{-2}(t)dt^2 + g_S$, where $M = R \times S$, and (S, g_S) is a 3D Riemannian manifold (possibly **satisfying some interesting equations**, or as in the previous case **symmetry conditions**). Again play the game with polytropic perfect fluids for $\check{g} = \Omega^2 g_E$ and $\hat{g} = \Omega^{-2} g_E$

- Consider (special) **Bianchi models**: $\check{g} = -dt^2 + \Omega^2(t)h_{ij}\sigma^i\sigma^j$, where σ^i , $i, j = 1, 2, 3$, are left invariant forms on a 3-dimensional Lie group G , and (h_{ij}) is a symmetric positive definite matrix. Here the Universe manifold is $M = \mathbb{R} \times G$, and the time variable t is along the \mathbb{R} factor. For each **Bianchi type** of G , decide which metric should play the role of g_{Einst} . In other words: find a preferred basis σ^i of the left invariant forms on G , so that the counterpart of g_{Einst} is $g_E = -\Omega^{-2}(t)dt^2 + (\sigma^1)^2 + (\sigma^2)^2 + (\sigma^3)^2$; and then play the game with $\check{g} = \Omega^2 g_E$ and $\hat{g} = \Omega^{-2} g_E$, similar to this I was describing in this talk. My game was **Bianchi IX**, i.e. I took $G = SU(2) = \mathbb{S}^3$.
- More generally, take as g_E the metric $g_E = -\Omega^{-2}(t)dt^2 + g_S$, where $M = R \times S$, and (S, g_S) is a 3D Riemannian manifold (possibly **satisfying some interesting equations**, or as in the previous case **symmetry conditions**). Again play the game with polytropic perfect fluids for $\check{g} = \Omega^2 g_E$ and $\hat{g} = \Omega^{-2} g_E$

- Consider (special) **Bianchi models**: $\check{g} = -dt^2 + \Omega^2(t)h_{ij}\sigma^i\sigma^j$, where σ^i , $i, j = 1, 2, 3$, are left invariant forms on a 3-dimensional Lie group G , and (h_{ij}) is a symmetric positive definite matrix. Here the Universe manifold is $M = \mathbb{R} \times G$, and the time variable t is along the \mathbb{R} factor. For each **Bianchi type** of G , decide which metric should play the role of g_{Einst} . In other words: find a preferred basis σ^i of the left invariant forms on G , so that the counterpart of g_{Einst} is $g_E = -\Omega^{-2}(t)dt^2 + (\sigma^1)^2 + (\sigma^2)^2 + (\sigma^3)^2$; and then play the game with $\check{g} = \Omega^2 g_E$ and $\hat{g} = \Omega^{-2} g_E$, similar to this I was describing in this talk. My game was **Bianchi IX**, i.e. I took $G = SU(2) = \mathbb{S}^3$.
- More generally, take as g_E the metric $g_E = -\Omega^{-2}(t)dt^2 + g_S$, where $M = R \times S$, and (S, g_S) is a 3D Riemannian manifold (possibly **satisfying some interesting equations**, or as in the previous case **symmetry conditions**). Again play the game with polytropic perfect fluids for $\check{g} = \Omega^2 g_E$ and $\hat{g} = \Omega^{-2} g_E$

- Consider (special) **Bianchi models**: $\check{g} = -dt^2 + \Omega^2(t)h_{ij}\sigma^i\sigma^j$, where σ^i , $i, j = 1, 2, 3$, are left invariant forms on a 3-dimensional Lie group G , and (h_{ij}) is a symmetric positive definite matrix. Here the Universe manifold is $M = \mathbb{R} \times G$, and the time variable t is along the \mathbb{R} factor. For each **Bianchi type** of G , decide which metric should play the role of g_{Einst} . In other words: find a preferred basis σ^i of the left invariant forms on G , so that the counterpart of g_{Einst} is $g_E = -\Omega^{-2}(t)dt^2 + (\sigma^1)^2 + (\sigma^2)^2 + (\sigma^3)^2$; and then play the game with $\check{g} = \Omega^2 g_E$ and $\hat{g} = \Omega^{-2} g_E$, similar to this I was describing in this talk. My game was **Bianchi IX**, i.e. I took $G = SU(2) = \mathbb{S}^3$.
- More generally, take as g_E the metric $g_E = -\Omega^{-2}(t)dt^2 + g_S$, where $M = R \times S$, and (S, g_S) is a 3D Riemannian manifold (possibly **satisfying some interesting equations**, or as in the previous case **symmetry conditions**). Again play the game with polytropic perfect fluids for $\check{g} = \Omega^2 g_E$ and $\hat{g} = \Omega^{-2} g_E$

- Consider (special) **Bianchi models**: $\check{g} = -dt^2 + \Omega^2(t)h_{ij}\sigma^i\sigma^j$, where σ^i , $i, j = 1, 2, 3$, are left invariant forms on a 3-dimensional Lie group G , and (h_{ij}) is a symmetric positive definite matrix. Here the Universe manifold is $M = \mathbb{R} \times G$, and the time variable t is along the \mathbb{R} factor. For each **Bianchi type** of G , decide which metric should play the role of g_{Einst} . In other words: find a preferred basis σ^i of the left invariant forms on G , so that the counterpart of g_{Einst} is $g_E = -\Omega^{-2}(t)dt^2 + (\sigma^1)^2 + (\sigma^2)^2 + (\sigma^3)^2$; and then play the game with $\check{g} = \Omega^2 g_E$ and $\hat{g} = \Omega^{-2} g_E$, similar to this I was describing in this talk. My game was **Bianchi IX**, i.e. I took $G = SU(2) = \mathbb{S}^3$.
- More generally, take as g_E the metric $g_E = -\Omega^{-2}(t)dt^2 + g_S$, where $M = R \times S$, and (S, g_S) is a 3D Riemannian manifold (possibly **satisfying some interesting equations**, or as in the previous case **symmetry conditions**). Again play the game with polytropic perfect fluids for $\check{g} = \Omega^2 g_E$ and $\hat{g} = \Omega^{-2} g_E$

- Consider (special) **Bianchi models**: $\check{g} = -dt^2 + \Omega^2(t)h_{ij}\sigma^i\sigma^j$, where σ^i , $i, j = 1, 2, 3$, are left invariant forms on a 3-dimensional Lie group G , and (h_{ij}) is a symmetric positive definite matrix. Here the Universe manifold is $M = \mathbb{R} \times G$, and the time variable t is along the \mathbb{R} factor. For each **Bianchi type** of G , decide which metric should play the role of g_{Einst} . In other words: find a preferred basis σ^i of the left invariant forms on G , so that the counterpart of g_{Einst} is $g_E = -\Omega^{-2}(t)dt^2 + (\sigma^1)^2 + (\sigma^2)^2 + (\sigma^3)^2$; and then play the game with $\check{g} = \Omega^2 g_E$ and $\hat{g} = \Omega^{-2} g_E$, similar to this I was describing in this talk. My game was **Bianchi IX**, i.e. I took $G = SU(2) = \mathbb{S}^3$.
- More generally, take as g_E the metric $g_E = -\Omega^{-2}(t)dt^2 + g_S$, where $M = R \times S$, and (S, g_S) is a 3D Riemannian manifold (possibly **satisfying** some **interesting equations**, or as in the previous case **symmetry conditions**). Again play the game with polytropic perfect fluids for $\check{g} = \Omega^2 g_E$ and $\hat{g} = \Omega^{-2} g_E$

- **H. W. Brinkman** (1925), 'Einstein spaces which are mapped conformally on each other', *Math. Ann.* **94**, 119-145
- **P. Tod** (2015), 'The equations of Conformal Cyclic Cosmology', *Gen. Rel. Grav.* **47**, <https://doi.org/10.1007/s10714-015-1859-7>
- **P. Tod** (2018), 'Conformal methods in General Relativity with application to Conformal Cyclic Cosmology: A minicourse at IX International Meeting on Lorentz Geometry held in Warsaw' (ask Paul Tod for a copy)
- **K. Meissner, P. Nurowski** (2017), 'Conformal transformations and the beginning of the Universe', *Phys. Rev. D* **95**, Issue 8, 84016, 1-5.

- **H. W. Brinkman** (1925), 'Einstein spaces which are mapped conformally on each other', *Math. Ann.* **94**, 119-145
- **P. Tod** (2015), 'The equations of Conformal Cyclic Cosmology', *Gen. Rel. Grav.* **47**, <https://doi.org/10.1007/s10714-015-1859-7>
- **P. Tod** (2018), 'Conformal methods in General Relativity with application to Conformal Cyclic Cosmology: A minicourse at IX International Meeting on Lorentz Geometry held in Warsaw' (ask Paul Tod for a copy)
- **K. Meissner, P. Nurowski** (2017), 'Conformal transformations and the beginning of the Universe', *Phys. Rev. D* **95**, Issue 8, 84016, 1-5.

- **H. W. Brinkman** (1925), 'Einstein spaces which are mapped conformally on each other', *Math. Ann.* **94**, 119-145
- **P. Tod** (2015), 'The equations of Conformal Cyclic Cosmology', *Gen. Rel. Grav.* **47**, <https://doi.org/10.1007/s10714-015-1859-7>
- **P. Tod** (2018), 'Conformal methods in General Relativity with application to Conformal Cyclic Cosmology: A minicourse at IX International Meeting on Lorentz Geometry held in Warsaw' (ask Paul Tod for a copy)
- **K. Meissner, P. Nurowski** (2017), 'Conformal transformations and the beginning of the Universe', *Phys. Rev. D* **95**, Issue 8, 84016, 1-5.

- **H. W. Brinkman** (1925), 'Einstein spaces which are mapped conformally on each other', *Math. Ann.* **94**, 119-145
- **P. Tod** (2015), 'The equations of Conformal Cyclic Cosmology', *Gen. Rel. Grav.* **47**, <https://doi.org/10.1007/s10714-015-1859-7>
- **P. Tod** (2018), 'Conformal methods in General Relativity with application to Conformal Cyclic Cosmology: A minicourse at IX International Meeting on Lorentz Geometry held in Warsaw' (ask Paul Tod for a copy)
- **K. Meissner, P. Nurowski** (2017), 'Conformal transformations and the beginning of the Universe', *Phys. Rev. D* **95**, Issue 8, 84016, 1-5.

- **H. W. Brinkman** (1925), 'Einstein spaces which are mapped conformally on each other', *Math. Ann.* **94**, 119-145
- **P. Tod** (2015), 'The equations of Conformal Cyclic Cosmology', *Gen. Rel. Grav.* **47**, <https://doi.org/10.1007/s10714-015-1859-7>
- **P. Tod** (2018), 'Conformal methods in General Relativity with application to Conformal Cyclic Cosmology: A minicourse at IX International Meeting on Lorentz Geometry held in Warsaw' (ask Paul Tod for a copy)
- **K. Meissner, P. Nurowski** (2017), 'Conformal transformations and the beginning of the Universe', *Phys. Rev. D* **95**, Issue 8, 84016, 1-5.

- **H. W. Brinkman** (1925), 'Einstein spaces which are mapped conformally on each other', *Math. Ann.* **94**, 119-145
- **P. Tod** (2015), 'The equations of Conformal Cyclic Cosmology', *Gen. Rel. Grav.* **47**, <https://doi.org/10.1007/s10714-015-1859-7>
- **P. Tod** (2018), 'Conformal methods in General Relativity with application to Conformal Cyclic Cosmology: A minicourse at IX International Meeting on Lorentz Geometry held in Warsaw' (ask Paul Tod for a copy)
- **K. Meissner, P. Nurowski** (2017), 'Conformal transformations and the beginning of the Universe', *Phys. Rev. D* **95**, Issue 8, 84016, 1-5.

THANK YOU!